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A CATALOG OF NEIGHBORING GALAXIES

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ABSTRACT

We present an all-sky catalog of 451 nearby galaxies, each having an individual distance estimate $D \lesssim 10$ Mpc or the radial velocity $V_{\mathrm{LG}} < 550$ km s^{-1}. The catalog contains data on basic optical and H i properties of the galaxies, in particular, their diameters, absolute magnitudes, morphological types, circumnuclear region types, optical and H i surface brightnesses, rotational velocities, indicative mass-to-luminosity and H i mass-to-luminosity ratios, as well as a so-called tidal index, which quantifies the galaxy environment. We expect the catalog completeness to be roughly 70%–80% within 8 Mpc. About 85% of the Local Volume population are dwarf (dIr, dIm, and dSph) galaxies with $M_B > -17.0$, which contribute about 4% to the local luminosity density, and roughly 10%–15% to the local H i mass density. The H i mass-to-luminosity and the H i mass-to-total (indicative) mass ratios increase systematically from giant galaxies toward dwarfs, reaching maximum values about 5 in solar units for the most tiny objects. For the Local Volume disklike galaxies, their H i masses and angular momentum follow Zasov's linear relation, expected for rotating gaseous disks being near the threshold of gravitational instability, favorable for active star formation. We found that the mean local luminosity density exceeds 1.7–2.0 times the global density, in spite of the presence of the Tully void and the absence of rich clusters in the Local Volume. The mean local H i density is 1.4 times its "global" value derived from the H i
Q1 Parkes Sky Survey. However, the mean local baryon density $\Omega_b(<8\ \mathrm{Mpc}) = 2.3\%$ consists of only a half of the global baryon density, $\Omega_b = (3.9 \pm 0.8)\%$ (Turner 2002). The mean-square pairwise difference of radial velocities is about 100 km s^{-1} for spatial separations within 1 Mpc, increasing to $\sim$300 km s^{-1} on a scale of $\sim$3 Mpc. We calculated also the integral area of the sky occupied by the neighboring galaxies. Assuming the H i size of spiral and irregular galaxies to be 2.5 times their standard optical diameter and ignoring any evolution effect, we obtain the expected number of the line-of-sight intersections with the H i galaxy images to be $dn/dz \sim 0.4$, which does not contradict the observed number of absorptions in QSO spectra.

Key words: galaxies: distances and redshifts — galaxies: fundamental parameters — galaxies: luminosity function, mass function

1. INTRODUCTION

To solve numerous problems of extragalactic astronomy and cosmology, one needs to have a quite complete and representative collection of galaxies limited by a fixed distance. In distinction to the galaxy samples limited by the apparent magnitude (Sandage & Tammann 1981; Tully 1988), creation of the distance-limited sample is an exceptionally difficult task because of the huge difference of galaxies according to their luminosity, surface brightness, and other global parameters.

It seems plausible that we know at present all (or almost all) galaxies within 1 Mpc around the Milky Way. However, it will be recalled that among these 34 galaxies five were discovered only a few years ago: Sag dSph (Ibata, Gilmore, & Irwin 1995), And V (Armandroff, Davis, & Jacoby 1998), And VI (=PegdSph; Armandroff, Jacoby, & Davies 1999; Karachentsev & Karachentseva 1999), Cas dSph (Karachentsev & Karachentseva 1999), and Cetus (Whiting, Hau, & Irwin 1999). It has been suggested that about 80% of nearby galaxies still remain unrevealed because of their extremely low surface brightness (McGaugh 1996; Impey & Bothun 1997; O'Neil & Bothun 2000) or small dimension (Klypin et al. 1999) comparable with the size of a globular cluster. Choosing between these two extreme suggestions (we know about 80% or only 20% of galaxies in our immediate surroundings) allows one to verify modern theories of galaxy formation. Naturally, the process of discovery of new galaxies is infinite. However, over the last decade, tremendous advance has been made in searching for nearby ($D < 10$ Mpc) galaxies owing to purpose-oriented efforts of different observational teams. Since the time of publication of a list of 226 galaxies having radial velocities less than 500 km s^{-1} (Karachentsev 1994), the total number of nearby galaxies has already doubled. In addition, the employment of the *Hubble Space Telescope*, as well as new large ground-based telescopes, makes it possible to study the stellar population and measure accurate distances of galaxies that are situated outside the Local Group. Therefore, it is time for summarizing the wealth of new observational data on neighboring galaxies, allowing us to establish some basic statistical properties of the Local Volume (LV) sample.

The main purpose of our paper is to present a catalog of nearby galaxies with references to their distance and redshift determinations because both these quantities are used for selection of galaxies for the catalog. The recent observational papers and selection criteria are described in § 2. The catalog itself is given in § 3. The completeness and reliability of the sample are discussed in § 4. The spatial distribution of the neighboring galaxies is considered in § 5. The basic optical and H I properties of the Local sample are explored in §§ 6 and 7, and, finally, some cosmologic parameters extracted from the catalog are presented in § 8.

2. RECENT OBSERVATIONAL WORK AND THE LV SAMPLE CRITERION

The nearest rich cluster of galaxies in Virgo is situated at a distance of $D_{\rm Virgo} = 17$ Mpc from us (Tonry et al. 2001). Therefore, the LV of radius $D = 8$ Mpc, undisturbed by high virial motions, may be considered to be a representative volume with regard to its population, as well as to the presence of Q2 basic structural properties (groups, voids, etc.) within it. The first step toward compiling such a "fair" sample was made by Kraan-Korteweg & Tammann (1979), who published a catalog of nearby galaxies within 10 Mpc. The catalog contains 179 galaxies satisfying the condition $V_{\rm LG} < 500$ km s^{-1}, where $V_{\rm LG}$ is the galaxy radial velocity corrected for the observer movement with respect to the Local Group centroid.

It is apparent that many members of the Virgo Cluster with a systemic velocity ~1100 km s^{-1} and a velocity dispersion ~650 km s^{-1} satisfy the same condition too. For this reason, Kraan-Korteweg & Tammann excluded from their catalog all probable Virgo members within a radius of 6° from the cluster center. Global parameters in this sample were studied by Huchtmeier & Richter (1988). Later, Karachentsev (1994) published an updated version of the LV list, which contains 226 galaxies with $V_{\rm LG} < 500$ km s^{-1}. Following the same condition, Karachentsev, Makarov, & Huchtmeier (1999b) enlarged the initial LV sample to 303 objects basing on new optical and H I observations. A compilation of nearby galaxies was also produced by Schmidt & Boller (1992), however, it contains a lot of fictitious objects camouflaging as nearby galaxies.

Over the past few years, special searches for new nearby dwarf galaxies have been undertaken basing on the optical sky survey POSS-II/ESO/SERC, H I and infrared surveys of the zone of avoidance, "blind" sky surveys in the 21 cm line, HIPASS and HIJASS. Armandroff et al. (1998, 1999) carried out automated searches for dwarf companions of low surface brightness in a wide vicinity of the Andromeda galaxy (M 31). Karachentsev & Karachentseva with collaborators undertook a visual inspection of all POSS-II/ESO/SERC plates and found more than 500 nearby dwarf galaxy candidates, mainly of low surface brightness (Karachentseva & Karachentsev 1998 [KK]; Karachentseva, Karachentsev, & Richter 1999 [KKR]; Karachentseva & Karachentsev 2000 [KKs]; Karachentsev et al. 2000b [KKSG], and Karachentsev, Karachentseva, & Huchtmeier 2001a [KKH]). H I observations of these objects (Huchtmeier, Karachentsev, & Karachentseva 1997, 2001, 2003; Huchtmeier et al. 2000), as well as optical spectral observations (Makarov, Karachentsev, & Burenkov 2003) revealed about 100 new galaxies with radial velocities less than 500 km s^{-1}. A similar search for new Local Group members was carried out by Whiting, Hau, & Irwin (2002), however, with a nearly zero result. The nearest galaxy group around IC 342/Maffei behind the Milky Way was studied in the H I line (Huchtmeier et al. 1995; Kraan-Korteweg et al. 1994) and in IR (Buta & McCall 1999) that resulted in finding some new members of the group. In the region of another nearby group in Centaurus, new gas-rich dwarf galaxies were found by Kilborn et al. 2002 and Staveley-Smith et al. 1998. As a result of the joint efforts, the local number density of galaxies has been increased 2 times (!). More than 80% of new nearby dwarf galaxies were found on the POSS-II/ESO/SERC plates with subsequent survey of the candidates in the H I line. Will the high rate of discovery of new neighboring galaxies be preserved in the future or has the epoch of "great geographic discoveries" in the LV completed yet? The reply to this question is important for testing different scenarios of galaxy formation.

It is obvious that the condition $V_{\rm LG} < 500$ km s^{-1} for selection of nearby galaxies seems to be too simplified. Indeed, the kinematic distance of a galaxy in a group may be in error by several megaparsecs because of virial motions. Apart from chaotic (peculiar) motions, the kinematic distance is also affected by anisotropic expansion of the LV. According to Karachentsev & Makarov (1996, 2001), the local velocity field on a scale of ~8 Mpc is characterized by the Hubble tensor H_{ij}, which has the main values of $H_{xx}:H_{yy}:H_{zz} = (81 \pm 3):(62 \pm 3):(48 \pm 5)$ in km s^{-1} Mpc^{-1}. The minor axis of the corresponding ellipsoid is aligned with the polar axis of the Local Supercluster, and the major axis has an angle of $(29 \pm 5)^\circ$ with respect to the direction to the Virgo Cluster core. Therefore, the effect of local anisotropy generates a considerable difference in kinematic distances of galaxies seen in different directions.

Until recently, the majority of very nearby galaxies had no reliable direct distance estimates. In the 1990s, many neighboring spiral and irregular galaxies were resolved into stars for the first time, which allowed determination of their distances from the luminosity of blue and red supergiants with an accuracy of ~25%. During the last few years, about a hundred of nearby galaxies have been imaged with the WFPC2 on the *Hubble Space Telescope* (*HST*). Accurate distances to them (with an error ~10%) were measured from the luminosity of the tip of the red giant branch (TRGB). A compilation of 223 galaxies situated within 5.5 Mpc from us has been presented by Karachentsev et al. (2003b). About half of them have accurate distance estimates via TRGB or Cepheids. The ACS camera that has recently been installed on the *HST* is able to measure distances of galaxies to within 7–8 Mpc in a fast "snapshot" mode (one object per one orbit). This stimulated us to prepare a sample of galaxies with the expected distances within 7–8 Mpc, being as complete as possible. Taking into account the presence of non-Hubble motions, as well as distance measurement errors, we selected galaxies for our sample on the basis of two simple conditions:

$$D < 10 \text{ Mpc}, \tag{1}$$

if a galaxy has an individual distance estimate, or

$$V_{\rm LG} < 550 \text{ km s}^{-1}, \tag{2}$$

if the galaxy distance has been estimated from its radial velocity alone. Such an approach permits us to save a maximum number of galaxies in the sample with true distances $D <$ 8 Mpc, although it includes inevitably a fraction of more distant objects too. (We included also in the sample probable members of Leo 1 group having the mean distance of $D =$ 10.4 Mpc).

Here, the following property of any LV sample should be noted. In different catalogs of redshifts of galaxies, for instance, in the *Z* catalog by Huchra (1995), there are a lot of objects with radial velocities within $\sim[-200, +200]$ km s^{-1}. Part of them are superpositions of distant galaxy with a foreground star having a nearly zero velocity. For other objects the measured radial velocity corresponds to Galactic hydrogen but not to a galaxy. In the zone of avoidance in the Milky Way confusions often occur when a planetary nebula or cometary nebula looks like a nearby galaxy. For this reason, any sample of the LV galaxies contains a fraction of spurious cases and establishing their nature requires new special observations.

3. CATALOG

A list of 451 galaxies with measured and expected distances within 10 Mpc is presented in Table 1. Its columns contain the following data:

1. Galaxy name (Cetus, And III, etc.) or number in the known catalogs: NGC (N), IC (I), UGC (U, Nilson 1973), UGCA (UA, Nilson 1974), DDO (D, van den Bergh 1966), ESO/Uppsala (E, Lauberts 1982), PGC (P, Paturel et al. 1996), and Karachentsev's lists (KK, KKR, KKH, KKSG, KKs) usually in chronological preference.

2. and 3. Equatorial coordinates (J2000.0).

4. *a*.—Major angular diameter in arcminutes measured at a level of $\sim$25 mag arcsec^{-2}. For dwarf galaxies of low surface brightness from the KK lists, their diameters correspond to the Holmberg system ($\sim$26.5 mag arcsec^{-2}).

5. *b*/*a*.—Apparent axial ratio.

6.. B_t.—integrated apparent magnitude in the *B* band. As the main source of data on B_t, we use the NASA Extragalactic Database (NED) with additions from the recent publications by Makarova (1999), Bremnes, Binggeli, & Prugniel (1998, 1999, 2000), Parodi, Barazza, & Binggeli (2002), etc.. For some dwarf low surface brightness (LSB) galaxies, their magnitudes are known with an accuracy of $\sim$0.5 mag
Q3 only.

7. A_g.—Galactic extinction in the *B* band from Schlegel, Finkbeiner, & Davis (1998). Estimates of A_g for some galaxies situated in the zone of avoidance (for instance, Maffei I) were made from their individual color.

8. *T*.—Morphological type in numerical code according to de Vaucouleurs et al. (1991, RC3).

9. T_n.—Type of circumnuclear region: (K) kern, a stellar-like nucleus; (B) bulge, where a stellar-like nucleus may be hidden because of strong extinction or steep brightness gradient. (A) absence of a distinct nucleus. The main sources for the classification were 2MASS images (Jarrett et al. 2000) and high-resolution images of galaxies from the *HST* archive (see illustrations and details in Karachentsev & Karachentseva 2002).

10. V_h.—Heliocentric radial velocity in kilometers per second from the NED or from recent sources (Huchtmeier et al. 2000, 2001, 2003; Makarov et al., 2003).

11. V_{LG}.—Radial velocity with respect to the LG centroid with the apex parameters according to Karachentsev & Makarov (1996).

12. W_{50}.—H I line width at the 50% level from the maximum in kilometers per second taken from the Lyon Extragalactic Database (LEDA) or from Huchtmeier et al. (2000, 2001, 2003). For some southern dwarf galaxies their line widths were measured on the basis of the H I Parkes Sky Survey (HIPASS) data.[1]

13. log *F*.—Logarithm of the total H I flux (in Jy km s^{-1}) from the LEDA, Huchtmeier et al. 2000, 2001, 2003, or from the HIPASS.

14. and 15. *D*.—Distance to the galaxy in megaparsecs with indication of the method used: (cep) from the luminosity of cepheids; (rgb) from the luminosity of the tip of RGB; (sbf) via fluctuation of surface brightness; (bs) from the luminosity of the brightest stars; (tf) from the Tully-Fisher relation; (mem) from membership in the known groups; (h) from the Hubble relation $V_{LG} = H_0 D$ with the mean local Hubble parameter $H_0 = 72$ km s^{-1} Mpc^{-1} (Karachentsev et al. 2003b), coinciding with its global value (72 ± 8) km s^{-1} Mpc^{-1} (Freedman et al. 2001).

Table 2 contains references to the distance estimates and to the galaxy membership in the known groups. The references to photometric data are indicated as "ph". Q4

Table 3 lists 61 cases where the corrected radial velocity of an object satisfy condition (2), but the object is not a galaxy or, being a galaxy, has $D > 12$ Mpc. The column definitions are similar to the definitions in Table 1. A large group among them ($N = 13$) consists of probable high-velocity clouds (Henning et al. 2000; Juraszek et al. 2000; Kilborn et al. 2002), which have not been identified with galaxies brighter than $B_t \sim 20$ mag. However, some of the supposed high-velocity clouds (HVCs) are situated in a zone of strong extinction and may turn out to be nearby dwarf galaxies. Another category presents objects ($N = 15$) classified as Galactic H II regions (e.g., Cam C), reflection nebulae (e.g., MB 2) or planetary nebulae (e.g., KKH33), judging by their CCD images or spectra. Some objects in Table 3 are distant galaxies (Sc 18, UGC 4491, etc.), whose published low radial velocities originating from Galactic H I emission or from a foreground star. There are also six galaxies, NGC 1400, UGC 6782, UGC 7131, NGC 4150, UGC 7321, and UGC 7369, with low intrinsic velocities, but having distances within a range of 13–27 Mpc. Being located in the vicinity of the Virgo and Fornax clusters, they have probably a significant component of velocity toward the corresponding cluster.

Figure 1 presents the distribution of 451 galaxies from Table 1 on their distances. The bottom panel shows *N*(*D*) for the galaxies with accurate ($\sim$5%) distance estimates from cepheids. Surprisingly, in the LV there are only 24 such galaxies. Figures 1*b*–1*e* correspond to the galaxy samples with less accurate distance measurements. Figure 1*f* presents 68 galaxies whose distances have not been measured as yet, but have been estimated via their radial velocities with $H_0 = 72$ km s^{-1} Mpc^{-1}. The top panel gives the distance distribution for 80 galaxies (out of 451) without radial velocities. As can be seen, only $(24 + 154)/451 = 39\%$ of the whole LV population within 8 Mpc have quite reliable distance estimates (cep + rgb). In other words, our present understanding of the nearest universe is still rather incomplete.

For each galaxy from Table 1, we determined its linear diameter, luminosity, mass, and other integrated characteristics,

[1] The data accessible at http://www.atnf.CSIRO.AU.

TABLE 1
Catalog of LV Galaxies

Name (1)	R.A. (J2000.0) (2)	Decl. (J2000.0) (3)	a (4)	b/a (5)	B_t (6)	A_g (7)	T (8)	T_n (9)	V_h (10)	V_{LG} (11)	W_{50} (12)	log F (13)	D (14)	Method (15)
WLM, DDO 221	00 01 58.1	−15 27 40	11.5	0.35	11.03	0.16	9	A	−116	−10	57	2.30	0.92	rgb
E349−031, SDIG	00 08 13.3	−34 34 42	1.1	0.82	15.48	0.05	10	A	207	216	20	0.36	4.1	tf
N55	00 15 08.5	−39 13 13	32.4	0.17	8.84	0.06	8	A	129	111	172	3.43	1.8	tf
N59, KK 2	00 15 25.1	−21 26 38	2.7	0.48	12.97	0.09	−3	KB	361	431	50	0.43	5.30	sbf
E410−005, KK 3	00 15 31.4	−32 10 48	1.3	0.77	14.90	0.06	−1	A	...	...	...	...	1.92	rgb
I10	00 20 24.5	+59 17 30	6.8	0.87	12.20	3.65	10	A	−344	−60	63	2.98	0.66	cep
Sc 22	00 23 51.7	−24 42 18	0.9	0.78	17.73	0.06	−3	A	...	...	...	...	4.21	rgb
Cetus, KKSG 1	00 26 11.0	−11 02 40	5.0	0.86	14.4	0.12	−2	A	...	...	...	...	0.78	rgb
E294−010	00 26 33.3	−41 51 20	1.1	0.64	15.53	0.02	−3	A	117	81	...	...	1.92	rgb
U288	00 29 04.0	+43 25 54	1.3	0.62	15.64	0.33	10	A	187	464	36	0.71	6.7	bs
N147	00 33 11.6	+48 30 28	13.5	0.61	10.36	0.75	−3	A	−193	85	...	...	0.76	rgb
And III, KK 5	00 35 33.8	+36 29 52	3.0	0.60	15.20	0.10	−3	A	−355	−92	...	...	0.76	rgb
N185	00 38 58.0	+48 20 10	12.2	0.85	9.99	0.79	−3	A	−202	73	...	0.23	0.62	rgb
N205	00 40 22.5	+41 41 11	19.4	0.61	8.72	0.27	−5	K	−244	24	13	0.34	0.83	rgb
And IV	00 42 30.1	+40 34 33	1.3	0.77	16.60	0.27	10	A	256	521	...	...	6.11	rgb
N221	00 42 42.1	+40 51 59	9.0	0.71	8.73	0.26	−5	B	−145	121	...	...	0.77	rgb
M31	00 42 44.5	+41 16 09	189.0	0.33	4.17	0.27	3	B	−301	−35	510	4.53	0.77	cep
DDO 226	00 43 03.8	−22 15 01	2.2	0.36	14.36	0.07	10	A	357	408	48	1.01	4.92	rgb
And I, KK 8	00 45 40.0	+38 02 14	4.0	0.95	13.90	0.23	−3	A	−380	−120	...	...	0.81	rgb
N247	00 47 08.3	−20 45 36	21.4	0.32	9.86	0.08	7	K	160	215	210	2.83	4.1	tf
N253	00 47 34.3	−25 17 32	26.7	0.22	7.92	0.08	5	B	241	274	410	2.90	3.94	rgb
KDG 2, E540−030	00 49 21.1	−18 04 28	1.2	0.92	16.37	0.10	−1	A	...	...	...	...	3.40	rgb
DDO 6	00 49 49.3	−21 00 58	1.7	0.41	15.19	0.07	10	A	295	348	22	0.53	3.34	rgb
E540−032, FG 24	00 50 24.6	−19 54 25	1.3	0.92	16.44	0.09	−3	A	...	...	...	...	3.42	rgb
SMC	00 52 38.0	−72 48 01	319.	0.64	2.75	0.17	9	A	158	−22	90	4.95	0.06	cep
N300	00 54 53.5	−37 40 57	21.9	0.71	8.95	0.06	7	K	144	114	149	3.31	2.15	cep
Sculptor	01 00 09.4	−33 42 33	35.2	0.79	10.05	0.05	−3	A	110	96	...	0.11	0.09	rgb
LGS−3	01 03 56.6	+21 53 41	1.2	0.58	16.18	0.18	−1	A	−286	−74	18	0.23	0.62	rgb
I1613	01 04 54.1	+02 07 60	16.4	0.90	9.92	0.11	10	A	−232	−89	25	2.67	0.73	cep
U685	01 07 22.3	+16 41 02	1.4	0.71	14.22	0.25	9	A	155	349	72	1.00	4.79	rgb
KKH 5	01 07 32.5	+51 26 25	0.6	0.62	17.1	1.22	10	A	39	304	37	0.24	4.26	rgb
N404	01 09 26.9	+35 43 03	3.5	0.95	11.21	0.25	−1	B	−48	195	78	1.63	3.06	rgb
And V	01 10 17.1	+47 37 41	2.0	0.74	16.67	0.54	−3	A	−403	−143	...	...	0.81	rgb
And II, KK 12	01 16 29.8	+33 25 09	3.0	0.63	15.10	0.27	−3	A	−188	46	...	...	0.68	rgb
U1104	01 32 42.3	+18 18 57	1.0	0.60	14.48	0.27	10	A	685	866	67	0.99	7.5	bs
M33	01 33 50.8	+30 39 37	69.2	0.61	6.19	0.18	5	K	−180	36	184	4.09	0.85	cep
KKH 6	01 34 51.6	+52 05 30	0.8	0.60	17.0	1.52	10	A	17	270	31	0.61	3.8	h
N625	01 35 05.0	−41 26 11	6.4	0.28	11.59	0.07	9	A	405	335	90	1.55	4.07	rgb
N628	01 36 41.3	+15 47 12	10.1	0.94	9.79	0.30	5	K	657	827	53	2.74	7.3	bs
U1171	01 39 41.6	+15 54 11	1.3	0.92	15.71	0.25	10	A	738	906	25	0.15	7.3	bs
DDO 13, U1176	01 40 09.6	+15 54 17	4.6	0.78	14.33	0.28	10	A	630	798	37	1.45	9.0	bs
KK 13	01 42 16.8	+26 22 04	0.7	0.57	16.57	0.39	10	A	357	556	31	−0.05	7.2	mem
KDG 10	01 43 41.4	+15 41 23	1.5	0.67	16.28	0.32	10	A	789	953	32	0.55	7.9	bs
KK 14	01 44 42.7	+27 17 16	1.6	0.37	17.47	0.31	10	A	423	622	31	0.43	7.2	mem
E245−05	01 45 03.6	−43 35 53	3.8	0.89	12.73	0.07	9	A	394	308	61	1.90	4.43	rgb
KK 15	01 46 41.6	+26 48 05	0.6	0.33	18.22	0.36	10	A	366	563	22	−0.10	7.2	mem
I1727, U1249	01 47 30.1	+27 19 52	6.9	0.36	12.06	0.34	9	A	338	535	118	2.08	7.2	bs
N672	01 47 53.2	+27 26 01	7.2	0.36	11.33	0.34	6	A	421	618	205	2.16	7.2	bs
U1281	01 49 32.3	+32 35 33	4.5	0.18	13.03	0.20	7	A	157	367	119	1.55	5.4	bs
Phoenix, P6830	01 51 06.3	−44 26 41	4.7	0.85	13.07	0.07	−1	A	−13	−106	21	0.14	0.44	rgb
KK 16	01 55 20.6	+27 57 15	0.8	0.37	16.3	0.29	10	A	207	400	24	−0.15	4.74	rgb
KK 17	02 00 09.9	+28 49 57	0.6	0.50	17.20	0.24	10	A	168	360	20	−0.40	4.72	rgb
N784	02 01 16.8	+28 50 37	6.6	0.23	12.16	0.26	8	A	194	386	92	1.78	5.0	bs
Cas 1, KK 19	02 06 07.9	+69 00 36	2.2	0.77	15.29	4.40	10	A	35	284	50	1.70	3.3	mem
N855	02 14 03.7	+27 52 38	2.6	0.38	13.30	0.31	−5	BK	595	774	115	0.49	9.73	sbf
N891	02 22 32.8	+42 20 48	13.1	0.29	10.71	0.28	3	B	528	738	448	2.08	9.77	sbf
KKH 11	02 24 35.0	+56 00 42	1.4	0.54	16.2	2.13	10	A	75	308	92	1.36	3.0	mem
KKs 3	02 24 42.7	−73 30 46	2.5	0.40	16.0:	0.22	−3	A	...	...	...	...	4.1	mem
N925	02 27 16.0	+33 34 41	11.1	0.56	10.58	0.33	7	A	553	738	200	2.44	9.29	cep
KKH 12	02 27 27.0	+57 29 16	1.7	0.29	17.80	3.44	10	A	70	303	52	1.21	3.0	mem
MB 1, KK 21	02 35 35.6	+59 22 47	2.4	0.42	16.80	4.22	7	A	190	421	50	0.91	3.0	mem
Maffei 1	02 36 35.5	+59 39 18	2.6	0.85	13.47	5.05	−3	B	15	246	...	...	3.0	fj
E115−21	02 37 45.0	−61 20 28	7.3	0.14	13.34	0.11	6	A	513	337	114	1.98	4.66	rgb

TABLE 1—*Continued*

Name (1)	R.A. (J2000.0) (2)	Decl. (J2000.0) (3)	a (4)	b/a (5)	B_t (6)	A_g (7)	T (8)	T_n (9)	V_h (10)	V_{LG} (11)	W_{50} (12)	log F (13)	D (14)	Method (15)
Fornax	02 39 54.7	−34 31 33	30.0	0.78	9.28	0.05	−3	A	53	−32	. . .	−0.40	0.14	rgb
Maffei 2	02 41 54.5	+59 36 11	3.8	0.50	14.77	7.19	4	B	−17	212	305	2.47	2.8	tf
Dw 2	02 54 08.5	+59 00 19	1.0	0.40	17.97	5.13	10	A	94	316	90	1.98	3.0	mem
MB 3, KK 22	02 55 43.6	+58 51 42	1.6	0.31	19.38	5.64	10	A	59	280	15	0.00	3.0	mem
E154−23	02 56 52.1	−54 34 23	8.2	0.18	12.69	0.07	8	A	578	412	124	2.15	5.7	h
Dw 1, KK 23	02 56 56.1	+58 54 42	2.0	0.65	15.01	6.34	3	A	112	333	187	2.31	2.8	tf
N1156	02 59 42.4	+25 14 15	3.3	0.85	12.75	0.97	9	A	376	510	63	1.78	7.8	bs
KKH 18	03 03 05.9	+33 41 40	0.7	0.57	16.7	0.86	10	A	216	375	34	0.48	4.43	rgb
N1313	03 18 15.4	−66 29 51	9.0	0.80	9.66	0.47	7	K	475	270	156	2.66	4.15	rgb
N1311	03 20 07.4	−52 11 06	3.2	0.25	13.18	0.09	9	A	570	398	110	0.86	5.5	h
U2684	03 20 23.7	+17 17 42	1.8	0.50	16.29	0.60	10	A	350	438	79	1.00	6.5	bs
KK 27	03 21 05.7	−66 19 22	1.2	0.33	16.5	0.33	−3	A	. . .	. . .	. . .	. . .	3.98	rgb
U2716	03 24 07.2	+17 45 15	1.6	0.56	14.64	0.59	8	A	381	467	44	0.84	6.5	h
U2773, KK 28	03 32 07.1	+47 47 37	1.2	0.75	15.04	2.43	10	B	218	397	79	1.42	5.5	h
I1959	03 33 11.8	−50 24 38	3.0	0.23	13.26	0.05	9	A	639	464	130	1.13	6.4	h
KK 35	03 45 12.6	+67 51 50	2.5	0.68	15.7	2.50	10	A	105	320	34	−0.10	3.16	rgb
I342	03 46 49.7	+68 05 45	21.2	0.91	9.37	2.41	6	B	31	245	151	3.69	3.28	cep
U2905	03 57 00.6	+16 31 28	1.0	0.60	15.76	1.35	10	A	292	344	48	0.36	5.8	bs
UA 86	03 59 49.5	+67 07 31	4.5	0.69	13.5	4.06	10	A	67	275	83	2.48	2.65	bs
I2038	04 08 54.1	−55 59 32	1.7	0.24	14.98	0.05	7	A	712	505	110	1.38	7.0	h
Cam A, KK 41	04 25 15.6	+72 48 21	3.7	0.57	14.84	0.93	10	A	−47	164	44	1.66	3.93	rgb
N1569	04 30 49.1	+64 50 53	3.6	0.51	11.30	3.02	9	A	−104	88	74	1.87	1.95	bs
UA 92	04 32 00.3	+63 36 50	2.0	0.50	15.22	3.42	10	A	−99	89	77	2.02	1.80	bs
N1560	04 32 49.9	+71 52 52	9.8	0.15	11.90	0.81	7	A	−36	171	125	2.65	3.45	rgb
Cam B, KK 44	04 53 06.9	+67 05 57	2.2	0.50	16.71	0.94	10	A	77	266	20	0.70	3.34	rgb
N1705	04 54 13.7	−53 21 41	1.9	0.74	12.76	0.03	9	A	627	400	94	0.90	5.10	rgb
UA 105	05 14 15.1	+62 34 51	5.5	0.64	12.14	1.35	9	A	111	279	118	2.14	3.15	rgb
LMC	05 23 34.6	−69 45 22	646.	0.85	0.9	0.32	9	A	278	28	65	6.00	0.05	cep
U3303	05 24 59.5	+04 30 18	3.6	0.78	13.95	0.56	8	A	521	446	163	1.47	7.2	bs
KK 49	05 41 41.5	+06 40 54	0.7	0.71	16.1	2.46	10	A	455	376	55	0.95	5.2	h
Orion	05 45 02.0	+05 04 06	2.0	0.48	15.40	3.16	8	A	365	276	159	1.89	6.4	bs
A0554+07	05 57 36.7	+07 29 31	0.5	0.40	18.4	2.55	10	A	428	340	50	0.45	5.5	bs
KKH 34, Mai 13	05 59 41.2	+73 25 39	0.9	0.56	17.1	1.08	10	A	110	299	24	0.38	4.61	rgb
E364−29, KKs 18	06 05 45.4	−33 04 54	3.5	0.71	13.58	0.19	10	A	787	555	65	1.20	7.7	h
E121−20, KKs 19	06 15 54.5	−57 43 35	1.4	0.71	15.85	0.17	10	A	577	309	65	1.15	4.3	h
E489−56, KK 54	06 26 17.0	−26 15 56	0.6	0.50	15.70	0.28	10	A	492	263	24	0.32	4.99	rgb
U3476	06 30 29.2	+33 18 07	1.0	0.30	15.98	1.02	10	A	469	477	83	1.09	7.0	bs
E490−17	06 37 56.6	−25 59 59	1.7	0.76	14.01	0.34	10	A	504	268	41	0.81	4.23	rgb
KK 55	06 39 32.9	−40 43 13	0.7	0.64	16.05	0.33	10	A	821	556	51	0.10	7.7	h
Carina	06 41 36.7	−50 57 58	24.3	0.67	11.30	0.27	−3	A	223	−53	. . .	. . .	0.10	rgb
KKSG 9	06 46 56.9	−17 56 29	0.5	0.70	17.2	1.91	9	A	696	477	187	1.41	6.6	h
KKH 37, Mai 16	06 47 45.8	+80 07 26	1.2	0.67	16.4	0.33	10	A	1	204	20	0.23	2.8	h
U3600	06 55 40.0	+39 05 42	1.4	0.36	16.18	0.39	10	A	412	435	84	0.74	7.3	bs
HIZSS 003	07 00 29.3	−04 12 30	2.	0.60	18	4.66	10	A	280	101	65	1.51	1.4	h
FG 202	07 05 17.1	−58 31 14	3.4	0.47	14.95	0.51	10	A	554	269	78	1.42	4.90	rgb
E558−11, KK 57	07 06 56.8	−22 02 26	1.8	0.56	14.43	1.64	9	A	738	496	116	0.86	6.9	h
U3698	07 09 18.8	+44 22 48	1.0	0.60	15.41	0.42	10	A	421	464	38	0.70	7.2	bs
N2337	07 10 13.6	+44 27 25	2.0	0.75	13.48	0.38	9	A	434	477	144	1.60	7.9	bs
U3755	07 13 51.8	+10 31 19	1.7	0.59	14.07	0.38	10	A	315	190	33	0.83	5.22	rgb
U3817	07 22 44.5	+45 06 30	1.8	0.50	15.96	0.44	10	A	437	478	35	1.00	8.6	bs
DDO 43, U3860	07 28 17.2	+40 46 13	1.3	0.69	14.96	0.25	10	A	354	371	35	1.06	7.81	rgb
N2366	07 28 52.0	+69 12 19	7.3	0.36	11.68	0.16	10	A	99	253	96	2.47	3.19	rgb
E059−01	07 31 19.3	−68 11 10	2.1	0.81	13.98	0.63	9	A	528	245	110	1.13	3.4	h
DDO 44, KK 61	07 34 11.3	+66 53 10	3.0	0.67	15.64	0.19	−3	A	. . .	. . .	. . .	. . .	3.19	rgb
N2403	07 36 54.4	+65 35 58	21.9	0.56	8.82	0.18	6	A	131	268	231	3.15	3.30	cep
DDO 46, U3966	07 41 25.8	+40 06 45	1.8	0.94	14.40	0.22	10	A	360	369	71	1.40	6.8	bs
DDO 47, U3974	07 41 55.0	+16 48 02	3.1	0.85	13.62	0.14	10	A	270	160	56	1.88	5.18	rgb
KK 65	07 42 31.2	+16 33 40	0.9	0.56	15.26	0.14	10	A	279	168	37	0.42	4.51	rgb
U4115	07 57 01.8	+14 23 27	1.8	0.56	15.23	0.12	10	A	338	210	83	1.32	5.49	rgb
N2537	08 13 14.7	+45 59 26	1.7	0.88	12.82	0.23	9	A	447	477	96	0.98	6.9	bs
Ho II, DDO 50	08 19 05.9	+70 42 51	7.9	0.86	11.09	0.14	10	A	157	311	66	2.56	3.39	rgb
E006−01, KK 68	08 19 23.3	−85 08 44	1.1	0.91	15.13	0.83	9	A	738	489	. . .	. . .	6.8	h
KDG 52	08 23 56.0	+71 01 46	1.3	0.92	16.35	0.09	10	A	113	268	22	0.65	3.55	rgb
DDO 52, U4426	08 28 28.5	+41 51 24	2.0	0.50	15.27	0.16	10	A	393	398	83	0.95	5.7	bs
DDO 53	08 34 06.5	+66 10 45	1.6	0.87	14.55	0.16	10	A	20	151	25	1.14	3.56	rgb

TABLE 1—*Continued*

Name (1)	R.A. (J2000.0) (2)	Decl. (J2000.0) (3)	a (4)	b/a (5)	B_t (6)	A_g (7)	T (8)	T_n (9)	V_h (10)	V_{LG} (11)	W_{50} (12)	log F (13)	D (14)	Method (15)
U4483	08 37 03.0	+69 46 31	1.2	0.58	14.95	0.15	10	A	156	304	33	1.13	3.21	rgb
N2683	08 52 40.9	+33 25 02	9.3	0.24	10.34	0.14	3	B	409	363	425	1.73	7.73	sbf
KK 69	08 52 50.7	+33 47 52	2.0	0.75	16.8	0.13	10	A	463	419	17	0.45	7.7	mem
KK 70	08 55 22.0	+33 33 33	0.5	0.80	17.7	0.13	−3	A	...	...	...	...	7.7	mem
D564−08	09 02 54.0	+20 04 31	0.7	0.71	16.90	0.12	10	A	483	365	49	0.30	5.1	h
KKH 46	09 08 36.6	+05 17 32	0.7	0.86	17.05	0.20	10	A	598	409	25	0.43	5.7	h
D634−03	09 08 53.5	+14 34 55	0.4	0.88	17.5	0.16	10	A	319	173	34	...	2.4	h
KK 71	09 09 10.9	−23 22 04	0.4	0.78	18.7	0.77	−3	A	...	...	...	...	9.8	mem
DDO 56, UA 148	09 09 46.6	−23 00 33	1.3	0.77	15.63	0.72	10	A	724	438	76	0.49	9.8	mem
KK 72	09 11 42.3	−23 58 56	0.6	0.83	18	0.80	−3	A	...	...	...	...	9.8	mem
N2784 dw 1	09 12 18.5	−24 12 40	0.9	0.89	17.27	0.89	−3	A	...	...	...	...	9.8	mem
N2784	09 12 19.4	−24 10 18	5.5	0.40	11.17	0.93	−2	K	691	402	...	...	9.82	sbf
KK 73	09 12 29.3	−24 14 28	0.8	0.62	16.35	0.85	−3	A	...	...	...	...	9.8	mem
E564−30, UA 153	09 13 12.1	−19 24 31	2.0	0.60	15.40	0.38	9	A	765	488	125	1.66	6.8	h
N2787	09 19 18.6	+69 12 12	3.2	0.63	11.82	0.56	1	B	696	838	358	0.93	7.48	sbf
D565−06	09 19 29.4	+21 36 12	0.7	0.86	16.95	0.17	10	A	498	386	14	−0.52	5.4	h
E565−03, KKs 30	09 23 09.9	−20 10 03	0.6	0.67	15.53	0.27	10	A	829	549	31	0.20	7.6	h
N2915	09 26 11.5	−76 37 35	1.9	0.53	13.20	1.19	10	A	460	184	135	1.94	3.78	rgb
KKH 51, KDG 56	09 30 12.9	+19 59 30	0.7	0.80	17.02	0.19	10	A	561	440	25	−0.10	8.9	mem
N2903	09 32 09.6	+21 30 02	12.6	0.48	9.47	0.13	4	B	555	442	371	2.35	8.9	bs
U5086	09 32 48.9	+21 27 54	1.0	0.90	15.94	0.14	−3	A	...	...	...	0.38	8.9	mem
Ho I, DDO 63	09 40 28.2	+71 11 11	3.6	0.81	13.64	0.21	10	A	139	291	29	1.59	3.84	rgb
U5186	09 42 59.8	+33 15 52	1.3	0.23	16.27	0.06	10	A	551	500	42	0.18	6.9	h
F8D1	09 44 50.0	+67 28 32	2.2	0.91	15.7	0.41	−3	A	...	...	...	...	3.77	rgb
U5209, KKH 54	09 45 04.2	+32 14 18	0.9	0.96	16.06	0.08	10	A	538	482	27	0.04	6.7	h
FM 1	09 45 10.0	+68 45 54	0.9	0.89	17.50	0.31	−3	A	...	...	...	...	3.42	rgb
N2976	09 47 15.6	+67 54 49	5.9	0.46	11.01	0.30	5	K	3	139	97	1.80	3.56	rgb
KK 77	09 50 10.0	+67 30 24	2.4	0.75	16.3	0.62	−3	A	...	...	...	...	3.48	rgb
KK 78	09 50 19.6	+31 27 24	0.5	0.60	17.68	0.09	10	A	541	481	...	...	7.1	mem
DDO 64, U5272	09 50 22.4	+31 29 15	2.1	0.38	15.41	0.10	10	A	520	460	82	1.13	7.1	bs
U5288	09 51 17.2	+07 49 38	1.3	0.62	14.62	0.15	10	A	556	377	93	1.40	6.8	bs
BK 3N	09 53 48.5	+68 58 09	0.5	0.80	18.78	0.35	10	A	−40	101	15	...	4.02	rgb
KKSG 15	09 55 10.5	−06 16 12	1.3	0.23	15.14	0.18	10	A	791	554	54	0.43	9.7	mem
M81	09 55 33.5	+69 03 60	26.9	0.52	7.69	0.36	3	B	−35	107	422	2.97	3.63	cep
M82	09 55 53.9	+69 40 57	11.2	0.38	9.06	0.69	8	B	202	347	147	2.43	3.53	rgb
DDO 68, U5340	09 56 45.7	+28 49 30	2.7	0.37	14.60	0.08	10	A	502	428	78	1.28	5.9	bs
KDG 61, KK 81	09 57 02.7	+68 35 30	2.4	0.58	15.24	0.31	−1	A	−116	23	...	...	3.60	rgb
A0952+69	09 57 29.0	+69 16 20	1.8	0.90	16.80	0.37	10	A	100	243	...	...	3.87	rgb
Ho IX, DDO 66	09 57 32.4	+69 02 35	2.5	0.80	14.53	0.35	10	A	46	188	69	2.20	3.7	mem
Leo A, DDO 69	09 59 26.4	+30 44 47	5.1	0.61	12.92	0.09	10	A	24	−40	19	1.68	0.69	rgb
KKSG 16	09 59 46.5	−09 21 04	0.9	0.78	18	0.31	−3	A	...	...	...	...	9.7	mem
Sex B, DDO 70	10 00 00.1	+05 19 56	5.1	0.69	11.85	0.14	10	A	301	111	38	2.03	1.36	rgb
KKH 57, HS 108	10 00 16.0	+63 11 06	0.6	0.83	17.87	0.09	−3	A	...	...	...	...	3.93	rgb
KKSG 17	10 01 38.4	−08 14 56	1.4	0.21	15.5	0.26	10	A	447	203	49	0.30	9.7	mem
UA 193	10 02 36.2	−06 00 43	4.3	0.07	14.83	0.17	7	A	661	425	130	1.33	9.7	mem
N3109	10 03 07.2	−26 09 36	19.7	0.19	10.39	0.29	9	A	403	110	116	3.22	1.33	rgb
N3077	10 03 21.0	+68 44 02	5.4	0.83	10.46	0.29	10	B	13	153	65	2.26	3.82	rgb
Antlia, P29194	10 04 04.0	−27 19 55	2.0	0.75	16.19	0.34	10	A	362	66	22	0.49	1.32	rgb
U5427	10 04 41.0	+29 21 59	1.2	0.67	14.89	0.10	8	A	495	424	78	0.42	7.1	bs
BK 5N	10 04 40.3	+68 15 20	0.8	0.75	17.53	0.25	−3	A	...	...	...	...	3.78	rgb
KDG 63, KK 83	10 05 07.3	+66 33 18	1.7	0.88	16.01	0.41	−3	A	−129	0	19	...	3.50	rgb
N3115	10 05 14.0	−07 43 07	7.2	0.35	9.86	0.20	−3	K	720	478	238	0.30	9.68	sbf
U5423	10 05 30.6	+70 21 52	0.9	0.67	14.42	0.34	10	A	348	496	44	0.58	5.3	bs
KK 84	10 05 34.4	−07 44 57	1.3	0.69	16.36	0.03	−3	A	...	...	...	...	9.7	mem
KKSG 18, N3115B	10 05 41.6	−07 58 53	1.5	0.80	13.61	0.23	−2	KB	698	456	...	−0.40	9.7	mem
KDG 64, KK 85	10 07 01.9	+67 49 39	1.9	0.47	15.52	0.25	−3	A	...	...	...	...	3.70	rgb
U5456	10 07 19.7	+10 21 44	1.6	0.50	13.84	0.18	9	A	544	377	61	0.47	3.8:	rgb
IKN	10 08 05.9	+68 23 57	2.7	0.85	17.0	0.60	−3	A	...	...	...	...	3.7	mem
Leo I, DDO 74	10 08 26.9	+12 18 29	9.8	0.76	11.18	0.16	−3	A	285	128	...	...	0.25	rgb
Sex A, DDO 75	10 11 00.8	−04 41 34	5.9	0.83	11.86	0.19	10	A	324	94	63	2.32	1.32	cep
Sex dSph	10 13 03.0	−01 36 52	30.	0.40	12.00	0.21	−3	A	226	8	...	...	0.09	rgb
KKH 60	10 15 59.4	+06 48 17	0.8	0.55	18	0.09	10	A	290	108	...	...	1.5	h
HIJASS	10 21 00.2	+68 41 60	...	...	20	0.09	13	A	46	187	...	...	3.7	mem
HS 117	10 21 25.2	+71 06 58	1.5	0.60	16.5	0.49	10	A	−37	116	13	...	3.7	mem
KKSG 19	10 24 28.3	−12 25 57	0.7	0.86	16.79	0.30	10	A	628	373	28	0.52	5.2	h

TABLE 1—*Continued*

Name (1)	R.A. (J2000.0) (2)	Decl. (J2000.0) (3)	a (4)	b/a (5)	B_t (6)	A_g (7)	T (8)	T_n (9)	V_h (10)	V_{LG} (11)	W_{50} (12)	log F (13)	D (14)	Method (15)
DDO 78, KK 89	10 26 27.9	+67 39 24	2.0	0.95	15.8	0.12	−3	A	55	191	. . .	. . .	3.72	rgb
U5672	10 28 20.9	+22 34 17	1.8	0.28	14.45	0.10	10	A	531	428	82	0.43	6.3	bs
I2574	10 28 22.4	+68 24 58	13.2	0.41	10.84	0.16	9	A	57	197	115	2.65	4.02	rgb
DDO 82	10 30 35.0	+70 37 10	3.4	0.59	13.57	0.19	9	A	56	207	. . .	. . .	4.00	rgb
LeG 01	10 31 53.8	+12 55 35	0.5	0.60	18.7	0.14	10	A	. . .	. . .	. . .	. . .	10.4	mem
N3274	10 32 17.1	+27 40 07	2.1	0.48	13.25	0.10	6	KB	537	461	157	1.74	6.5	bs
LeG 02	10 33 19.4	+10 11 21	0.3	0.95	19.1	0.13	10	A	. . .	. . .	. . .	. . .	10.4	mem
BK 6N, KK 91	10 34 31.9	+66 00 42	1.1	0.64	16.9	0.05	−3	A	. . .	. . .	. . .	. . .	3.85	rgb
LeG 03	10 35 48.8	+08 28 47	1.1	0.23	17.8	0.11	9	A	. . .	. . .	. . .	. . .	10.4	mem
N3299	10 36 23.8	+12 42 27	2.2	0.77	13.30	0.11	8	A	641	490	120	0.46	10.4	mem
LeG 04	10 39 40.2	+12 44 06	0.6	0.50	18.7	0.12	10	A	. . .	. . .	. . .	. . .	10.4	mem
FS 01, LeG 05	10 39 43.3	+12 38 04	0.7	0.53	16.77	0.11	−3	A	. . .	. . .	. . .	. . .	10.4	mem
P031727	10 39 55.7	+13 54 28	0.6	0.65	18.3	0.15	10	A	1010	865	27	−0.26	10.4	mem
U5812	10 40 56.6	+12 28 19	1.5	0.40	15.10	0.10	10	A	1008	857	56	0.17	10.4	mem
P031877	10 42 00.3	+12 20 06	1.3	0.46	15.7	0.11	0	A	778	627	47	−0.03	10.4	mem
LeG 09	10 42 34.6	+12 09 02	0.8	0.83	18.5	0.11	−2	A	. . .	. . .	. . .	. . .	10.4	mem
N3344	10 43 30.2	+24 55 25	7.0	0.90	10.38	0.14	4	BK	586	498	154	2.27	6.9	h
LeG 10	10 43 55.4	+12 08 07	0.1	0.95	19.2	0.11	10	A	. . .	. . .	. . .	. . .	10.4	mem
N3351	10 43 57.7	+11 42 13	7.4	0.68	10.53	0.12	4	B	778	624	268	1.45	10.05	cep
LeG 11	10 44 02.1	+15 35 20	0.5	0.90	18.8	0.14	10	A	. . .	. . .	. . .	. . .	10.4	mem
LeG 12	10 44 07.8	+11 31 59	0.2	0.80	19.1	0.13	10	A	. . .	. . .	. . .	. . .	10.4	mem
FS 09, LeG 13	10 44 57.6	+11 54 58	0.4	0.44	17.43	0.10	10	A	. . .	. . .	. . .	. . .	10.4	mem
FS 13, LeG 14	10 46 14.4	+12 57 36	0.4	0.88	18.7	0.10	−3	A	886	739	. . .	. . .	10.4	mem
FS 14, KK 93	10 46 24.8	+14 01 30	0.6	0.92	18.3	0.13	10	A	. . .	. . .	. . .	. . .	10.4	mem
FS 15, LeG 16	10 46 30.0	+11 45 21	0.4	0.98	19.0	0.11	−2	A	. . .	. . .	. . .	. . .	10.4	mem
FS 17, LeG 17	10 46 41.5	+12 19 35	0.7	0.78	16.98	0.10	10	A	. . .	. . .	. . .	. . .	10.4	mem
N3368	10 46 45.8	+11 49 11	7.6	0.68	10.11	0.11	3	BK	897	744	341	1.53	10.42	cep
LeG 18	10 46 53.3	+12 44 40	0.8	0.75	18.9	0.10	10	A	617	469	28	−0.52	10.4	mem
FS 20, LeG 19	10 46 54.8	+12 47 17	0.5	0.70	18.2	0.10	10	A	. . .	. . .	. . .	. . .	10.4	mem
FS 21, KK 94	10 46 57.3	+12 59 54	0.8	0.85	17.5	0.13	10	A	832	685	32	0.20	10.4	mem
LeG 21	10 47 00.8	+12 57 35	0.4	0.75	18.6	0.13	10	A	. . .	. . .	. . .	. . .	10.4	mem
DDO 88, U5889	10 47 22.3	+14 04 13	2.2	0.95	14.25	0.15	9	A	573	431	43	0.79	9.3	bs
FS 23	10 47 27.4	+13 53 23	0.4	0.75	17.79	0.15	10	A	573	431	. . .	. . .	10.4	mem
P32250	10 47 41.9	−38 51 15	2.2	0.14	14.97	0.33	7	A	720	415	55	0.87	5.8	h
N3377	10 47 42.4	+13 59 08	5.2	0.58	11.24	0.15	−5	B	665	523	. . .	. . .	10.91	sbf
N3379	10 47 49.6	+12 34 54	5.4	0.89	10.24	0.11	−3	B	911	762	. . .	. . .	11.12	sbf
N3384	10 48 16.9	+12 37 45	5.5	0.45	10.85	0.12	−1	B	704	556	39	0.27	11.38	sbf
P032327	10 48 43.4	+12 18 56	0.2	0.80	17.51	0.11	−2	A	887	737	64	−0.26	10.4	mem
P032348	10 48 53.7	+14 07 28	1.2	0.50	15.5	0.17	−1	A	637	496	. . .	. . .	10.4	mem
DDO 87, U5918	10 49 36.5	+65 31 50	2.4	0.90	15.00	0.05	10	A	340	468	64	0.92	7.4	bs
FS 40, LeG 22	10 49 37.1	+11 21 06	0.6	0.66	18.0	0.13	10	A	. . .	. . .	. . .	. . .	10.4	mem
P1424345	10 49 52.2	+13 09 42	0.1	0.90	20	0.15	10	A	754	608	. . .	. . .	10.4	mem*
LeG 23	10 50 09.1	+13 29 01	0.3	0.95	19.1	0.14	−3	A	. . .	. . .	. . .	. . .	10.4	mem
U5944	10 50 19.1	+13 16 19	1.1	0.97	14.8	0.13	−3	A	. . .	. . .	. . .	. . .	10.4	mem
KK 96	10 50 27.1	+12 21 39	1.0	0.80	18.3	0.11	−3	A	. . .	. . .	. . .	. . .	10.4	mem
N3412	10 50 53.3	+13 24 44	3.6	0.56	11.45	0.12	−1	B	841	697	. . .	. . .	10.4	mem
LeG 26	10 51 21.1	+12 50 57	0.9	0.89	17.2	0.10	−2	A	. . .	. . .	. . .	. . .	10.4	mem
LeG 27	10 52 20.1	+14 42 26	0.4	0.67	18.6	0.09	10	A	. . .	. . .	. . .	. . .	10.4	mem
KDG 73	10 52 55.3	+69 32 45	0.6	0.83	17.09	0.08	10	A	116	263	18	0.00	3.70	rgb
LeG 28	10 53 00.7	+10 22 45	0.6	0.50	18.3	0.11	10	A	. . .	. . .	. . .	. . .	10.4	mem
D640−16	10 55 03.6	+14 05 36	0.3	0.83	18.6	0.10	10	A	. . .	. . .	. . .	. . .	10.4	mem
D640−12	10 55 57.1	+12 20 20	0.8	0.50	18.4	0.08	10	A	. . .	. . .	. . .	. . .	10.4	mem
D640−13	10 56 14.0	+12 00 37	0.7	0.57	17.66	0.08	10	A	990	841	. . .	. . .	10.4	mem
E215−09, KKs 40	10 57 30.2	−48 10 44	2.0	0.60	16.03	0.95	10	A	598	290	52	1.86	4.0	h
D640−14	10 58 10.5	+11 59 57	0.6	0.66	18.5	0.07	10	A	. . .	. . .	. . .	. . .	10.4	mem
LeG 32	10 59 17.3	+15 05 07	0.4	0.75	18.7	0.09	10	A	. . .	. . .	. . .	. . .	10.4	mem
N3489	11 00 18.6	+13 54 04	3.5	0.57	11.12	0.07	1	B	677	538	307	−0.06	10.4	mem
LeG 33	11 00 45.2	+14 10 20	0.5	0.55	18.6	0.08	10	A	. . .	. . .	. . .	. . .	10.4	mem
D640−08	11 00 52.1	+13 52 51	1.2	0.58	16.96	0.07	−3	A	. . .	. . .	. . .	. . .	10.4	mem
LeG 35	11 03 02.1	+08 02 54	0.6	0.50	18.1	0.16	10	A	. . .	. . .	. . .	. . .	10.4	mem
CGCG 66−109	11 04 26.5	+11 45 18	1.4	0.29	15.7	0.06	10	A	. . .	. . .	. . .	. . .	10.4	mem
U6145, KKSG 21	11 05 35.0	−01 51 49	1.7	0.47	16.45	0.24	10	A	740	533	56	0.30	7.4	h
Leo II, DDO 93	11 13 29.2	+22 09 17	12.0	0.92	12.45	0.07	−3	A	76	−18	. . .	. . .	0.21	rgb
N3593	11 14 35.9	+12 49 07	5.3	0.43	11.83	0.08	0	B	628	489	194	0.93	6.8	h
N3621	11 18 16.1	−32 48 42	12.4	0.44	10.18	0.35	7	K	726	436	266	2.89	6.70	cep

TABLE 1—*Continued*

Name (1)	R.A. (J2000.0) (2)	Decl. (J2000.0) (3)	a (4)	b/a (5)	B_t (6)	A_g (7)	T (8)	T_n (9)	V_h (10)	$V_{\rm LG}$ (11)	W_{50} (12)	log F (13)	D (14)	Method (15)
N3627	11 20 15.0	+12 59 29	9.1	0.46	9.65	0.14	4	B	723	586	340	1.28	10.28	cep
U6456	11 28 00.6	+78 59 29	1.5	0.53	14.32	0.16	10	A	−103	89	22	1.15	4.34	rgb
U6541	11 33 29.1	+49 14 17	1.4	0.57	14.32	0.08	10	A	250	304	29	0.48	3.89	rgb
N3738	11 35 48.6	+54 31 22	2.6	0.73	11.92	0.05	10	A	228	305	78	1.34	4.90	rgb
N3741	11 36 06.4	+45 17 07	2.0	0.55	14.38	0.10	10	A	230	264	81	1.72	3.03	rgb
E320−14, KKs 44	11 37 53.4	−39 13 14	1.4	0.64	15.85	0.62	10	A	654	362	20	0.42	5.0	h
KK 109	11 47 11.2	+43 40 19	0.6	0.67	18.62	0.08	10	A	212	241	19	−0.16	4.51	rgb
DDO 99, U6817	11 50 53.0	+38 52 50	4.1	0.37	13.70	0.11	10	A	242	248	37	1.50	2.64	rgb
E379−07, KK 112	11 54 43.0	−33 33 29	1.1	0.73	16.60	0.32	10	A	640	363	60	0.41	5.22	rgb
BTS 76	11 58 44.1	+27 35 06	0.3	0.60	16.5	0.10	10	A	499	451	. . .	. . .	6.3	h
N4068	12 04 02.4	+52 35 19	3.2	0.53	13.19	0.09	10	A	210	290	51	1.49	5.2	bs
MCG 6−27−17	12 09 56.4	+36 26 07	0.8	0.31	15.5	0.11	10	A	339	341	. . .	. . .	4.7	h
N4144	12 09 59.3	+46 27 26	6.1	0.21	12.16	0.06	6	A	267	319	150	1.68	9.8	bs
N4163	12 12 08.9	+36 10 10	1.9	0.84	13.63	0.09	10	A	163	164	30	0.98	3.6	bs
KK 127	12 13 22.7	+29 55 18	1.0	0.40	15.37	0.08	10	A	135	105	38	0.34	1.5	h
N4190	12 13 44.6	+36 37 60	1.7	0.88	13.52	0.13	10	A	230	234	46	1.37	3.5	bs
E321−014	12 13 49.6	−38 13 53	1.4	0.43	15.22	0.40	10	A	613	337	25	0.46	3.19	rgb
U7242, KKH 77	12 14 07.4	+66 05 32	1.9	0.42	14.60	0.08	10	A	68	213	61	0.86	4.3	mem
DDO 113, KDG 90	12 14 57.9	+36 13 08	1.5	0.87	15.70	0.09	10	A	280	283	. . .	. . .	2.86	rgb
N4214	12 15 38.9	+36 19 39	8.5	0.78	10.24	0.09	10	A	291	295	62	2.51	2.94	rgb
MCG 9−20−131	12 15 46.7	+52 23 15	1.2	0.30	15.3	0.10	10	A	162	245	. . .	. . .	3.4	h
U7298	12 16 28.6	+52 13 38	1.1	0.55	15.95	0.10	10	A	173	255	23	0.69	4.21	rgb
N4236	12 16 43.3	+69 27 56	22.6	0.32	10.06	0.06	8	A	0	160	162	2.79	4.45	rgb
N4244	12 17 29.9	+37 48 27	19.4	0.11	10.67	0.09	6	K	243	255	204	2.63	4.49	rgb
I3104, E020−04	12 18 46.1	−79 43 34	3.8	0.47	13.63	1.70	9	A	430	171	41	1.08	2.27	rgb
N4258	12 18 57.5	+47 18 14	18.6	0.39	9.10	0.07	4	B	448	507	421	2.65	7.83	cep
U7356	12 19 09.1	+47 05 23	0.9	0.99	15.75	0.08	10	A	272	330	. . .	. . .	6.7	sbf
I3308, U7505	12 25 17.9	+26 42 53	1.5	0.13	15.41	0.10	7	A	316	277	128	1.10	3.8	h
KK 144	12 25 27.9	+28 28 57	1.5	0.33	16.5	0.09	10	A	483	453	44	0.93	6.3	h
N4395	12 25 49.8	+33 32 46	13.2	0.83	10.61	0.07	9	K	320	315	109	2.52	4.61	rgb
UA 281	12 26 16.0	+48 29 31	0.8	0.75	15.14	0.06	10	A	281	349	46	1.00	5.7	bs
DDO 126, U7559	12 27 05.1	+37 08 33	3.2	0.63	14.12	0.06	10	A	218	231	59	1.45	4.87	rgb
DDO 125, U7577	12 27 41.8	+43 29 38	4.3	0.56	12.95	0.09	10	A	195	240	30	1.34	2.54	rgb
U7584, KK 148	12 28 02.9	+22 35 22	0.8	0.50	16.2	0.10	9	A	603	545	35	0.46	7.6	h
KKH 80	12 28 05.4	+22 17 27	0.9	0.67	17.0	0.09	10	A	601	542	13	−0.10	7.5	h
N4449	12 28 11.2	+44 05 40	6.2	0.71	10.06	0.08	9	BK	201	249	136	2.90	4.21	rgb
DDO 127, U7599	12 28 28.5	+37 14 00	2.0	0.50	14.98	0.07	9	A	277	291	66	1.10	6.9	bs
U7605	12 28 39.0	+35 43 05	1.1	0.73	14.76	0.06	10	A	310	317	29	0.76	4.43	rgb
N4460	12 28 45.8	+44 51 52	4.0	0.30	12.26	0.08	−1	B	490	542	121	0.32	9.59	sbf
KK 149	12 28 52.3	+42 10 40	0.8	0.63	15.01	0.11	10	A	407	446	60	0.48	6.2	h
U7639	12 29 53.0	+47 31 48	2.3	0.70	14.13	0.05	10	A	381	446	40	0.72	8.0	bs
KK 151	12 30 23.8	+42 54 05	1.2	0.42	15.8	0.08	9	A	436	479	34	0.30	6.7	h
DDO 133, U7698	12 32 53.0	+31 32 21	6.5	0.69	13.53	0.07	10	A	332	321	53	1.64	6.1	bs
N4523	12 33 48.0	+15 10 05	2.0	0.67	14.42	0.17	8	A	262	173	134	1.31	6.4	bs
KKSG 29	12 37 14.1	−10 29 51	1.4	0.43	16.54	0.12	10	A	. . .	. . .	. . .	. . .	9.3	mem
Arp 211, KK 154	12 37 21.3	+38 44 43	0.7	0.57	15.74	0.08	10	A	458	484	27	0.11	6.70	rgb
KKSG 30	12 37 35.9	−08 52 02	1.0	0.50	17.3	0.14	10	A	1105	917	36	0.00	9.3	mem
KKSG 31	12 38 33.7	−10 29 25	0.8	0.88	17.7	0.13	−3	A	. . .	. . .	. . .	. . .	9.3	mem
UA 292	12 38 40.0	+32 46 00	1.0	0.70	16.10	0.07	10	A	308	306	27	1.16	3.1	bs
KKSG 32	12 39 54.4	−11 45 03	0.6	0.83	18.5	0.21	−3	A	. . .	. . .	. . .	. . .	9.3	mem
N4594	12 39 59.1	−11 37 23	8.5	0.51	8.98	0.22	1	B	1024	828	743	1.15	9.33	sbf
N4605	12 40 00.3	+61 36 29	5.9	0.37	10.89	0.06	8	K	143	276	133	1.69	5.2	bs
KKSG 33	12 40 08.9	−12 21 53	0.4	0.99	18.5	0.20	−3	A	. . .	. . .	. . .	. . .	9.3	mem
N4600	12 40 23.0	+03 07 04	1.2	0.67	13.73	0.12	1	K	787	648	. . .	. . .	7.35	sbf
KKSG 34	12 41 18.9	−11 55 39	0.8	0.88	18.0	0.23	−3	A	. . .	. . .	. . .	. . .	9.3	mem
U7857	12 41 54.2	+13 46 22	1.3	0.62	14.70	0.10	8	A	18	−73	. . .	. . .	6.3	bs
N4627	12 41 59.7	+32 34 25	2.6	0.69	13.06	0.07	−3	A	542	541	. . .	. . .	9.38	sbf
I3687, DDO 141	12 42 15.1	+38 30 07	3.4	0.88	13.75	0.09	10	A	358	385	46	1.34	4.57	rgb
KK 160	12 43 57.4	+43 39 41	0.8	0.75	17	0.11	10	A	293	346	22	−0.15	4.8	h
KKs 51	12 44 21.5	−42 56 23	0.9	0.44	16.7	0.38	−3	A	. . .	. . .	. . .	. . .	3.6	mem
E381−018	12 44 42.7	−35 58 00	1.2	0.58	15.72	0.27	10	A	610	353	50	0.30	4.9	h
E381−20	12 46 00.4	−33 50 17	3.0	0.40	14.44	0.28	10	A	585	332	79	1.58	4.6	h
DDO 147, U7949	12 46 59.8	+36 28 35	1.9	0.79	15.12	0.08	10	A	331	351	29	1.26	9.9	bs
HIPASS J1247−77	12 47 32.6	−77 35 01	0.8	0.38	17	3.20	10	A	414	155	41	0.60	2.2	h
KK 166	12 49 13.3	+35 36 45	0.7	0.59	17.62	0.06	−3	A	. . .	. . .	. . .	. . .	4.74	rgb

TABLE 1—*Continued*

Name (1)	R.A. (J2000.0) (2)	Decl. (J2000.0) (3)	a (4)	b/a (5)	B_t (6)	A_g (7)	T (8)	T_n (9)	V_h (10)	V_{LG} (11)	W_{50} (12)	log F (13)	D (14)	Method (15)
U7990	12 50 27.0	+28 21 07	1.0	0.70	16.2	0.06	10	A	512	495	...	...	6.9	h
N4736	12 50 53.5	+41 07 10	11.2	0.81	8.74	0.08	2	B	309	353	214	1.97	4.66	rgb
DDO 154, U8024	12 54 05.2	+27 08 55	3.0	0.73	14.17	0.04	10	A	375	355	85	2.16	4.3	bs
E443−09, KK 170	12 54 53.6	−28 20 27	0.8	0.75	17.06	0.28	10	A	645	410	29	0.00	5.7	h
U8055	12 56 04.0	+03 48 41	0.9	0.78	15.9	0.14	10	A	618	493	90	0.96	6.8	h
E219−010	12 56 09.6	−50 08 38	1.4	0.79	16.42	0.96	−3	A	...	...	...	...	4.28	sbf
U8061, KK 172	12 56 43.4	+11 55 52	0.9	0.80	15.2	0.13	10	A	560	471	53	0.30	6.5	h
N4826	12 56 44.2	+21 41 05	10.3	0.52	9.12	0.18	2	B	408	364	304	1.75	7.48	sbf
GR 8, DDO 155	12 58 40.4	+14 13 03	1.1	0.91	14.61	0.11	10	A	214	136	26	0.89	2.10	rgb
UA 319	13 02 14.4	−17 14 15	1.6	0.69	15.33	0.35	9	A	747	547	110	0.50	7.6	h
DDO 161	13 03 16.8	−17 25 23	8.0	0.13	13.78	0.35	8	A	744	545	135	1.91	7.6	h
E269−37, KK 179	13 03 33.6	−46 35 03	0.8	0.75	16.26	0.57	−3	A	...	...	...	...	3.48	rgb
KK 182, Cen 6	13 05 02.9	−40 04 58	1.0	0.60	16.33	0.44	10	A	613	360	16	0.33	3.6	mem
N4945	13 05 26.1	−49 28 16	19.8	0.19	9.27	0.76	6	B	560	296	360	2.60	3.6	mem
I4182	13 05 49.3	+37 36 21	6.0	0.92	12.02	0.06	9	A	320	356	40	1.82	4.70	cep
DDO 165	13 06 26.8	+67 42 15	3.5	0.54	13.31	0.10	10	A	31	196	46	1.45	4.57	rgb
U8215	13 08 03.6	+46 49 41	1.0	0.70	16.03	0.04	10	A	218	297	26	0.63	5.6	bs
P45628	13 09 36.6	−27 08 26	0.6	0.67	15.11	0.33	10	A	693	470	55	0.83	6.5	h
E269−058	13 10 32.9	−46 59 27	3.0	0.60	13.29	0.46	10	A	402	142	70	0.54	3.6	mem
KKs 53, Cen 7	13 11 14.2	−38 54 22	0.8	0.87	17.3	0.38	−3	A	...	...	...	...	3.6	mem
N5023	13 12 11.9	+44 02 19	7.3	0.11	12.82	0.08	6	A	407	476	179	1.69	5.4	bs
KK 189	13 12 45.0	−41 49 55	0.6	0.67	17.75	0.49	−3	A	...	...	...	...	3.6	mem
E269−66, KK 190	13 13 09.2	−44 53 24	1.4	0.86	14.59	0.40	−1	A	784	528	...	...	3.54	sbf
DDO 167, U8308	13 13 22.8	+46 19 11	1.1	0.55	15.45	0.04	10	A	163	243	26	0.66	4.19	rgb
DDO 168, U8320	13 14 28.6	+45 55 10	3.6	0.39	12.97	0.07	10	A	194	273	60	1.78	4.33	rgb
DDO 169, U8331	13 15 30.7	+47 29 47	2.7	0.33	14.46	0.04	10	A	259	345	50	1.18	8.2	bs
N5068	13 18 55.3	−21 02 21	7.2	0.88	10.52	0.44	6	KB	673	473	66	2.05	6.6	h
KK 195	13 21 08.2	−31 31 47	1.3	0.46	17.1	0.27	10	A	564	338	21	0.74	5.22	rgb
KKs 54	13 21 32.4	−31 53 11	1.1	0.68	18.13	0.29	−3	A	...	...	...	...	4.6	mem
KK 196	13 21 47.1	−45 03 48	1.3	0.77	16.14	0.36	10	A	741	490	...	...	3.6	mem
N5102	13 21 57.8	−36 37 47	8.7	0.32	10.28	0.24	1	K	467	230	196	1.94	3.40	rgb
KK 197	13 22 01.8	−42 32 08	0.9	0.78	15.68	0.66	−3	A	...	...	...	...	3.6	mem
KKs 55	13 22 12.4	−42 43 51	0.8	0.88	18.5	0.63	−3	A	...	...	...	...	3.6	mem
KK 198	13 22 56.1	−33 34 03	0.6	0.83	17.65	0.30	−3	A	...	...	...	...	4.6	mem
KK 200	13 24 36.0	−30 58 20	1.3	0.62	16.67	0.30	9	A	487	264	22	0.23	4.63	rgb
N5128	13 25 28.9	−43 00 60	25.7	0.78	7.84	0.50	−2	B	547	301	527	1.94	3.66	rgb
I4247, E444−34	13 26 44.4	−30 21 45	1.3	0.38	14.4	0.27	10	A	415	195	45	0.48	4.6	mem
KK 203	13 27 28.1	−45 21 09	0.5	0.99	18	0.44	−3	A	...	...	...	...	3.6	mem
E324−24	13 27 37.4	−41 28 50	3.2	0.72	12.90	0.47	10	A	513	270	75	1.85	3.73	rgb
P170257	13 29 21.0	−21 10 45	0.6	0.75	15.3	0.46	10	A	650	457	...	...	6.3	h
N5204	13 29 36.4	+58 25 04	5.0	0.60	11.73	0.05	9	A	203	341	110	2.06	4.65	rgb
N5194	13 29 56.0	+47 14 04	11.2	0.62	8.57	0.15	5	B	463	555	195	2.34	8.0	mem
N5195	13 29 58.7	+47 16 05	5.0	0.92	10.45	0.15	−1	B	465	558	...	1.00	8.02	sbf
U8508	13 30 44.4	+54 54 36	1.7	0.59	14.12	0.06	10	A	62	186	49	1.15	2.56	rgb
N5206	13 33 43.9	−48 09 04	3.7	0.86	11.64	0.52	−3	K	571	322	...	...	3.6	mem
N5229	13 34 02.9	+47 54 55	3.6	0.13	14.33	0.08	7	A	363	460	127	1.39	5.1	bs
N5238	13 34 42.7	+51 36 50	1.7	0.82	13.55	0.05	8	A	232	345	36	0.65	5.2	bs
P47885	13 35 08.1	−30 07 04	0.5	0.80	15.8	0.21	10	K	570	360	110	1.12	5.0	h
E444−78, UA 365	13 36 30.8	−29 14 11	1.2	0.42	15.53	0.23	10	A	573	363	34	0.60	4.6	mem
KK 208	13 36 35.5	−29 34 15	6.0	0.42	14.3	0.19	−3	A	...	...	...	...	4.68	rgb
N5236	13 37 00.1	−29 52 04	12.9	0.89	8.20	0.28	5	B	516	304	212	3.25	4.47	cep
DEEP J1337−33	13 37 00.6	−33 21 47	0.2	0.95	17.3	0.21	10	A	591	371	20	0.02	4.51	rgb
E444−84	13 37 20.2	−28 02 46	1.3	0.77	15.06	0.30	10	A	587	380	60	1.32	4.61	rgb
HIPASS J1337−39	13 37 25.1	−39 53 52	0.6	0.83	16.5	0.32	10	A	492	258	37	0.82	4.90	rgb
N5237	13 37 38.9	−42 50 51	1.9	0.84	13.23	0.41	−3	A	370	131	80	0.66	3.6	mem
U8638	13 39 19.4	+24 46 33	1.2	0.67	14.44	0.06	10	A	274	273	32	0.59	2.3	bs
DDO 181, U8651	13 39 53.8	+40 44 21	2.3	0.57	14.45	0.03	10	A	202	272	39	1.06	3.01	rgb
N5253	13 39 55.8	−31 38 24	5.0	0.38	10.87	0.24	8	A	404	190	79	1.73	4.00	cep
I4316, E445−06	13 40 18.1	−28 53 40	1.6	0.63	14.56	0.24	10	A	589	382	55	0.32	4.41	rgb
N5264	13 41 37.0	−29 54 50	2.5	0.60	12.60	0.22	10	A	477	268	41	1.09	4.53	rgb
KKs 57	13 41 38.1	−42 34 55	0.8	0.87	18.1	0.39	−3	A	...	...	...	...	3.6	mem
KK 211	13 42 05.6	−45 12 18	0.8	0.87	16.32	0.48	−5	A	...	...	...	...	3.58	rgb
KK 213	13 43 35.8	−43 46 09	0.6	0.50	18.5	0.42	−3	A	...	...	...	...	3.63	rgb
E325−11	13 45 00.8	−41 51 32	2.7	0.48	13.99	0.38	10	A	540	307	66	1.60	3.40	rgb
KKs 58	13 46 00.8	−36 19 44	1.0	0.70	17.41	0.27	−3	A	...	...	...	...	3.6	mem

TABLE 1—*Continued*

Name (1)	R.A. (J2000.0) (2)	Decl. (J2000.0) (3)	a (4)	b/a (5)	B_t (6)	A_g (7)	T (8)	T_n (9)	V_h (10)	V_{LG} (11)	W_{50} (12)	log F (13)	D (14)	Method (15)
KK 217	13 46 17.2	−45 41 05	0.6	0.83	17.57	0.52	−3	A	. . .	. . .	. . .	. . .	3.84	rgb
KK 218	13 46 39.5	−29 58 45	1.1	0.45	17.60	0.26	−3	A	. . .	. . .	. . .	. . .	4.6	mem
E174−01, KKs 59	13 47 57.7	−53 21 04	3.0	0.43	14.2	2.13	10	A	680	432	59	1.68	6.0	h
CenN	13 48 09.2	−47 33 54	0.9	0.78	17.5	0.61	−3	A	. . .	. . .	. . .	. . .	3.6	mem
HIPASS J1348−37	13 48 33.9	−37 58 03	0.5	0.80	16.9	0.33	10	A	570	347	40	0.00	3.6	mem
KK 221	13 48 46.4	−46 59 49	1.5	0.67	18.	0.60	−3	A	. . .	. . .	. . .	. . .	3.98	rgb
E383−87	13 49 18.8	−36 03 41	4.5	0.80	11.03	0.31	8	A	326	108	32	1.54	1.5	h
DDO 183, U8760	13 50 51.1	+38 01 16	2.2	0.32	14.45	0.07	10	A	191	257	30	0.99	5.1	bs
HIPASS J1351−47	13 51 22.0	−47 00 00	0.5	0.60	17.5	0.62	10	A	530	292	35	0.54	3.6	mem
KKH 86	13 54 33.6	+04 14 35	0.7	0.71	16.88	0.12	10	A	287	209	14	−0.30	2.61	rgb
U8837, Ho IV	13 54 45.1	+53 54 17	4.3	0.30	13.76	0.06	10	A	144	276	77	1.30	8.3	bs
U8833	13 54 48.7	+35 50 15	0.9	0.89	15.15	0.05	10	A	226	285	28	0.78	3.19	rgb
E384−016	13 57 01.6	−35 20 02	1.3	0.69	15.11	0.32	10	A	561	350	. . .	. . .	3.72	sbf
M101	14 03 12.8	+54 21 02	28.8	0.93	8.20	0.04	6	BK	241	379	143	3.25	7.38	cep
N5408	14 03 21.5	−41 22 35	1.9	0.42	12.21	0.30	10	A	509	288	68	1.81	4.81	rgb
N5474	14 05 02.1	+53 39 47	4.8	0.90	11.60	0.05	8	AB	277	413	40	2.04	7.2	bs
N5477	14 05 33.1	+54 27 39	1.7	0.76	14.19	0.05	9	A	304	443	52	1.12	7.7	bs
KK 230	14 07 10.7	+35 03 37	0.6	0.83	17.9	0.06	10	A	62	126	21	0.41	1.90	rgb
Circinus, E97−13	14 13 09.3	−65 20 21	6.9	0.43	12.1	4.2	3	B	449	204	296	3.28	2.8	h
KKH 87	14 15 09.4	+57 05 15	0.8	0.37	16.1	0.05	10	A	320	473	38	0.53	7.4	mem
DDO 187	14 15 56.5	+23 03 19	1.7	0.76	14.38	0.10	10	A	152	172	34	1.08	2.50	rgb
N5585	14 19 48.3	+56 43 49	5.8	0.64	11.19	0.07	4	AB	305	459	146	2.10	5.7	bs
DDO 190, U9240	14 24 43.5	+44 31 33	1.8	0.89	13.10	0.05	10	A	150	263	45	1.37	2.79	rgb
P51659	14 28 03.7	−46 18 06	2.4	0.38	16.50	0.56	10	A	386	171	37	1.23	3.58	rgb
E222−10	14 35 03.0	−49 25 18	0.9	0.44	16.33	1.11	10	A	632	415	65	0.63	5.8	h
DDO 194, U9405	14 35 24.6	+57 15 24	1.7	0.35	14.52	0.05	10	A	222	385	85	0.82	8.0	bs
E272−25	14 43 25.5	−44 42 19	3.0	0.50	14.77	0.69	8	A	624	422	75	0.48	5.9	h
E223−09	15 01 08.5	−48 17 33	2.6	0.92	13.82	1.12	10	A	586	387	63	2.10	5.4	h
UMin	15 09 11.3	+67 12 52	30.2	0.63	11.9	0.14	−3	A	−247	−44	. . .	. . .	0.06	rgb
E274−01	15 14 13.5	−46 48 45	13.4	0.10	11.71	1.11	7	A	522	335	177	2.18	4.7	h
KKR 25	16 13 47.6	+54 22 16	1.1	0.59	16.45	0.04	10	A	−139	68	15	0.18	1.86	rgb
E137−18	16 20 59.3	−60 29 15	3.2	0.34	12.20	1.05	9	A	606	421	145	1.56	5.8	h
Draco	17 20 01.4	+57 54 34	35.5	0.69	10.9	0.12	−3	A	−293	−48	. . .	. . .	0.08	rgb
Milky Way	17 45 09.2	−28 01 12	. . .	. . .	−4.60	. . .	4	B	0	−88	440	8.30	0.01	cep
I4662, E102−14	17 47 06.3	−64 38 25	2.8	0.57	11.74	0.30	9	A	308	145	100	2.10	2.0	h
N6503	17 49 27.6	+70 08 41	7.1	0.34	11.05	0.14	6	B	43	301	173	2.20	5.27	rgb
Sag dSph	18 55 03.1	−30 28 42	450.	0.48	4.50	0.66	−3	A	142	161	. . .	. . .	0.02	rgb
N6789, KK 245	19 16 41.8	+63 58 22	1.4	0.77	13.76	0.30	10	A	−141	144	. . .	. . .	3.60	rgb
Sag DIG, E594−4	19 29 59.0	−17 40 41	2.9	0.72	14.12	0.52	10	A	−77	23	32	1.40	1.04	rgb
N6822	19 44 57.7	−14 48 11	15.5	0.87	9.32	1.03	10	A	−57	64	63	3.38	0.50	cep
KK 246	20 03 57.4	−31 40 54	1.2	0.42	17.06	1.28	10	A	358	401	53	0.92	5.6	h
U11583, KK 250	20 30 15.3	+60 26 25	1.8	0.44	15.90	1.32	10	A	127	430	90	1.30	5.9	bs
KK 251	20 30 32.0	+60 21 13	1.6	0.50	16.5	1.28	10	A	130	433	49	1.12	5.9	bs
KK 252	20 31 32.6	+60 48 48	0.9	0.99	16.7	1.95	10	A	138	441	28	0.15	5.9	bs
N6946	20 34 51.1	+60 09 15	11.5	0.85	9.58	1.48	6	KB	51	355	169	2.95	5.9	bs
KKR 55	20 45 20.8	+60 24 40	0.6	0.67	17.0	2.94	10	A	32	337	37	0.73	5.9	bs
DDO 210	20 46 51.8	−12 50 53	2.2	0.50	14.0	0.22	10	A	−137	13	21	1.05	0.94	rgb
KKR 56	20 48 24.1	+58 37 06	0.7	0.64	17.6	3.14	10	A	−47	260	. . .	0.75	5.9	bs
Cepheus 1	20 51 09.7	+56 53 24	3.0	0.50	15.4	4.05	9	A	58	367	90	2.13	5.9	mem
I5052, E074−15	20 52 06.2	−69 12 14	5.9	0.14	11.68	0.22	7	A	598	454	182	2.02	6.3	h
KKR 59	21 03 24.2	+57 17 14	2.3	0.61	15.7	3.86	10	A	1	311	63	1.56	5.9	bs
KKR 60	21 05 53.0	+57 12 19	0.7	0.70	18	4.57	10	A	−35	275	. . .	. . .	5.9	mem
I5152, E237−27	22 02 41.9	−51 17 43	5.2	0.62	11.06	0.11	10	A	124	75	91	1.99	2.07	rgb
KK 258	22 40 43.9	−30 47 59	1.6	0.50	17.36	0.06	−3	A	. . .	. . .	. . .	. . .	3.9	mem
Tucana, P69519	22 41 49.0	−64 25 12	2.9	0.41	15.7	0.14	−2	A	130:	9	. . .	. . .	0.88	rgb
UA 438	23 26 27.5	−32 23 26	1.5	0.80	13.86	0.06	10	A	62	99	34	1.18	2.23	rgb
Cas dSph, KKH 96	23 26 31.8	+50 40 32	2.5	0.80	13.65	0.83	−3	A	−307	−5	. . .	. . .	0.79	rgb
Pegasus, D216	23 28 34.1	+14 44 48	5.0	0.54	13.21	0.28	10	A	−184	60	23	1.28	0.76	rgb
UA 442	23 43 46.0	−31 57 33	6.4	0.14	13.58	0.07	8	A	267	299	93	1.52	4.27	rgb
KKH 98	23 45 34.0	+38 43 04	1.1	0.55	16.7	0.53	10	A	−137	151	22	0.61	2.45	rgb
Peg dSph, KKH 99	23 51 46.4	+24 35 10	4.0	0.50	14.05	0.28	−3	A	−354	−94	. . .	. . .	0.82	cep
E149−003	23 52 02.8	−52 34 39	2.2	0.18	15.0	0.06	9	A	577	501	55	0.77	6.4	tf
N7793	23 57 49.4	−32 35 24	9.3	0.68	9.70	0.08	7	K	229	252	174	2.47	3.91	rgb

TABLE 2
NOTES TO TABLE 1

R.A. (J2000.0)	Decl. (J2000.0)	Notes
00 01 58.1......	−15 27 40	Dolphin 2000
00 08 13.3......	−34 34 42	Karachentsev et al. 2003a
00 15 08.5......	−39 13 13	Karachentsev et al. 2003a
00 15 25.1......	−21 26 38	Jerjen, Freeman, & Binggeli 1998, and +0.9 Mpc correction; photometry(ph): Jerjen, Binggeli, & Freeman 2000a
00 15 31.4......	−32 10 48	Karachentsev et al. 2000b; ph: Parodi et al. 2002
00 20 24.5......	+59 17 30	Ferrarese et al. 2000
00 23 51.7......	−24 42 18	Karachentsev et al. 2003a; ph: Jerjen et al. 2000a
00 26 11.0......	−11 02 40	Sarajedini et al. 2002
00 26 33.3......	−41 51 20	Karachentsev et al. 2002c; ph: Jerjen et al. 2000a
00 29 04.0......	+43 25 54	Georgiev, Karachentsev, & Tikhonov 1997; ph: Barazza, Binggeli, & Prugniel 2001
00 33 11.6......	+48 30 28	Han et al. 1997
00 35 33.8......	+36 29 52	Armandroff et al. 1993
00 38 58.0......	+48 20 10	Lee, Freedman, & Madore 1993
00 40 22.5......	+41 41 11	Lee 1996
00 42 30.1......	+40 34 33	Ferguson, Gallagher, & Wyse 2000
00 42 42.1......	+40 51 59	Evans et al. 2000
00 42 44.5......	+41 16 09	Ferrarese et al. 2000
00 43 03.8......	−22 15 01	Karachentsev et al. 2003a
00 45 40.0......	+38 02 14	Da Costa et al. 1996
00 47 08.3......	−20 45 36	Karachentsev et al. 2003a
00 47 34.3......	−25 17 32	Karachentsev et al. 2003a
00 49 21.1......	−18 04 28	Karachentsev et al. 2003a; ph: Jerjen et al. 2000a
00 49 49.3......	−21 00 58	Karachentsev et al. 2003a
00 50 24.6......	−19 54 25	Karachentsev et al. 2003a; ph: Jerjen et al. 2000a
00 52 38.0......	−72 48 01	Ferrarese et al. 2000
00 54 53.5......	−37 40 57	Freedman et al. 1992
01 00 09.4......	−33 42 33	van den Bergh 2000
01 03 56.6......	+21 53 41	Miller et al. 2001; ph: Karachentseva et al. 1996
01 04 54.1......	+02 07 60	Dolphin et al. 2001a
01 07 22.3......	+16 41 02	Maiz−Apellaniz, Cieza, & MacKenty 2002; ph: Barazza et al. 2001
01 07 32.5......	+51 26 25	Karachentsev et al. 2003b
01 09 26.9......	+35 43 03	Karachentsev et al. 2002c
01 10 17.1......	+47 37 41	Armandroff et al. 1998
01 16 29.8......	+33 25 09	Da Costa et al. 2000
01 32 42.3......	+18 18 57	Sharina, Karachentsev, & Tikhonov 1996
01 33 50.8......	+30 39 37	Ferrarese et al. 2000
01 34 51.6......	+52 05 30	Karachentsev et al. 2001a
01 35 05.0......	−41 26 11	M. E. Sharina 2004, private communication
01 36 41.3......	+15 47 12	Sharina et al. 1996
01 39 41.6......	+15 54 11	Sharina et al. 1996
01 40 09.6......	+15 54 17	Sharina et al. 1996
01 42 16.8......	+26 22 04	N672 group
01 43 41.4......	+15 41 23	Sharina et al. 1996
01 44 42.7......	+27 17 16	N672 group
01 45 03.6......	−43 35 53	Karachentsev et al. 2003a
01 46 41.6......	+26 48 05	N672 group
01 47 30.1......	+27 19 52	I. O. Drozdovsky 1999, private communication
01 47 53.2......	+27 26 01	I. O. Drozdovsky 1999, private communication
01 49 32.3......	+32 35 33	Makarova & Karachentsev 1998; ph: Barazza et al. 2001
01 51 06.3......	−44 26 41	Held, Saviane, & Momany 1999. New $V_h = -13$ km s^{-1} from Irwin & Tolstoy 2002
01 55 20.6......	+27 57 15	Karachentsev et al. 2003b
02 00 09.9......	+28 49 57	Karachentsev et al. 2003b
02 01 16.8......	+28 50 37	Drozdovsky & Karachentsev 2000
02 06 07.9......	+69 00 36	IC 342 group; ph: Buta & McCall 1999 (calculated from V, I)
02 14 03.7......	+27 52 38	Tonry et al. 2001
02 22 32.8......	+42 20 48	Ferrarese et al. 2000
02 24 35.0......	+56 00 42	Maffei group
02 24 42.7......	−73 30 46	Karachentseva & Karachentsev 2000
02 27 16.0......	+33 34 41	Ferrarese et al. 2000
02 27 27.0......	+57 29 16	Maffei group
02 35 35.6......	+59 22 47	Maffei group
02 36 35.5......	+59 39 18	Fingerhut et al. 2003
02 37 45.0......	−61 20 28	Karachentsev et al. 2003b; ph: Parodi et al. 2002
02 39 54.7......	−34 31 33	van den Bergh 2000

TABLE 2—*Continued*

R.A. (J2000.0)	Decl. (J2000.0)	Notes
02 41 54.5......	+59 36 11	Karachentsev et al. 2003c
02 54 08.5......	+59 00 19	Maffei group
02 55 43.6......	+58 51 42	Maffei group
02 56 52.1......	−54 34 23	Parodi et al. 2002
02 56 56.1......	+58 54 42	Karachentsev et al. 2003c
02 59 42.4......	+25 14 15	Karachentsev, Musella, & Grimaldi 1996; ph: Barazza et al. 2001
03 03 05.9......	+33 41 40	Karachentsev et al. 2003b
03 18 15.4......	−66 29 51	Mendez et al. 2002
03 20 07.4......	−52 11 06	Parodi et al. 2002
03 20 23.7......	+17 17 42	Sharina, Karachentsev, & Tikhonov 1999
03 21 05.7......	−66 19 22	Karachentsev et al. 2003b
03 24 07.2......	+17 45 15	Parodi et al. 2002; ph: Barazza et al. 2001
03 32 07.1......	+47 47 37	Huchtmeier et al. 2000
03 33 11.8......	−50 24 38	Parodi et al. 2002
03 45 12.6......	+67 51 50	Karachentsev et al. 2003c
03 46 49.7......	+68 05 45	Saha, Claver, & Hoessel 2002
03 57 00.6......	+16 31 28	Georgiev et al. 1997; ph: Barazza et al. 2001
03 59 49.5......	+67 07 31	Karachentsev et al. 1997
04 08 54.1......	−55 59 32	In a pair. *W* from HIPASS; ph: Parodi et al. 2002
04 25 15.6......	+72 48 21	Karachentsev et al. 2003c; ph: Karachentseva et al. 1996
04 30 49.1......	+64 50 53	Makarova & Karachentsev 2003
04 32 00.3......	+63 36 50	Karachentsev et al. 1997; ph: Karachentseva et al. 1996
04 32 49.9......	+71 52 52	Karachentsev et al. 2003c
04 53 06.9......	+67 05 57	Karachentsev et al. 2003c
04 54 13.7......	−53 21 41	Tosi et al. 2001
05 14 15.1......	+62 34 51	Karachentsev et al. 2002c
05 23 34.6......	−69 45 22	van den Bergh 2000
05 24 59.5......	+04 30 18	Makarova & Karachentsev 1998; ph: Barazza et al. 2001
05 41 41.5......	+06 40 54	Huchtmeier et al. 2000
05 45 02.0......	+05 04 06	Karachentsev & Musella 1996; ph: Barazza et al. 2001
05 57 36.7......	+07 29 31	Karachentsev & Musella 1996
05 59 41.2......	+73 25 39	Karachentsev et al. 2003b
06 05 45.4......	−33 04 54	Huchtmeier et al. 2001
06 15 54.5......	−57 43 35	Huchtmeier et al. 2001
06 26 17.0......	−26 15 56	Karachentsev et al. 2003b; ph: Parodi et al. 2002
06 30 29.2......	+33 18 07	Makarova & Karachentsev 1998; ph: Barazza et al. 2001
06 37 56.6......	−25 59 59	Karachentsev et al. 2003b; ph: Parodi et al. 2002
06 39 32.9......	−40 43 13	Huchtmeier et al. 2000; ph: Parodi et al. 2002
06 41 36.7......	−50 57 58	van den Bergh 2000
06 46 56.9......	−17 56 29	Probable H I confusion from N2283 at 25′ southwest.
06 47 45.8......	+80 07 26	Should be confirmed as the LV member.
06 55 40.0......	+39 05 42	Makarova & Karachentsev 1998; ph: Barazza et al. 2001
07 00 29.3......	−04 12 30	Massey, Henning, & Kraan-Korteweg 2003
07 05 17.1......	−58 31 14	Karachentsev et al. 2003b; ph: Parodi et al. 2002
07 06 56.8......	−22 02 26	Huchtmeier et al. 2000; ph: Parodi et al. 2002
07 09 18.8......	+44 22 48	Makarova & Karachentsev 1998; ph: Karachentseva et al. 1996
07 10 13.6......	+44 27 25	Makarova & Karachentsev 1998; ph: Barazza et al. 2001
07 13 51.8......	+10 31 19	Karachentsev et al. 2003b; ph: Makarova 1999
07 22 44.5......	+45 06 30	Makarova & Karachentsev 1998 (ph: calculated from *V*, *I*)
07 28 17.2......	+40 46 13	M. E. Sharina 2004, private communication; ph: Makarova & Karachentsev 1998 (calculated from *V*, *I*)
07 28 52.0......	+69 12 19	Karachentsev et al. 2003b; ph: Makarova 1999
07 31 19.3......	−68 11 10	Putman et al. 2002; ph: Parodi et al. 2002
07 34 11.3......	+66 53 10	Karachentsev et al. 1999c
07 36 54.4......	+65 35 58	Freedman & Madore 1988
07 41 25.8......	+40 06 45	Georgiev et al. 1997
07 41 55.0......	+16 48 02	Karachentsev et al. 2003b; ph: Makarova 1999
07 42 31.2......	+16 33 40	Karachentsev et al. 2003b; ph: Makarova et al. 2002
07 57 01.8......	+14 23 27	Karachentsev et al. 2003b; ph: Barazza et al. 2001
08 13 14.7......	+45 59 26	Sharina et al. 1999; ph: Barazza et al. 2001
08 19 05.9......	+70 42 51	Karachentsev et al. 2003b
08 19 23.3......	−85 08 44	Huchtmeier et al. 2000; ph: Parodi et al. 2002
08 23 56.0......	+71 01 46	Karachentsev et al. 2003b; ph: Makarova 1999
08 28 28.5......	+41 51 24	Makarova & Karachentsev 1998
08 34 06.5......	+66 10 45	Karachentsev et al. 2003b; ph: Makarova 1999
08 37 03.0......	+69 46 31	Dolphin et al. 2001b; ph: Makarova 1999

TABLE 2—*Continued*

R.A. (J2000.0)	Decl. (J2000.0)	Notes
08 52 40.9......	+33 25 02	Tonry et al. 2001
08 52 50.7......	+33 47 52	N2683 group
08 55 22.0......	+33 33 33	N2683 group
09 02 54.0......	+20 04 31	Schombert, Pildis, & Eder 1997; ph: Makarova 1999
09 08 36.6......	+05 17 32	Karachentsev et al. 2001a; ph: Makarova et al. 2002
09 08 53.5......	+14 34 55	V_h and W from Huchtmeier et al. 2003
09 09 10.9......	−23 22 04	N2784 group
09 09 46.6......	−23 00 33	N2784 group; ph: Parodi et al. 2002
09 11 42.3......	−23 58 56	N2784 group
09 12 18.5......	−24 12 40	N2784 group. In the KK(1998) list it is KK 73, but its coordinates were confused with PGC 16 60 99 (see also Parodi et al. 2002)
09 12 19.4......	−24 10 18	Tonry et al. 2001
09 12 29.3......	−24 14 28	N2784 group; ph: Parodi et al. 2002
09 13 12.1......	−19 24 31	Parodi et al. 2002
09 19 18.6......	+69 12 12	Tonry et al. 2001
09 19 29.4......	+21 36 12	Schombert et al. 1997; ph: Makarova & Karachentsev 1998
09 23 09.9......	−20 10 03	Huchtmeier et al. 2001
09 26 11.5......	−76 37 35	Karachentsev et al. 2003b; ph: Parodi et al. 2002
09 30 12.9......	+19 59 30	N2903 group; ph: Makarova et al. 2002
09 32 09.6......	+21 30 02	Drozdovsky & Karachentsev 2000
09 32 48.9......	+21 27 54	N2903 group; ph: Makarova & Karachentsev 1998 (calculated from V, I)
09 40 28.2......	+71 11 11	Karachentsev et al. 2003b; ph: Makarova 1999
09 42 59.8......	+33 15 52	Schneider et al. 1990
09 44 50.0......	+67 28 32	Caldwell et al. 1998
09 45 04.2......	+32 14 18	Karachentsev et al. 2001a; ph: Makarova et al. 2002
09 45 10.0......	+68 45 54	Karachentsev et al. 2001b; ph: Makarova et al. 2002
09 47 15.6......	+67 54 49	Karachentsev et al. 2003b; ph: Bremnes et al. 1998
09 50 10.0......	+67 30 24	Karachentsev et al. 2000a
09 50 19.6......	+31 27 24	H I confusion with U5272; ph: Makarova et al. 2002
09 50 22.4......	+31 29 15	Makarova & Karachentsev 1998; ph: Barazza et al. 2001
09 51 17.2......	+07 49 38	Sharina et al. 1999; ph: Makarova 1999
09 53 48.5......	+68 58 09	Karachentsev et al. 2003b; ph: Makarova 1999
09 55 10.5......	−06 16 12	N3115 group. $V_h = 1021 \pm 89$ km s^{-1} in the NED.
09 55 33.5......	+69 03 60	Freedman et al. 1994
09 55 53.9......	+69 40 57	Sakai & Madore 1999
09 56 45.7......	+28 49 30	Makarova & Karachentsev 1998
09 57 02.7......	+68 35 30	Karachentsev et al. 2000a; ph: Bremnes et al. 1998
09 57 29.0......	+69 16 20	Karachentsev et al. 2003b; ph: Makarova et al. 2002
09 57 32.4......	+69 02 35	M81 group; ph: Makarova 1999
09 59 26.4......	+30 44 47	Tolstoy et al. 1998
09 59 46.5......	−09 21 04	N3115 group
10 00 00.1......	+05 19 56	Karachentsev et al. 2002c
10 00 16.0......	+63 11 06	Karachentsev et al. 2001b; ph: Makarova et al. 2002
10 01 38.4......	−08 14 56	N3115 group
10 02 36.2......	−06 00 43	N3115 group; ph: Parodi et al. 2002
10 03 07.2......	−26 09 36	Karachentsev et al. 2002c
10 03 21.0......	+68 44 02	Karachentsev et al. 2003b
10 04 04.0......	−27 19 55	Aparicio et al. 1997
10 04 41.0......	+29 21 59	Makarova & Karachentsev 1998
10 04 40.3......	+68 15 20	Caldwell et al. 1998; ph: Bremnes et al. 1998
10 05 07.3......	+66 33 18	Karachentsev et al. 2000a; ph: Bremnes et al. 1998
10 05 14.0......	−07 43 07	Tonry et al. 2001
10 05 30.6......	+70 21 52	Sharina et al. 1999; ph: Makarova 1999
10 05 34.4......	−07 44 57	N3115 group; ph: Parodi et al. 2002
10 05 41.6......	−07 58 53	N3115 group; ph: Parodi et al. 2002
10 07 01.9......	+67 49 39	Karachentsev et al. 2000a; ph: Bremnes et al. 1998
10 07 19.7......	+10 21 44	Maiz−Apellaniz et al. 2002; ph: Makarova & Karachentsev 1998
10 08 05.9......	+68 23 57	M81 group
10 08 26.9......	+12 18 29	van den Bergh 2000
10 11 00.8......	−04 41 34	Dolphin et al. 2003
10 13 03.0......	−01 36 52	van den Bergh 2000
10 15 59.4......	+06 48 17	Should be confirmed as the LV member.
10 21 00.2......	+68 41 60	Boyce et al. 2001. H I cloud?
10 21 25.2......	+71 06 58	M81 group
10 24 28.3......	−12 25 57	Huchtmeier et al. 2001; ph: Makarova et al. 2002

TABLE 2—*Continued*

R.A. (J2000.0)	Decl. (J2000.0)	Notes
10 26 27.9......	+67 39 24	Karachentsev et al. 2000a. V_h from Sharina, Karachentsev, & Burenkov 2001.
10 28 20.9......	+22 34 17	Sharina et al. 1999
10 28 22.4......	+68 24 58	Karachentsev et al. 2003b
10 30 35.0......	+70 37 10	Karachentsev et al. 2003b; ph: Makarova 1999
10 31 53.8......	+12 55 35	The "LeG" means Leo group. Karachentsev & Karachentseva 2004
10 32 17.1......	+27 40 07	Makarova & Karachentsev 1998
10 33 19.4......	+10 11 21	Leo group. Karachentsev & Karachentseva 2004
10 34 31.9......	+66 00 42	Karachentsev et al. 2000a
10 35 48.8......	+08 28 47	Leo group. Karachentsev & Karachentseva 2004
10 36 23.8......	+12 42 27	Leo group
10 39 40.2......	+12 44 06	Leo group. Karachentsev & Karachentseva 2004
10 39 43.3......	+12 38 04	Leo group. The "FS" means the object from Ferguson & Sandage 1990.
10 39 55.7......	+13 54 28	Leo group. Karachentsev & Karachentseva 2004
10 40 56.6......	+12 28 19	Leo group. Karachentsev & Karachentseva 2004
10 42 00.3......	+12 20 06	Leo group. Ferguson & Sandage 1990
10 42 34.6......	+12 09 02	Leo group. Karachentsev & Karachentseva 2004
10 43 55.4......	+12 08 07	Leo group. Karachentsev & Karachentseva 2004
10 43 57.7......	+11 42 13	Ferrarese et al. 2000
10 44 02.1......	+15 35 20	Leo group. Karachentsev & Karachentseva 2004
10 44 07.8......	+11 31 59	Leo group. Karachentsev & Karachentseva 2004
10 44 57.6......	+11 54 58	Leo group. Ferguson & Sandage 1990
10 46 14.4......	+12 57 36	Leo group. Ferguson & Sandage 1990. V_h from Flint et al. 2001.
10 46 24.8......	+14 01 30	Leo group. Ferguson & Sandage 1990
10 46 30.0......	+11 45 21	Leo group. Ferguson & Sandage 1990
10 46 41.5......	+12 19 35	Leo group. Ferguson & Sandage 1990
10 46 45.8......	+11 49 11	Ferrarese et al. 2000
10 46 53.3......	+12 44 40	Leo group. V_h from Schneider 1989
10 46 54.8......	+12 47 17	Leo group. Ferguson & Sandage 1990
10 46 57.3......	+12 59 54	Leo group. Ferguson & Sandage 1990. V_h from Huchtmeier et al. 2003.
10 47 00.8......	+12 57 35	Leo group. Karachentsev & Karachentseva 2004
10 47 22.3......	+14 04 13	Leo group. Makarova & Karachentsev 1998
10 47 27.4......	+13 53 23	Leo group. Ferguson & Sandage 1990
10 47 41.9......	−38 51 15	V_h and W from HIPASS, V_h = 17 km s^{-1} in the NED belongs to a star.
10 47 42.4......	+13 59 08	Ferrarese et al. 2000
10 47 49.6......	+12 34 54	Ferrarese et al. 2000
10 48 16.9......	+12 37 45	Ferrarese et al. 2000
10 48 43.4......	+12 18 56	Leo group
10 48 53.7......	+14 07 28	Leo group
10 49 36.5......	+65 31 50	Sharina et al. 1999; ph: Makarova 1999
10 49 37.1......	+11 21 06	Leo group. Ferguson & Sandage 1990
10 49 52.2......	+13 09 42	V_h, W from Rosenberg & Schneider 2000
10 50 09.1......	+13 29 01	Leo group. Karachentsev & Karachentseva 2004
10 50 19.1......	+13 16 19	Leo group
10 50 27.1......	+12 21 39	Undetected in H I by Huchtmeier et al. 2001
10 50 53.3......	+13 24 44	Leo group
10 51 21.1......	+12 50 57	Leo group. Karachentsev & Karachentseva 2004
10 52 20.1......	+14 42 26	Leo group. Karachentsev & Karachentseva 2004
10 52 55.3......	+69 32 45	Karachentsev et al. 2003b; ph: Makarova 1999
10 53 00.7......	+10 22 45	Leo group. Karachentsev & Karachentseva 2004
10 55 03.6......	+14 05 36	Leo group. Schombert et al. 1997
10 55 57.1......	+12 20 20	Leo group. Schombert et al. 1997
10 56 14.0......	+12 00 37	Leo group. Schombert et al. 1997
10 57 30.2......	−48 10 44	Huchtmeier et al. 2001
10 58 10.5......	+11 59 57	Schombert et al. 1997. Undetected in H I by Huchtmeier et al. 2001
10 59 17.3......	+15 05 07	Leo group. Karachentsev & Karachentseva 2004
11 00 18.6......	+13 54 04	Leo group
11 00 45.2......	+14 10 20	Leo group. Karachentsev & Karachentseva 2004
11 00 52.1......	+13 52 51	Leo group. Schombert et al. 1997
11 03 02.1......	+08 02 54	Leo group. Karachentsev & Karachentseva 2004
11 04 26.5......	+11 45 18	Leo group. Karachentsev & Karachentseva 2004
11 05 35.0......	−01 51 49	Huchtmeier et al. 2001
11 13 29.2......	+22 09 17	van den Bergh 2000
11 18 16.1......	−32 48 42	Ferrarese et al. 2000
11 20 15.0......	+12 59 29	Ferrarese et al. 2000
11 28 00.6......	+78 59 29	Mendez et al. 2002

TABLE 2—*Continued*

R.A. (J2000.0)	Decl. (J2000.0)	Notes
11 33 29.1......	+49 14 17	Karachentsev et al. 2003b; ph: Bremnes et al. 2000
11 35 48.6......	+54 31 22	Karachentsev et al. 2003b; ph: Bremnes et al. 2000
11 36 06.4......	+45 17 07	Karachentsev et al. 2003b; ph: Bremnes et al. 2000
11 37 53.4......	−39 13 14	Huchtmeier et al. 2001
11 47 11.2......	+43 40 19	Karachentsev et al. 2003b
11 50 53.0......	+38 52 50	Karachentsev et al. 2002c; ph: Makarova 1999
11 54 43.0......	−33 33 29	Karachentsev et al. 2002b
11 58 44.1......	+27 35 06	Looks like a distant galaxy.
12 04 02.4......	+52 35 19	Makarova, Karachentsev, & Georgiev 1997; ph: Makarova 1999
12 09 56.4......	+36 26 07	Looks like a distant galaxy.
12 09 59.3......	+46 27 26	Karachentsev & Drozdovsky 1998; ph: Makarova 1999
12 12 08.9......	+36 10 10	Tikhonov & Karachentsev 1998; ph: Makarova 1999
12 13 22.7......	+29 55 18	Probably, a distant galaxy; ph: Makarova 1999
12 13 44.6......	+36 37 60	Tikhonov & Karachentsev 1998
12 13 49.6......	−38 13 53	Karachentsev et al. 2002b
12 14 07.4......	+66 05 32	NGC 4236 group
12 14 57.9......	+36 13 08	Karachentsev et al. 2002c; H I conf. w. N4214; ph: Bremnes et al. 2000
12 15 38.9......	+36 19 39	Maiz-Apellaniz et al. 2002
12 15 46.7......	+52 23 15	Looks like a distant galaxy.
12 16 28.6......	+52 13 38	Karachentsev et al. 2003b; ph: Makarova 1999
12 16 43.3......	+69 27 56	Karachentsev et al. 2002a
12 17 29.9......	+37 48 27	Karachentsev et al. 2003b
12 18 46.1......	−79 43 34	Karachentsev et al. 2002c
12 18 57.5......	+47 18 14	Newman et al. 2001
12 19 09.1......	+47 05 23	Jerjen et al. 2001. H I conf. w. NGC 4258; ph: Makarova et al. 1998
12 25 17.9......	+26 42 53	Peripheric Virgo Cluster member? ph: Makarova 1999
12 25 27.9......	+28 28 57	Karachentseva & Karachentsev 1998
12 25 49.8......	+33 32 46	Karachentsev et al. 2003b
12 26 16.0......	+48 29 31	Makarova et al. 1997
12 27 05.1......	+37 08 33	Karachentsev et al. 2003b; ph: Makarova et al. 1998
12 27 41.8......	+43 29 38	Karachentsev et al. 2002c; ph: Makarova 1999
12 28 02.9......	+22 35 22	Huchtmeier et al. 2000
12 28 05.4......	+22 17 27	Karachentsev et al. 2001a
12 28 11.2......	+44 05 40	Karachentsev et al. 2003b; ph: Makarova 1999
12 28 28.5......	+37 14 00	Makarova et al. 1998
12 28 39.0......	+35 43 05	Karachentsev et al. 2003b; ph: Makarova et al. 1998
12 28 45.8......	+44 51 52	Tonry et al. 2001
12 28 52.3......	+42 10 40	Huchtmeier et al. 2000
12 29 53.0......	+47 31 48	Makarova et al. 1998
12 30 23.8......	+42 54 05	Huchtmeier et al. 2000
12 32 53.0......	+31 32 21	Makarova et al. 1998; ph: Makarova 1999
12 33 48.0......	+15 10 05	Tikhonov, Galazutdinova, & Drozdovskii 2000. In Virgo Cluster?
12 37 14.1......	−10 29 51	N4594 group. Irr or dSph? (unseen in HIPASS)
12 37 21.3......	+38 44 43	Crone et al. 2002; ph: Makarova et al. 1998
12 37 35.9......	−08 52 02	N4594 group
12 38 33.7......	−10 29 25	N4594 group
12 38 40.0......	+32 46 00	Makarova et al. 1998
12 39 54.4......	−11 45 03	N4594 group
12 39 59.1......	−11 37 23	Ferrarese et al. 2000
12 40 00.3......	+61 36 29	Karachentsev, Kopylov, & Kopylova 1994
12 40 08.9......	−12 21 53	N4594 group
12 40 23.0......	+03 07 04	Tonry et al. 2001
12 41 18.9......	−11 55 39	N4594 group
12 41 54.2......	+13 46 22	Tikhonov et al. 2000. In Virgo Cluster?
12 41 59.7......	+32 34 25	Tonry et al. 2001
12 42 15.1......	+38 30 07	Karachentsev et al. 2003b; ph: Makarova et al. 1998
12 43 57.4......	+43 39 41	Karachentseva & Karachentsev 1998
12 44 21.5......	−42 56 23	Cen A group
12 44 42.7......	−35 58 00	Karachentsev et al. 2002b; ph: Jerjen et al. 2000a
12 46 00.4......	−33 50 17	Karachentsev et al. 2002b
12 46 59.8......	+36 28 35	Makarova et al. 1998
12 47 32.6......	−77 35 01	Kilborn et al. 2002
12 49 13.3......	+35 36 45	Karachentsev et al. 2003b
12 50 53.5......	+41 07 10	Karachentsev et al. 2003b
12 54 05.2......	+27 08 55	Makarova et al. 1998

TABLE 2—*Continued*

R.A. (J2000.0)	Decl. (J2000.0)	Notes
12 54 53.6......	−28 20 27	Huchtmeier et al. 2000
12 56 09.6......	−50 08 38	Jerjen, Freeman, & Binggeli 2000b, and −0.5 Mpc correction; ph: Jerjen et al. 2000a
12 56 43.4......	+11 55 52	Huchtmeier et al. 2000
12 56 44.2......	+21 41 05	Tonry et al. 2001
12 58 40.4......	+14 13 03	Dohm-Palmer et al. 1998; ph: Gallagher & Hunter 1986
13 02 14.4......	−17 14 15	Côté et al. 1997
13 03 16.8......	−17 25 23	Côté et al. 1997
13 03 33.6......	−46 35 03	Karachentsev et al. 2002b
13 05 02.9......	−40 04 58	Cen A group
13 05 26.1......	−49 28 16	Cen A group
13 05 49.3......	+37 36 21	Sandage & Tammann 1982; ph: Makarova 1999
13 06 26.8......	+67 42 15	Karachentsev et al. 2003b; ph: Makarova 1999
13 08 03.6......	+46 49 41	Makarova et al. 1997; ph: Bremnes et al. 1999
13 09 36.6......	−27 08 26	Da Costa et al. 1998
13 10 32.9......	−46 59 27	Cen A group. V_h from HIPASS, in the NED V_h = 1853 km s^{-1}.
13 11 14.2......	−38 54 22	Cen A group
13 12 11.9......	+44 02 19	Sharina et al. 1999
13 12 45.0......	−41 49 55	Cen A group; ph: Jerjen et al. 2000a
13 13 09.2......	−44 53 24	Jerjen et al. 2000b, and −0.5 Mpc correction; ph: Jerjen et al. 2000a
13 13 22.8......	+46 19 11	Karachentsev et al. 2003b; ph: Makarova 1999
13 14 28.6......	+45 55 10	Karachentsev et al. 2003b; ph: Makarova 1999
13 15 30.7......	+47 29 47	Karachentsev & Drozdovsky 1998; ph: Makarova 1999
13 21 08.2......	−31 31 47	Pritzl et al. 2003
13 21 32.4......	−31 53 11	M83 group; ph: Jerjen et al. 2000a
13 21 47.1......	−45 03 48	Cen A group; ph: Jerjen et al. 2000a
13 21 57.8......	−36 37 47	Karachentsev et al. 2002b
13 22 01.8......	−42 32 08	Cen A group; ph: Jerjen et al. 2000a
13 22 12.4......	−42 43 51	Cen A group
13 22 56.1......	−33 34 03	Karachentsev et al. 2002b; ph: Jerjen et al. 2000a
13 24 36.0......	−30 58 20	Karachentsev et al. 2002b
13 25 28.9......	−43 00 60	Soria et al. 1996
13 26 44.4......	−30 21 45	M83 group. V_h from HIPASS, in the NED V_h = 274 ± 65 km s^{-1}.
13 27 28.1......	−45 21 09	Cen A group
13 27 37.4......	−41 28 50	Karachentsev et al. 2002b
13 29 21.0......	−21 10 45	Da Costa et al. 1998
13 29 36.4......	+58 25 04	Karachentsev et al. 2003b
13 29 56.0......	+47 14 04	Pair with N5195
13 29 58.7......	+47 16 05	Tonry et al. 2001
13 30 44.4......	+54 54 36	Karachentsev et al. 2002c; ph: Makarova 1999
13 33 43.9......	−48 09 04	Cen A group
13 34 02.9......	+47 54 55	Sharina et al. 1999; ph: Makarova 1999
13 34 42.7......	+51 36 50	Karachentsev et al. 1994; ph: Bremnes et al. 1999
13 35 08.1......	−30 07 04	V_h and W from HIPASS. Probable H I confusion from N5236 at 30′.
13 36 30.8......	−29 14 11	M83 group
13 36 35.5......	−29 34 15	Karachentsev et al. 2002b. See also Pohlen et al. 2003
13 37 00.1......	−29 52 04	Thim et al. 2003
13 37 00.6......	−33 21 47	Grossi, Disney, & Minchin 2003
13 37 20.2......	−28 02 46	Karachentsev et al. 2002b
13 37 25.1......	−39 53 52	Grossi et al. 2003
13 37 38.9......	−42 50 51	Cen A group
13 39 19.4......	+24 46 33	Makarova et al. 1998. D = 2.3 Mpc is a lower distance limit.
13 39 53.8......	+40 44 21	Karachentsev et al. 2002c; ph: Makarova et al. 1998
13 39 55.8......	−31 38 24	Thim et al. 2003
13 40 18.1......	−28 53 40	Karachentsev et al. 2002b
13 41 37.0......	−29 54 50	Karachentsev et al. 2002b
13 41 38.1......	−42 34 55	Cen A group
13 42 05.6......	−45 12 18	Karachentsev et al. 2002b; ph: Jerjen et al. 2000a
13 43 35.8......	−43 46 09	Karachentsev et al. 2002b
13 45 00.8......	−41 51 32	Karachentsev et al. 2002b
13 46 00.8......	−36 19 44	Cen A group; ph: Jerjen et al. 2000a
13 46 17.2......	−45 41 05	Karachentsev et al. 2002b; ph: Jerjen et al. 2000a
13 46 39.5......	−29 58 45	M83 group; ph: Jerjen et al. 2000a
13 47 57.7......	−53 21 04	Karachentsev et al. 2002b
13 48 09.2......	−47 33 54	Cen A group
13 48 33.9......	−37 58 03	Cen A group

TABLE 2—*Continued*

R.A. (J2000.0)	Decl. (J2000.0)	Notes
13 48 46.4......	−46 59 49	Karachentsev et al. 2002b
13 50 51.1......	+38 01 16	Makarova et al. 1998
13 51 22.0......	−47 00 00	Cen A group
13 54 33.6+....	04 14 35	Karachentsev et al. 2002c; ph: Makarova et al. 2002
13 54 45.1......	+53 54 17	Karachentsev et al. 1994; ph: Makarova 1999
13 54 48.7......	+35 50 15	Karachentsev et al. 2003b; ph: Makarova et al. 1998
13 57 01.6......	−35 20 02	Jerjen et al. 2000b, and −0.5 Mpc correction; ph: Jerjen et al. 2000a
14 03 12.8......	+54 21 02	Ferrarese et al. 2000
14 03 21.5......	−41 22 35	Karachentsev et al. 2002b
14 05 02.1......	+53 39 47	Karachentsev et al. 1994; ph: Makarova 1999
14 05 33.1......	+54 27 39	Karachentsev et al. 1994; ph: Bremnes et al. 1999
14 07 10.7......	+35 03 37	E. K. Grebel 2001, private communication
14 13 09.3......	−65 20 21	Circinus. $A_g = 4.2^{\rm m}$ from the galaxy color.
14 15 09.4......	+57 05 15	M101 group
14 15 56.5......	+23 03 19	Aparicio, Tikhonov, & Karachentsev 2000
14 19 48.3......	+56 43 49	Karachentsev et al. 1994
14 24 43.5......	+44 31 33	Karachentsev et al. 2002c; ph: Bremnes et al. 1999
14 28 03.7......	−46 18 06	Karachentsev et al. 2002b
14 35 24.6......	+57 15 24	Makarova & Karachentsev 1998; ph: Bremnes et al. 1999
15 09 11.3......	+67 12 52	van den Bergh 2000
16 13 47.6......	+54 22 16	Karachentsev et al. 2001c
16 20 59.3......	−60 29 15	Woudt, Kraan-Korteweg, & Fairall 1999
17 20 01.4......	+57 54 34	van den Bergh 2000
17 45 09.2......	−28 01 12	The Milky Way
17 49 27.6......	+70 08 41	Karachentsev et al. 2003b; ph: Makarova 1999
18 55 03.1......	−30 28 42	van den Bergh 2000
19 16 41.8......	+63 58 22	Drozdovsky et al. 2001
19 29 59.0......	−17 40 41	Karachentsev et al. 2002c; ph: Karachentsev, Aparicio, & Makarova 1999a (from V, I)
19 44 57.7......	−14 48 11	van den Bergh 2000
20 03 57.4......	−31 40 54	Huchtmeier et al. 2000
20 30 15.3......	+60 26 25	N6946 gr. The average "bs"-distance from Sharina, Karachentsev, & Tikhonov 1997
20 30 32.0......	+60 21 13	N6946 gr. The average "bs"-distance from Sharina et al. 1997; ph: Karachentsev, Sharina, & Huchtmeier 2000c (calculated from R, I)
20 31 32.6......	+60 48 48	N6946 gr. The average "bs"-distance from Sharina et al. 1997; ph: Karachentsev et al. 2000c (calculated from V, I)
20 34 51.1......	+60 09 15	N6946 gr. The average "bs"-distance from Sharina et al. 1997
20 45 20.8......	+60 24 40	N6946 gr. The average "bs"-distance from Sharina et al. 1997; ph: Karachentsev et al. 2000c (calculated from R, I)
20 46 51.8......	−12 50 53	Karachentsev et al. 2002c
20 48 24.1......	+58 37 06	N6946 gr. The average "bs"-distance from Sharina et al. 1997; ph: Karachentsev et al. 2000c (calculated from R, I)
20 51 09.7......	+56 53 24	N6946 group, Burton et al. 1999
21 03 24.2......	+57 17 14	N6946 gr. The average "bs"-distance from Sharina et al. 1997; ph: Karachentsev et al. 2000c (calculated from R, I)
21 05 53.0......	+57 12 19	N6946 group? Should be confirmed as the LV galaxy
22 02 41.9......	−51 17 43	Karachentsev et al. 2002c
22 40 43.9......	−30 47 59	Semiresolved (Whiting et al. 2002)
22 41 49.0......	−64 25 12	V_h from Oosterloo, Da Costa, & Staveley-Smith 1996
23 26 27.5......	−32 23 26	Karachentsev et al. 2002c
23 26 31.8......	+50 40 32	Karachentsev et al. 2003c
23 28 34.1......	+14 44 48	van den Bergh 2000
23 43 46.0......	−31 57 33	Karachentsev et al. 2003a
23 45 34.0......	+38 43 04	Karachentsev et al. 2002c
23 51 46.4......	+24 35 10	Pritzl et al. 2002
23 52 02.8......	−52 34 39	Karachentsev et al. 2003a. V_h and W from HIPASS
23 57 49.4......	−32 35 24	Karachentsev et al. 2003a

TABLE 3
A LIST OF FALSE AND QUESTIONABLE LV OBJECTS

Name	R.A. (J2000.0)	Decl. (J2000.0)	a	B_t	A_g	T	V_h	V_{LG}	Notes
Sc 18	00 00 55.9	−41 08 58	0.7	17.5	0.57	Ir	151	129	Distant
KKH 1	00 14 27.2	+31 19 02	0.7	18.4	0.32	Ir	−104	160	Gal.+Local H I
Sc 24	00 36 39.2	−32 34 22	0.6	18.0	0.06	Ir	79	83	Distant
Sc 42	01 39 14.5	−47 17 58	0.8	16.7	0.06	Ir	162	64	Distant
MB 2	02 36 59.8	+59 14 14	1.1	18	4.03	Neb	...	...	Cirrus
U2397, KKH 15	02 56 15.3	+37 08 07	0.7	16.9	0.50	Ir	−142	32	Gal.+Local H I
U2689	03 21 27.7	+40 48 06	1.5	16.0	0.64	Ir	277	443	Distant
Cam C, KK 26	03 23 06.4	+62 47 08	1.8	15.6	4.80	H II	−64	151	...
N1400	03 39 30.6	−18 41 15	2.5	11.9	0.28	E	558	485	Distant, *
KKH 25	04 22 16.0	+30 26 21	0.7	18.3	2.33	Ir	−135	−49	Gal.+Local H I
KKH 26	04 38 22.9	+07 05 34	2.4	16.5	0.65	Neb	55	32	...
KKH 29	04 56 51.0	+37 56 60	1.2	17.	7.73	Neb	−39	50	...
P16744	05 06 25.2	−31 57 10	2.1	13.38	0.06	H II?	686	497	Knot in NGC 1800
U3327	05 36 40.9	+31 51 16	0.8	14.	19.11	Neb	187	223	Refl.
KKH 33	05 52 48.5	+28 05 56	1.1	16.	3.18	PN	49	57	...
N2242	06 34 07.2	+44 46 37	0.3	15.	0.57	PN	38	99	...
HIZSS 021	07 46 06.0	−28 25 35	...	20	4.34	HVC	494	219	...
VKN	08 40 08.9	+68 26 23	0.9	17.7	0.34	dSph?	...	...	...
HIPASS J0851−75	08 51 29.2	−75 31 52	...	20	0.60	HVC	482	204	...
U4991	09 23 21.3	+22 19 07	0.6	14	0.17	Str	147	39	Star+gal.
KKSG 13	09 35 27.7	−16 20 00	1.2	18.2	0.24	Ir	89	−181	Gal.+Local H I
KKH 59	10 10 15.2	+62 54 50	1.2	15.2	0.06	Ir	−160	−49	Gal.+Local H I?
E264−35	10 42 51.4	−47 36 56	1.5	14.02	0.80	Sd	770	459	V_h = 1099 in HIPASS
HIPASS J1114−77	11 14 14.0	−77 44 26	...	20	4.30	HVC	536	267	...
P35366	11 29 07.7	−29 34 57	1.4	16.6	0.29	SB	749	468	Distant
KKH 70	11 39 26.9	+60 10 18	0.9	16.7	0.06	Ir	−151	−41	Gal.+Local H I
DDO97, U6782	11 48 57.3	+23 50 16	2.0	15.2	0.12	Ir	524	452	Distant, *
A1154−1935	11 56 56.5	−19 52 09	0.2	18	0.12	E?	120	−126	Distant
KKH 75	12 01 55.4	+62 38 31	0.7	18.0	0.08	Neb	−153	−26	Local H I?
MCG 5−29−3	12 05 23.3	+28 21 57	0.6	15.4	0.10	Ir	568	527	Distant
U7131	12 09 11.8	+30 54 25	1.2	15.5	0.10	S	253	226	*
KKs 50	12 09 30.2	−30 21 49	0.6	18.0	0.25	dSph	150	−115	Gal.+Local H I?
N4150	12 10 33.6	+30 24 06	2.3	12.4	0.08	E	226	198	Distant, *
U7321	12 17 34.1	+22 32 22	5.5	14.1	0.12	Sd	409	345	Distant, *
I779, U7369	12 19 38.8	+29 52 59	1.0	14.91	0.08	S0	225	198	Distant
KKH 79	12 19 51.7	+61 31 06	0.5	17.7	0.08	Ir	516	642	*
KKSG 37	12 48 01.0	−12 39 22	0.6	18.7	0.22	dSph	84	−110	Gal.+Local H I
P43501	12 50 57.5	−13 27 31	1.1	14.2	0.23	SB	346	151	Distant
Cen 5	13 04 57.0	−49 24 04	0.8	17.	0.78	Ir	122	−143	Distant
KK 191	13 13 38.9	+42 02 39	0.8	18.2	0.06	Ir?	368	429	*
KK 201	13 25 12.1	−37 37 26	1.3	17.50	0.30	Sph?	...	...	Distant?, *
N5461	14 03 41.0	+54 19 03	0.6	16.	0.04	H II	298	436	H II in M101
KKR 4	14 16 21.5	+13 52 31	1.1	17.	0.11	PN	−17	−36	...
HIPASS J1441−62	14 41 37.0	−62 44 38	...	20	5.74	HVC	671	439	...
KK 236	15 05 31.9	+55 51 51	0.9	19.17	0.06	dSph	−150	24	Gal.+Local H I?
KK 237	15 08 06.6	+56 15 38	0.7	18.21	0.05	Ir	−177	0	Gal.+Local H I?
HIZOA J1526−51	15 26 20.0	−51 11 16	...	20	2.75	HVC	605	415	...
HIPASS J1614−72	16 14 22.5	−72 15 55	...	20	0.39	HVC	383	170	...
HIZOA J1618−55	16 18 16.0	−55 36 48	...	20	2.89	HVC	403	230	...
HIPASS J1712−64	17 12 35.0	−64 38 12	...	20	0.46	HVC	458	283	...
KKH 89	18 32 44.2	+33 23 17	1.5	18.0	0.28	Neb?	−99	150	Cirrus?
HIPASS J1858−67	18 58 25.3	−67 22 58	...	20	0.26	HVC	434	281	...
HIPASS J1913−67	19 13 00.2	−67 05 49	...	20	0.29	HVC	426	278	...
HIPASS J1919−68	19 19 58.4	−68 39 13	...	20	0.29	HVC	242	88	...
HIPASS J1931−68	19 31 32.7	−68 03 31	...	20	0.46	HVC	372	223	...
KKH 90	19 41 50.6	+68 34 14	1.3	18.0	0.67	Neb	−139	146	...
KKH 91	19 57 59.6	+04 47 26	0.9	17.4	0.46	Neb	−4	207	...
PGC 64622	20 24 11.0	+35 51 37	0.7	16	22.91	Neb	−114	186	...
KKH 93	20 58 00.0	+62 20 56	0.5	18.3	2.89	Ir	−112	192	Gal.+Local H I
HIPASS J2118−82	21 18 02.0	−82 32 45	...	20.	0.89	HVC	318	114	...
Sc 2	23 20 37.2	−31 54 22	0.4	18	0.87	Ir	68	109	Distant

NOTE—Notes to Table 3 are as follows: For $03^h39^m30\overset{s}{.}6$, $-18°41'15''$, $D = 26.4$ Mpc (Tonry et al. 2001); $11^h48^m57\overset{s}{.}3$, $+23°50'16''$, $D = 14$ Mpc (Makarova et al. 1998); $12^h09^m11\overset{s}{.}8$, $+30°54'25''$, $D = 14$ Mpc (Makarova et al. 1998); $12^h10^m33\overset{s}{.}6$, $+30°24'06''$, $D = 13.7$ Mpc (Tonry et al. 2001); $12^h17^m34\overset{s}{.}1$, $+22°32'22''$, unresolved with *HST*; $12^h19^m51\overset{s}{.}7$, $+61°31'06''$, $V_h = 320$ (Karachentsev et al. 2001a) should be read as 520 km s^{-1}; $13^h13^m38\overset{s}{.}9$, $+42°02'39''$, possible confusion from NGC 5055 at 24′; $13^h25^m12\overset{s}{.}1$, $-37°37'26''$, unresolved into stars with the WFPC2.

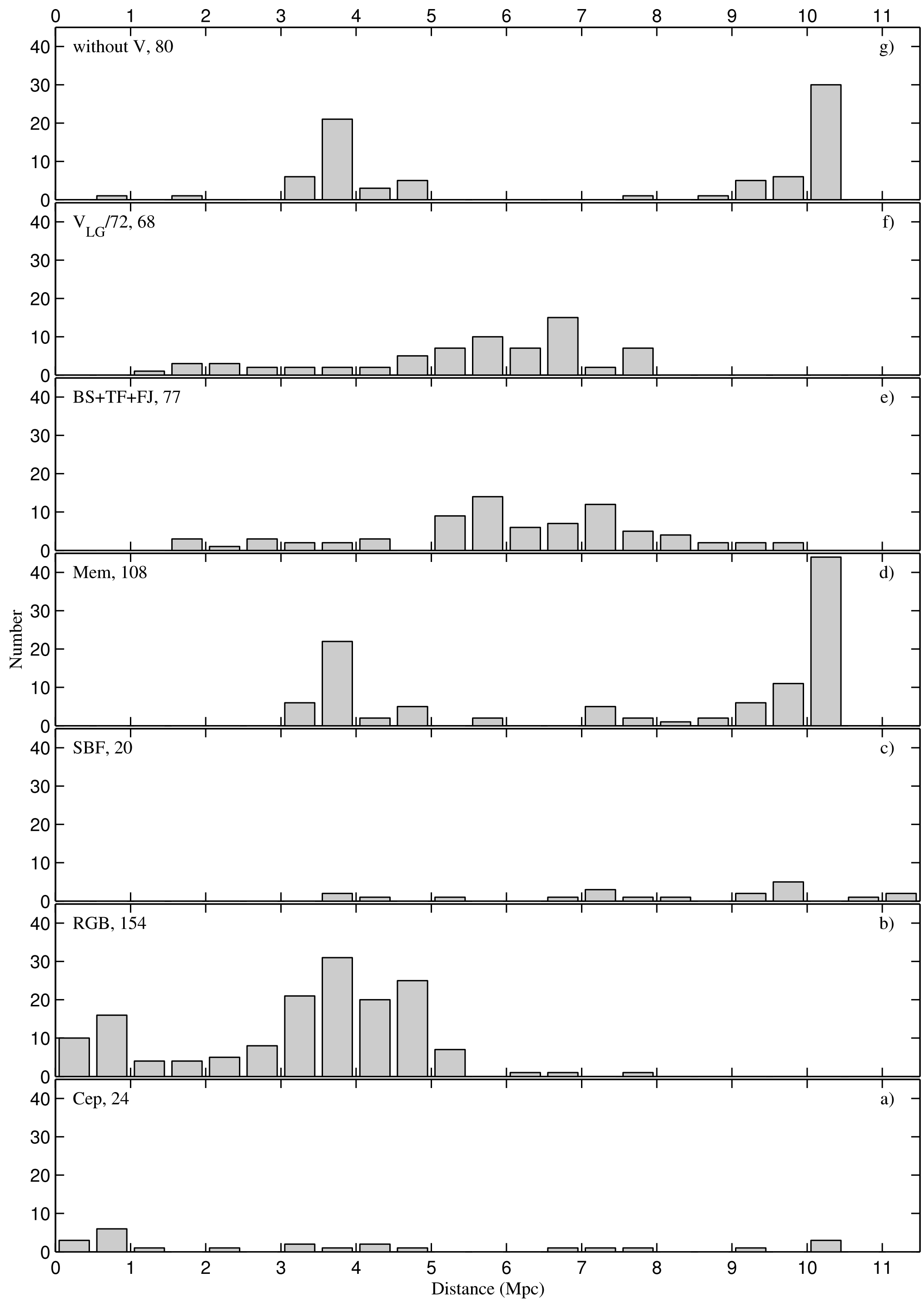

FIG. 1.—Distribution of 451 neighboring galaxies according to their distances derived by different methods: (*a*) from cepheids, (*b*) from the tip of red giant branch stars, (*c*) from surface brightness fluctuations, (*d*) from the galaxy membership in the known nearby groups, (*e*) from the luminosity of the brightest stars (BS) and from the Tully-Fisher (TF) or Faber-Jackson (FJ) relations. (*f*, *g*) Distance distributions for 68 galaxies with distance estimates from the Hubble relation with $H_0 = 72$ km s^{-1} Mpc^{-1} and for 80 galaxies without radial velocities, respectively.

which are collected in Table 4. Its columns contain the following:

1. Galaxy name.
2. Major linear diameter (in kiloparsecs), A_{25}, corrected for galaxy inclination and the Galactic extinction in the manner adopted in RC2 (de Vaucouleurs, de Vaucouleurs, & Corwin 1976):

$$\log A_{25} = \log A_{\rm kpc} - 0.20 \log (a/b) + 0.09 A_g. \tag{3}$$

3. Inclination from face-on in degrees

$$\sin^2 i = [1 - (b/a)^2]/[1 - (b/a)_0^2], \tag{4}$$

where $(b/a)_0 = 0.07$, if $T = 6$ or 7, $(b/a)_0 = 0.12$, if $T = 5$ or 8, $(b/a)_0 = 0.18$, if $T = 4$ or 9, and $(b/a)_0 = 0.20$, if otherwise.

4. H I rotational velocity $V_m = W_{50}/(2 \sin i)$, corrected for inclination and turbulent motions according to Tully & Fouqué (1985) with isotropic noncircular motion parameter $\sigma_z = 8$ km

TABLE 4

Global Parameters of the Neighboring Galaxies

Name (1)	A_{25} (2)	i (3)	V_m (4)	A_i (5)	M_B (6)	Σ_B (7)	$M_{\rm HI}/L$ (8)	$\Sigma_{\rm HI}$ (9)	log M_{25} (10)	M_{25}/L (11)	Θ (12)	MD (13)
WLM	2.58	72	26	0.00	−13.95	24.4	0.67	6.9	8.32	3.54	0.3	M31
E349−031	1.27	36	7	0.00	−12.63	24.2	0.51	6.9	6.88	0.43	0.5	N7793
N55	12.05	83	78	0.57	−18.06	23.7	0.78	7.3	9.93	3.26	−0.4	N300
N59	3.66	64	24	0.00	−15.74	23.4	0.06	6.2	8.41	0.83	−0.6	N253
E410−005	0.70	41	...	0.00	−11.58	24.0	...	...	...	...	0.4	N55
I10	2.71	30	55	0.02	−15.57	22.9	0.37	7.2	8.98	3.68	1.8	M31
Sc 22	1.06	40	...	0.00	−10.45	26.0	...	...	...	...	0.9	N253
Cetus	1.13	31	...	0.00	−10.18	26.4	...	...	...	...	0.5	M31
E294−010	0.56	52	...	0.00	−10.91	24.2	...	...	...	...	1.0	N55
U288	2.47	53	19	0.00	−13.82	24.4	1.03	7.1	8.00	1.90	−2.1	N6946
N147	3.16	54	...	0.00	−14.79	24.0	...	...	...	...	3.0	M31
And III	0.61	55	...	0.00	−9.30	25.9	...	...	...	...	3.5	M31
N185	2.51	33	...	0.00	−14.76	23.5	0.00	4.5	...	...	2.3	M31
N205	4.49	54	5	0.00	−16.15	23.4	0.00	4.3	7.12	0.03	3.7	M31
And IV	2.32	41	...	0.00	−12.60	25.5	...	...	...	...	−2.1	I342
N221	1.99	46	...	0.00	−15.96	21.8	...	...	...	...	6.8	M31
M31	35.87	74	255	1.05	−21.58	22.5	0.07	6.7	11.43	4.07	4.6	N221
DDO 226	2.60	72	22	0.00	−14.17	24.2	0.80	7.0	8.17	2.02	−0.1	N253
And I	0.98	19	...	0.00	−10.87	25.4	...	...	...	...	3.7	M31
N247	20.66	72	101	0.52	−18.81	24.1	0.51	6.9	10.39	4.75	1.3	N253
N253	22.98	79	199	1.24	−21.37	21.7	0.05	6.8	11.03	1.93	0.3	N247
KDG 2	1.19	24	...	0.00	−11.39	25.3	...	...	...	...	0.4	N253
DDO 6	1.40	69	7	0.00	−12.50	24.5	0.57	6.8	6.85	0.46	0.5	N253
E540−032	1.30	24	...	0.00	−11.32	25.6	...	...	...	...	0.6	N253
SMC	5.28	51	51	0.04	−16.35	23.6	0.14	6.5	9.20	2.94	3.5	LMC
N300	12.95	45	94	0.14	−17.92	24.0	0.97	7.2	10.12	5.84	−0.3	N55
Sculptor	0.89	39	...	0.00	−9.77	26.3	0.00	3.6	...	...	2.8	Milky Way
LGS−3	0.20	56	5	0.00	−7.96	24.9	0.64	6.7	5.77	2.46	1.7	M31
I1613	3.49	26	19	0.00	−14.51	24.5	0.59	6.8	8.16	1.47	0.9	M31
U685	1.92	46	45	0.01	−14.44	23.3	0.58	7.3	8.65	4.76	−1.6	N253
KKH 5	0.87	53	19	0.00	−12.27	23.7	0.59	7.1	7.58	3.00	−1.2	I342
N404	3.25	19	108	0.03	−16.49	22.4	0.15	7.1	9.65	7.20	−1.0	Maffei 2
And V	0.50	43	...	0.00	−8.41	26.4	...	...	...	...	2.8	M31
And II	0.57	52	...	0.00	−9.33	25.8	...	...	...	...	2.4	M31
U1104	2.08	55	36	0.00	−15.17	22.7	0.71	7.6	8.51	1.77	0.5	N628
M33	16.09	53	104	0.23	−18.87	23.5	0.38	7.0	10.31	3.70	2.0	M31
KKH 6	1.09	55	15	0.00	−12.42	24.1	0.95	7.2	7.45	1.95	−0.8	Maffei 2
N625	5.96	77	41	0.00	−16.53	23.7	0.22	6.7	9.06	1.81	−0.4	N253
N628	22.54	20	68	0.02	−19.84	23.2	0.51	7.2	10.08	0.90	0.0	U1171
U1171	2.86	24	21	0.00	−13.86	24.7	0.33	6.4	8.17	2.71	2.4	N628
DDO 13	12.14	40	24	0.00	−15.72	26.0	1.77	6.7	8.92	2.73	−1.4	N628
KK 13	1.42	57	14	0.00	−13.11	24.0	0.40	6.8	7.54	1.28	1.0	N672
KDG 10	3.40	49	17	0.00	−13.53	25.4	1.29	6.8	8.05	2.78	−0.1	N628
KK 14	2.93	71	13	0.00	−12.13	26.5	2.96	6.7	7.75	5.08	2.0	N672
E245−05	4.85	28	58	0.02	−15.59	24.1	1.36	7.3	9.28	7.13	−0.7	N253
KK 15	1.08	74	6	0.00	−11.43	25.1	1.67	7.0	6.71	0.88	2.1	N672
I1727	12.64	72	55	0.14	−17.71	24.1	0.78	7.1	9.65	2.36	4.2	N672
N672	13.19	69	100	0.46	−18.76	23.2	0.35	7.1	10.19	3.10	3.8	I1727
U1281	5.23	80	53	0.21	−16.04	23.9	0.60	7.1	9.24	4.28	−1.2	N784
Phoenix	0.59	33	10	0.00	−10.22	24.9	0.03	5.4	6.81	3.42	0.8	Milky Way
KK 16	0.96	71	8	0.00	−12.37	23.9	0.27	6.7	6.87	0.53	−0.4	N784
KK 17	0.75	62	5	0.00	−11.41	24.3	0.36	6.7	6.30	0.35	−0.3	N784
N784	7.55	79	41	0.00	−16.59	24.1	0.52	6.9	9.18	2.23	−1.5	U1281
Cas 1	4.99	41	34	0.00	−16.70	23.1	0.17	6.8	8.82	0.88	0.5	I342
N855	6.47	71	54	0.12	−17.07	23.3	0.07	6.3	9.34	2.08	−0.8	N925
N891	30.80	78	220	1.07	−20.59	23.2	0.10	6.6	11.24	6.45	−1.2	N925
KKH 11	1.68	59	47	0.03	−13.35	24.1	1.42	7.3	8.64	12.84	1.0	Maffei 2
KKs 3	2.60	69	...	0.00	−12.28	26.1	...	...	...	...	−0.3	N1313
N925	28.60	56	110	0.29	−19.88	23.7	0.40	6.9	10.60	2.87	−0.9	N891
KKH 12	2.36	78	23	0.00	−13.03	25.2	1.36	6.9	8.18	5.95	1.6	Maffei 1
MB 1	4.22	65	24	0.00	−14.81	24.6	0.13	6.1	8.45	2.19	4.4	Maffei 1
Maffei 1	6.25	33	...	0.00	−18.97	21.3	...	...	...	...	2.7	MB 1
E115−21	6.83	83	51	0.18	−15.30	25.2	2.38	7.1	9.31	10.01	−0.6	N1313
Fornax	1.17	40	...	0.00	−11.50	25.2	0.00	3.2	...	...	2.3	Milky Way

TABLE 4—*Continued*

Name (1)	A_{25} (2)	i (3)	V_m (4)	A_i (5)	M_B (6)	Σ_B (7)	M_{HI}/L (8)	Σ_{HI} (9)	log M_{25} (10)	M_{25}/L (11)	Θ (12)	MD (13)
Maffei 2	11.96	62	163	0.49	−20.15	21.6	0.03	6.7	10.57	2.07	1.9	Dw 1
Dw 2	2.10	69	43	0.00	−14.55	23.4	1.97	7.8	8.65	4.32	2.5	MB 3
MB 3	3.56	76	5	0.00	−13.65	25.4	0.05	5.3	7.01	0.23	2.9	Dw 2
E154−23	9.79	82	55	0.24	−16.40	24.9	1.91	7.2	9.54	6.19	−1.3	N1311
Dw 1	5.56	51	109	0.21	−18.78	21.3	0.07	7.2	9.89	1.52	2.5	Maffei 2
N1156	8.86	32	52	0.02	−17.70	23.3	0.46	7.1	9.45	1.50	−1.7	N925
KKH 18	0.96	57	17	0.00	−12.39	23.8	0.99	7.3	7.48	2.17	−1.6	I342
N1313	11.45	37	116	0.12	−19.02	22.6	0.29	7.3	10.25	2.84	−1.6	KK 27
N1311	3.95	80	49	0.11	−15.72	23.6	0.17	6.6	9.05	3.70	−1.0	E154−23
U2684	3.36	62	40	0.00	−13.37	25.6	2.84	7.0	8.79	17.53	0.2	U2716
KK 27	1.19	74	...	0.00	−11.83	24.9	...	...	...	...	1.3	N1313
U2716	3.04	57	23	0.00	−15.01	23.7	0.43	7.0	8.26	1.16	−0.4	U2684
U2773	3.00	42	52	0.03	−16.12	22.6	0.43	7.4	8.97	2.14	−1.6	I342
I1959	4.21	82	58	0.24	−16.06	23.4	0.31	7.0	9.22	4.00	−1.9	N253
KK 35	3.57	48	18	0.00	−14.30	24.8	0.02	5.3	8.15	1.74	2.4	I342
I342	32.71	25	162	0.07	−20.69	23.2	0.42	7.2	11.00	3.41	−0.1	KK 35
U2905	2.02	55	26	0.00	−14.41	23.4	0.20	6.8	8.19	1.71	−2.1	M31
UA86	7.47	48	50	0.03	−17.71	23.0	0.26	7.1	9.33	1.14	0.3	I342
I2038	2.63	77	50	0.12	−14.42	24.0	3.03	7.7	8.88	8.38	−2.2	N253
Cam A	4.58	57	23	0.00	−14.06	25.6	2.52	7.0	8.44	4.17	0.1	I342
N1569	3.34	61	37	0.00	−18.17	20.8	0.02	6.9	8.74	0.19	−0.4	UA 92
UA92	1.85	62	39	0.00	−14.48	23.2	0.83	7.5	8.51	3.33	1.1	N1569
N1560	7.96	82	56	0.27	−16.87	23.9	1.43	7.4	9.46	3.30	1.0	I342
Cam B	2.26	62	5	0.00	−11.85	26.2	1.53	6.5	6.77	0.69	1.0	I342
N1705	2.67	43	61	0.06	−15.86	22.6	0.14	6.9	9.06	3.32	−1.7	N1313
UA 105	6.10	51	67	0.11	−16.81	23.4	0.39	7.0	9.50	3.84	0.3	I342
LMC	9.75	32	54	0.02	−17.93	23.3	0.25	6.9	9.52	1.41	3.6	Milky Way
U3303	8.06	39	116	0.13	−16.03	24.8	0.89	6.8	10.10	31.40	−1.8	Orion
KK 49	1.65	46	34	0.00	−14.94	22.5	0.38	7.4	8.34	1.48	−2.0	Orion
Orion	6.19	62	81	0.25	−17.04	23.2	0.73	7.4	9.67	4.59	−2.2	U3303
A0554+07	1.13	69	23	0.00	−12.85	23.7	0.93	7.3	7.86	3.35	−1.7	KK 49
KKH 34	1.34	58	9	0.00	−12.30	24.7	0.92	6.9	7.11	1.00	−0.8	M81
E364−29	7.61	46	40	0.00	−16.04	24.7	0.54	6.7	9.15	3.49	−2.6	N253
E121−20	1.69	46	40	0.00	−12.49	25.0	3.98	7.4	8.50	20.53	−1.6	N1313
E489−56	0.80	62	9	0.00	−13.07	22.8	0.46	7.4	6.85	0.27	−2.1	M31
U3476	1.98	77	38	0.00	−14.27	23.5	1.79	7.7	8.51	4.13	−2.2	N2683
E490−17	2.12	42	26	0.00	−14.46	23.5	0.29	6.9	8.24	1.82	−1.9	M31
KK 55	1.54	52	29	0.00	−13.71	23.5	0.37	7.0	8.16	3.06	−2.6	E364−29
Carina	0.69	49	...	0.00	−8.97	26.5	...	...	...	...	2.7	Milky Way
KKSG 9	1.33	47	117	0.19	−14.00	22.9	4.25	8.3	9.32	34.01	−1.7	E558−11
KKH 37	0.97	49	6	0.00	−11.17	25.1	0.69	6.6	6.54	0.76	−0.3	I342
U3600	2.63	72	39	0.00	−13.53	24.9	1.71	7.1	8.67	11.61	−1.9	N2683
HIZSS 03	1.85	55	35	0.00	−12.19	25.5	1.27	6.7	8.43	22.88	−0.6	M31
FG 202	4.63	64	38	0.00	−14.01	25.6	2.37	6.9	8.90	12.64	−2.0	N1313
E558−11	4.52	57	61	0.11	−16.51	23.1	0.13	6.7	9.29	3.10	−2.5	M31
U3698	2.06	55	20	0.00	−14.30	23.6	0.75	7.3	7.96	1.13	−1.5	N2337
N2337	4.69	42	95	0.12	−16.51	23.2	0.93	7.5	9.70	7.92	−1.9	N2683
U3755	2.51	55	16	0.00	−14.90	23.4	0.30	6.9	7.88	0.54	−2.1	M31
U3817	4.29	62	16	0.00	−14.15	25.3	2.43	7.1	8.12	1.85	−1.7	N2337
DDO 43	2.89	48	19	0.00	−14.75	23.9	1.32	7.4	8.10	1.03	−1.5	N2337
N2366	5.71	72	45	0.02	−16.02	24.1	1.76	7.4	9.12	3.29	1.0	N2403
E059−01	2.27	37	81	0.07	−14.38	23.7	0.42	7.0	9.24	19.93	−1.5	Circinus
DDO 44	2.67	49	...	0.00	−12.07	26.4	...	...	...	...	1.7	N2403
N2403	19.43	56	128	0.34	−19.29	23.5	0.45	7.1	10.57	4.59	−0.0	M81
DDO 46	3.68	20	90	0.02	−15.01	24.1	1.74	7.4	9.54	22.29	−1.5	N2683
DDO 47	4.65	33	46	0.01	−15.10	24.6	2.80	7.4	9.06	6.70	−2.1	N2683
KK 65	1.08	58	18	0.00	−13.15	23.3	0.44	7.1	7.62	1.48	−2.0	M31
U4115	2.62	58	43	0.01	−13.59	24.8	3.46	7.4	8.76	13.50	−1.7	DDO 47
N2537	3.49	29	88	0.05	−16.65	22.4	0.15	7.0	9.50	4.38	−1.4	N2683
Ho II	7.78	31	56	0.02	−16.72	24.0	1.29	7.3	9.45	3.75	0.6	M81
E006−01	2.54	25	...	0.00	−14.86	23.5	...	...	...	...	−2.4	N4945
KDG 52	1.35	24	15	0.00	−11.49	25.5	2.15	7.0	7.57	5.98	0.7	M81
DDO 52	2.98	62	42	0.00	−13.67	25.0	1.49	7.0	8.78	13.11	−1.5	N2683
DDO 53	1.67	30	17	0.00	−13.37	24.1	1.19	7.3	7.73	1.56	0.7	M81
U4483	1.04	56	16	0.00	−12.73	23.7	1.69	7.6	7.49	1.60	0.5	M81

TABLE 4—*Continued*

Name (1)	A_{25} (2)	i (3)	V_m (4)	A_i (5)	M_B (6)	Σ_B (7)	M_{HI}/L (8)	Σ_{HI} (9)	log M_{25} (10)	M_{25}/L (11)	Θ (12)	MD (13)
N2683	16.18	82	205	1.19	−20.43	21.9	0.03	6.6	10.90	3.44	0.2	KK 69
KK 69	4.34	42	5	0.00	−12.76	26.7	1.98	6.4	7.10	0.64	3.2	N2683
KK 70	1.10	38	...	0.00	−11.86	24.7	...	...	...	...	2.8	N2683
D564−08	0.99	46	30	0.00	−11.76	24.5	1.55	7.2	8.01	13.07	−1.9	N2683
KKH 46	1.17	31	16	0.00	−11.93	24.7	2.23	7.3	7.55	3.87	−2.0	N2903
D634−03	0.28	29	29	0.00	−9.56	24.0	...	...	7.43	25.68	−1.3	M31
KK 71	1.27	40	...	0.00	−12.03	24.8	...	...	...	...	1.5	N2784
DDO 56	4.08	41	52	0.03	−15.07	24.3	0.42	6.7	9.10	7.62	1.2	N2784
KK 72	1.94	35	...	0.00	−12.76	25.0	...	...	...	...	3.3	N2784
N2784 dw 1	3.01	28	...	0.00	−13.58	25.1	...	...	...	...	4.3	N2784
N2784	15.86	69	...	0.00	−19.72	22.6	...	...	...	...	2.0	KK 73
KK 73	2.47	53	...	0.00	−14.46	23.8	...	...	...	...	4.1	N2784
E564−30	3.86	54	68	0.13	−14.27	25.0	6.24	7.6	9.32	26.19	−2.0	N3115
N2787	7.13	52	214	0.39	−18.50	22.1	0.03	6.4	10.58	9.71	−1.9	M81
D565−06	1.11	31	5	0.00	−11.88	24.6	0.23	6.3	6.51	0.36	−1.8	N2683
E565−03	1.29	49	16	0.00	−14.14	22.7	0.30	7.2	7.59	0.55	−1.8	N3115
N2915	2.35	60	69	0.16	−16.04	22.1	0.72	7.8	9.12	3.23	−1.3	Circinus
KKH 51	1.80	38	14	0.00	−12.92	24.7	0.65	6.8	7.60	1.73	1.6	N2903
N2903	28.94	63	198	0.60	−21.00	22.6	0.11	6.8	11.12	3.36	1.8	U5086
U5086	2.61	26	...	0.00	−13.95	24.4	0.76	6.9	...	...	4.6	N2903
Ho I	4.03	37	18	0.00	−14.49	24.8	1.38	7.0	8.20	1.62	1.5	M81
U5186	1.97	83	18	0.00	−12.98	24.8	0.70	6.7	7.88	3.10	−1.0	N2683
F8D1	2.58	25	...	0.00	−12.59	25.8	...	...	...	...	2.0	M81
U5209	1.77	17	34	0.00	−13.15	24.4	0.41	6.7	8.38	8.43	−1.1	N2683
FM 1	0.93	28	...	0.00	−10.48	25.7	...	...	...	...	1.8	M82
N2976	5.57	63	48	0.05	−17.10	22.9	0.17	6.9	9.17	1.38	2.7	M81
KK 77	2.61	42	...	0.00	−12.03	26.4	...	...	...	...	2.0	M81
KK 78	0.95	55	...	0.00	−11.67	24.5	...	...	...	...	4.2	DDO 64
DDO 64	3.65	71	38	0.00	−13.95	25.2	2.70	7.2	8.80	10.63	3.3	KK 78
U5288	2.41	53	51	0.05	−14.74	23.5	2.22	7.8	8.87	6.03	−1.6	N2903
BK 3N	0.60	38	5	0.00	−9.59	25.6	...	...	6.24	1.64	1.0	M81
KKSG 15	2.84	83	24	0.00	−14.97	23.6	0.39	7.0	8.28	1.24	0.7	N3115
M81	26.85	61	232	0.59	−21.06	22.4	0.07	6.7	11.22	4.07	2.2	M82
M82	10.93	69	70	0.26	−19.63	21.9	0.07	6.9	9.80	0.57	2.7	M81
DDO 68	3.86	71	36	0.00	−14.33	24.9	1.85	7.1	8.77	7.07	−1.6	N3344
KDG 61	2.40	56	...	0.00	−12.85	25.4	...	...	...	...	3.9	M81
A0952+69	2.14	26	...	0.00	−11.51	26.5	...	...	...	...	1.9	N3077
Ho IX	2.77	38	50	0.02	−13.68	24.8	11.03	7.9	8.90	17.37	3.3	M81
Leo A	0.94	54	4	0.00	−11.36	24.8	0.98	6.9	6.14	0.25	0.2	Milky Way
KKSG 16	2.58	40	...	0.00	−12.24	26.1	...	...	...	...	1.1	N3115
Sex B	1.93	48	22	0.00	−13.96	23.8	0.78	7.2	8.02	1.76	−0.7	Milky Way
KKH 57	0.67	35	...	0.00	−10.19	25.3	...	...	...	...	0.7	M81
KKSG 17	3.05	86	21	0.00	−14.69	24.0	0.37	6.8	8.21	1.39	2.0	N3115
UA 193	7.38	90	58	0.43	−15.70	25.0	1.59	7.0	9.45	9.60	1.3	N3115
N3109	5.81	86	51	0.17	−15.68	24.4	2.35	7.4	9.25	6.09	−0.1	Antlia
N3077	6.14	35	51	0.02	−17.76	22.5	0.32	7.3	9.26	0.93	1.9	M81
Antlia	0.78	42	9	0.00	−9.75	26.0	1.02	6.4	6.87	6.02	2.3	N3109
U5427	2.34	48	46	0.02	−14.48	23.7	0.32	6.9	8.76	5.99	−1.2	N3344
BK 5N	0.87	42	...	0.00	−10.61	25.4	...	...	...	...	2.4	N3077
KDG 63	1.84	29	6	0.00	−12.12	25.5	...	...	6.87	0.68	1.8	M81
N3115	17.13	73	115	0.55	−20.82	21.7	0.00	5.3	10.42	0.80	1.9	KKSG 18
U5423	1.37	49	25	0.00	−14.54	22.5	0.25	7.2	8.00	0.99	−0.9	M81
KK 84	3.43	48	...	0.00	−13.60	25.4	...	...	...	...	4.5	N3115
KKSG 18	4.25	38	...	0.00	−16.55	22.9	0.01	5.8	...	...	3.6	N3115
KDG 64	1.85	64	...	0.00	−12.57	25.1	...	...	...	...	2.5	M81
U5456	1.60	62	31	0.00	−14.24	23.1	0.13	6.7	8.24	2.26	−1.8	M31
IKN	3.19	33	...	0.00	−11.44	27.4	...	...	...	...	2.7	M81
Leo I	0.70	42	...	0.00	−10.97	24.6	...	...	...	...	1.5	Milky Way
Sex A	2.27	35	49	0.01	−13.95	24.1	1.45	7.3	8.80	10.74	−0.6	Milky Way
Sex dSph	0.68	69	...	0.00	−7.98	27.5	...	...	...	...	2.8	Milky Way
KKH 60	0.32	58	...	0.00	−7.97	25.8	...	...	...	...	−0.8	Sex B
HIJASS	...	...	...	0.00	−7.93	...	...	...	...	...	2.2	M81
HS 117	1.61	55	5	0.00	−11.83	25.5	...	...	6.67	0.56	1.9	M81
KKSG 19	1.09	31	20	0.00	−12.09	24.4	1.97	7.3	7.70	4.72	−2.0	N3621
DDO 78	2.20	19	...	0.00	−12.17	25.8	...	...	...	...	1.8	M81

TABLE 4—*Continued*

Name (1)	A_{25} (2)	i (3)	V_m (4)	A_i (5)	M_B (6)	Σ_B (7)	M_{HI}/L (8)	Σ_{HI} (9)	log M_{25} (10)	M_{25}/L (11)	Θ (12)	MD (13)
U5672	2.61	78	37	0.00	−14.65	23.7	0.22	6.7	8.62	3.70	−0.7	N3344
I2574	13.35	68	55	0.12	−17.46	24.5	1.13	7.1	9.67	3.11	0.9	M81
DDO 82	3.70	55	...	0.00	−14.63	24.5	...	...	...	...	0.9	M81
LeG 01	1.41	55	...	0.00	−11.53	25.5	...	...	...	...	0.3	N3299
N3274	3.50	62	80	0.24	−16.16	22.9	1.20	7.8	9.41	5.74	−0.3	N3344
LeG 02	0.92	19	...	0.00	−11.12	25.0	...	...	...	...	0.1	N3368
BK 6N	1.14	52	...	0.00	−11.08	25.5	...	...	...	...	1.1	M81
LeG 03	2.54	82	...	0.00	−12.40	25.9	...	...	...	...	−0.1	N3368
N3299	6.46	40	83	0.09	−16.99	23.4	0.08	6.3	9.71	5.28	0.5	N3368
LeG 04	1.62	62	...	0.00	−11.51	25.9	...	...	...	...	2.0	FS 01
FS 01	1.91	60	...	0.00	−13.43	24.3	...	...	...	...	1.2	LeG 04
P031727	1.72	51	13	0.00	−11.94	25.5	1.51	6.8	7.50	3.41	0.5	N3368
U5812	3.86	69	26	0.00	−15.09	24.2	0.22	6.5	8.50	1.86	1.2	N3368
P031877	3.44	65	23	0.00	−14.50	24.5	0.24	6.4	8.31	2.09	1.5	N3368
LeG 09	2.39	35	...	0.00	−11.70	26.5	...	...	...	...	1.7	N3368
N3344	14.16	26	155	0.07	−19.03	23.0	0.33	7.1	10.60	6.25	−1.5	N3274
LeG 10	0.31	19	...	0.00	−11.00	22.7	...	...	...	...	2.1	N3368
N3351	20.53	48	168	0.28	−19.88	23.0	0.05	6.3	10.83	4.81	0.8	N3368
LeG 11	1.52	26	...	0.00	−11.43	25.8	...	...	...	...	0.0	N3368
LeG 12	0.59	38	...	0.00	−11.12	24.1	...	...	...	...	2.2	N3368
FS 09	1.05	66	...	0.00	−12.76	23.7	...	...	...	...	2.8	N3368
FS 13	1.20	29	...	0.00	−11.49	25.2	...	...	...	...	1.6	N3368
FS 14	1.83	24	...	0.00	−11.92	25.7	...	...	...	...	0.8	N3412
FS 15	1.23	12	...	0.00	−11.20	25.6	...	...	...	...	4.3	N3368
FS 17	2.06	40	...	0.00	−13.21	24.7	...	...	...	...	2.6	N3368
N3368	21.82	48	216	0.33	−20.42	22.6	0.04	6.4	11.07	5.15	0.6	FS 15
LeG 18	2.33	42	15	0.00	−11.29	26.9	1.51	6.3	7.81	12.59	2.5	FS 20
FS 20	1.44	47	...	0.00	−11.99	25.1	...	...	...	...	2.2	LeG 18
FS 21	2.41	33	24	0.00	−12.72	25.5	2.12	6.9	8.20	8.27	2.4	LeG 21
LeG 21	1.17	42	...	0.00	−11.62	25.0	...	...	...	...	2.9	FS 21
DDO 88	6.08	19	58	0.01	−15.75	24.5	0.40	6.6	9.38	7.73	−0.5	N3351
FS 23	1.18	42	...	0.00	−12.45	24.2	...	...	...	...	1.2	N3412
P32250	2.68	83	24	0.00	−14.18	24.3	0.80	7.0	8.27	2.55	−1.0	N3621
N3377	15.27	56	...	0.00	−19.10	23.1	...	...	...	...	0.8	N3379
N3379	17.46	28	...	0.00	−20.10	22.4	...	...	...	...	1.0	N3384
N3384	15.91	66	18	0.00	−19.55	22.8	0.01	5.5	8.78	0.06	1.2	N3379
P032327	0.59	38	46	0.01	−12.70	22.5	0.75	7.7	8.17	7.91	2.3	N3368
P032348	3.27	62	...	0.00	−14.76	24.1	...	...	...	...	1.3	N3412
DDO 87	5.11	26	64	0.02	−14.42	25.4	1.17	6.7	9.38	26.46	−1.5	N2787
FS 40	1.72	50	...	0.00	−12.22	25.3	...	...	...	...	2.0	N3368
P1424345	0.31	26	...	0.00	−10.24	23.5	...	...	...	...	2.5	N3412
LeG 23	0.92	19	...	0.00	−11.13	25.0	...	...	...	...	3.3	N3412
U5944	3.40	14	...	0.00	−15.42	23.6	...	...	...	...	3.2	N3412
KK 96	2.96	38	...	0.00	−11.90	26.8	...	...	...	...	1.7	N3368
N3412	9.94	58	...	0.00	−18.76	22.5	...	...	...	...	1.9	U5944
LeG 26	2.72	28	...	0.00	−12.99	25.5	...	...	...	...	1.8	N3412
LeG 27	1.14	49	...	0.00	−11.58	25.0	...	...	...	...	0.7	N3412
KDG 73	0.63	35	5	0.00	−10.83	24.5	0.96	7.0	6.26	0.55	1.3	M81
LeG 28	1.62	62	...	0.00	−11.90	25.5	...	...	...	...	0.8	N3368
D640−16	0.89	35	...	0.00	−11.59	24.5	...	...	...	...	1.1	N3489
D640−12	2.14	62	...	0.00	−11.77	26.2	...	...	...	...	0.7	N3368
D640−13	1.92	57	...	0.00	−12.51	25.2	...	...	...	...	0.7	N3368
E215−09	2.56	55	28	0.00	−12.93	25.4	11.75	7.7	8.37	10.06	−0.9	N4945
D640−14	1.69	50	...	0.00	−11.66	25.8	...	...	...	...	0.5	N3489
LeG 32	1.16	42	...	0.00	−11.48	25.2	...	...	...	...	1.2	N3489
N3489	9.60	57	172	0.42	−19.45	21.8	0.00	5.5	10.52	3.53	1.5	D640−08
LeG 33	1.36	58	...	0.00	−11.57	25.4	...	...	...	...	3.0	N3489
D640−08	3.30	56	...	0.00	−13.20	25.7	...	...	...	...	4.0	N3489
LeG 35	1.63	62	...	0.00	−12.15	25.2	...	...	...	...	−0.4	N3627
CGCG 66−109	3.35	78	...	0.00	−14.45	24.5	...	...	...	...	0.3	N3489
U6145	3.31	64	27	0.00	−13.14	25.8	0.92	6.5	8.46	10.34	−1.7	N4594
Leo II	0.73	24	...	0.00	−9.23	26.4	...	...	...	...	1.7	Milky Way
N3593	9.00	67	96	0.36	−17.77	23.3	0.05	6.2	9.98	4.79	−1.7	N3344
N3621	22.05	64	138	0.51	−19.81	23.2	0.63	7.3	10.69	3.72	−1.9	N4594
N3627	23.98	65	178	0.59	−21.14	22.1	0.01	6.0	10.95	2.00	−0.7	N3489

TABLE 4—*Continued*

Name (1)	A_{25} (2)	i (3)	V_m (4)	A_i (5)	M_B (6)	Σ_B (7)	M_{HI}/L (8)	Σ_{HI} (9)	log M_{25} (10)	M_{25}/L (11)	Θ (12)	MD (13)
U6456	1.72	60	7	0.00	−14.03	23.5	0.98	7.4	7.00	0.16	−0.3	M81
U6541	1.44	57	13	0.00	−13.71	23.4	0.23	6.8	7.46	0.61	−0.7	M81
N3738	3.52	44	50	0.03	−16.61	22.4	0.18	7.1	9.00	1.47	−1.0	M81
N3741	1.60	58	42	0.00	−13.13	24.2	4.07	7.8	8.52	11.84	−0.8	M81
E320−14	2.12	52	5	0.00	−13.26	24.7	0.49	6.6	6.85	0.22	−1.2	N5128
KK 109	0.74	49	4	0.00	−9.73	25.9	2.72	6.9	6.09	1.01	−0.6	N4244
DDO 99	2.64	71	16	0.00	−13.52	24.9	1.30	7.0	7.91	2.05	−0.5	N4214
E379−07	1.68	44	38	0.00	−12.31	25.1	1.26	6.9	8.45	21.61	−1.3	N5128
BTS 76	0.51	55	...	0.00	−12.60	22.2	...	...	...	...	−1.2	N4826
N4068	4.34	60	26	0.00	−15.48	24.0	0.81	7.1	8.53	1.39	−1.0	N4736
MCG 6−27−17	0.89	76	...	0.00	−12.97	23.1	...	...	...	...	0.6	N4244
N4144	12.89	79	68	0.39	−18.25	23.6	0.35	6.9	9.84	2.26	−0.9	N4460
N4163	1.96	34	21	0.00	−14.24	23.5	0.37	7.0	8.00	1.29	0.1	N4190
KK 127	0.37	69	17	0.00	−10.59	23.6	0.43	7.0	7.10	4.67	−0.8	M31
N4190	1.73	29	41	0.00	−14.33	23.2	0.80	7.5	8.53	4.08	0.0	N4163
E321−014	1.19	67	9	0.00	−12.70	24.0	0.37	6.8	7.06	0.62	−0.3	N5128
U7242	2.03	68	29	0.00	−13.65	24.2	0.70	7.0	8.30	4.48	0.4	N4236
DDO 113	1.24	30	...	0.00	−11.67	25.1	...	...	...	...	1.6	N4214
N4214	7.05	40	43	0.00	−17.19	23.4	0.56	7.2	9.18	1.29	−0.7	DDO 113
MCG 9−20−131	0.95	77	...	0.00	−12.46	23.7	...	...	...	...	−0.7	M81
U7298	1.22	58	8	0.00	−12.27	24.5	1.61	7.2	6.97	0.75	−0.7	N4449
N4236	23.58	73	76	0.35	−18.59	24.6	0.67	6.8	10.20	3.73	−0.4	M81
N4244	16.60	85	93	0.92	−18.60	23.8	0.47	7.0	10.23	3.94	−0.0	N4736
I3104	3.07	64	20	0.00	−14.85	23.9	0.11	6.3	8.13	1.00	−0.5	Circinus
N4258	35.61	69	215	0.81	−21.25	22.8	0.13	6.8	11.28	3.91	−0.7	N5194
U7356	1.78	8	...	0.00	−13.46	24.1	...	...	...	...	−0.3	N4258
I3308	1.13	84	57	0.32	−12.91	23.7	1.89	7.6	8.63	18.83	−1.1	N4736
KK 144	2.24	74	20	0.00	−12.59	25.5	4.70	7.3	8.01	6.02	−0.9	N4826
N4395	17.30	35	85	0.07	−17.85	24.7	0.77	6.8	10.16	6.77	0.1	N4244
UA 281	1.27	42	30	0.00	−13.70	23.1	1.62	7.8	8.11	2.75	−1.0	N4736
DDO 126	4.18	52	33	0.00	−14.38	25.0	1.79	7.1	8.72	6.00	0.1	N4736
DDO 125	2.88	58	14	0.00	−14.16	24.4	0.46	6.7	7.80	0.87	−0.9	N4214
U7584	1.57	62	16	0.00	−13.30	24.0	1.20	7.3	7.69	1.49	−0.1	N4826
KKH 80	1.85	49	5	0.00	−12.47	25.2	0.70	6.6	6.73	0.36	−0.1	N4826
N4449	7.21	46	84	0.12	−18.27	22.3	1.05	7.9	9.78	1.89	−0.0	N4736
DDO 127	3.55	62	33	0.00	−14.28	24.8	1.75	7.2	8.66	5.65	−0.7	N4258
U7605	1.35	44	16	0.00	−13.53	23.4	0.66	7.3	7.59	0.96	0.7	N4244
N4460	8.92	77	55	0.16	−17.89	23.2	0.02	5.9	9.50	1.40	−0.7	N4144
KK 149	1.35	52	33	0.00	−14.06	22.9	0.41	7.3	8.24	2.67	−0.8	N4258
U7639	5.04	47	23	0.00	−15.44	24.4	0.34	6.6	8.50	1.37	1.4	N4258
KK 151	2.00	67	15	0.00	−13.41	24.4	0.58	6.8	7.72	1.45	−0.4	N4258
DDO 133	10.87	48	32	0.00	−15.47	26.0	1.59	6.6	9.10	5.26	−1.1	N4826
N4523	3.56	48	79	0.13	−14.91	24.2	1.36	7.3	9.42	18.21	−0.8	N4826
KKSG 29	3.28	67	...	0.00	−13.42	25.5	...	...	...	...	2.2	N4594
Arp 211	1.24	57	12	0.00	−13.47	23.3	0.36	7.1	7.29	0.51	−0.7	N4258
KKSG 30	2.42	62	17	0.00	−12.68	25.6	1.10	6.6	7.90	4.36	1.2	N4594
KKSG 31	2.17	29	...	0.00	−12.27	25.7	...	...	...	...	2.3	N4594
UA 292	0.85	47	13	0.00	−11.43	24.5	5.62	7.8	7.25	3.05	−0.4	N4214
KKSG 32	1.63	35	...	0.00	−11.55	25.8	...	...	...	...	4.5	N4594
N4594	21.10	61	413	0.81	−21.90	21.0	0.00	5.9	11.62	4.68	0.3	KKSG 32
N4605	7.41	69	63	0.21	−17.96	22.7	0.13	6.9	9.53	1.44	−1.1	M81
KKSG 33	1.13	8	...	0.00	−11.54	25.0	...	...	...	...	2.9	N4594
N4600	2.43	49	...	0.00	−15.72	22.5	...	...	...	...	−1.2	N4594
KKSG 34	2.21	29	...	0.00	−12.07	26.0	...	...	...	...	3.5	N4594
U7857	2.21	52	...	0.00	−14.40	23.6	...	...	...	...	−0.8	N4826
N4627	6.68	48	...	0.00	−16.87	23.6	...	...	...	...	−1.4	N4258
I3687	4.49	29	41	0.00	−14.64	24.9	0.96	6.8	8.95	7.94	1.1	N4736
KK 160	1.08	42	9	0.00	−11.52	25.0	0.61	6.6	7.01	1.65	1.0	N4736
KKs 51	0.87	66	...	0.00	−11.46	24.5	...	...	...	...	0.7	N5128
E381−018	1.62	56	26	0.00	−13.00	24.4	0.45	6.7	8.12	5.30	−0.6	N5236
E381−20	3.54	69	37	0.00	−14.15	24.9	2.64	7.3	8.76	8.06	−0.3	N5236
DDO 147	5.31	39	18	0.00	−14.94	25.0	2.84	7.3	8.28	1.29	−1.4	N5194
HIPASS J1247−77	0.82	71	19	0.00	−12.91	23.0	0.20	6.9	7.52	1.44	0.1	I3104
KK 166	0.88	55	...	0.00	−10.82	25.2	...	...	...	...	0.3	N4736
U7990	1.89	47	...	0.00	−13.05	24.6	...	...	...	...	−0.3	N4826

TABLE 4—*Continued*

Name (1)	A_{25} (2)	i (3)	V_m (4)	A_i (5)	M_B (6)	Σ_B (7)	M_{HI}/L (8)	Σ_{HI} (9)	log M_{25} (10)	M_{25}/L (11)	Θ (12)	MD (13)
N4736	14.80	37	164	0.15	−19.83	22.3	0.04	6.4	10.66	3.46	−0.5	N4244
DDO 154	3.55	44	54	0.04	−14.08	25.0	9.42	7.8	9.08	18.05	−0.9	N4736
E443−09	1.33	42	16	0.00	−12.00	24.9	0.78	6.7	7.61	4.16	−0.9	N5236
U8055	1.74	40	62	0.05	−13.45	24.1	2.64	7.6	8.90	21.02	−1.3	N4826
E219−010	2.03	39	...	0.00	−12.70	25.1	...	...	...	...	0.1	N4945
U8061	1.67	38	38	0.00	−13.99	23.4	0.32	7.0	8.45	4.57	−0.8	N4826
N4826	20.41	61	164	0.47	−20.90	22.0	0.02	6.4	10.80	1.80	−1.7	N5194
GR 8	0.67	25	21	0.00	−12.11	23.3	0.74	7.4	7.56	3.33	−1.2	M31
UA 319	3.53	47	66	0.09	−14.51	24.5	0.43	6.6	9.25	18.06	2.1	DDO 161
DDO 161	12.64	87	60	0.37	−16.34	25.5	2.05	6.9	9.72	9.85	1.4	UA 319
E269−37	0.86	42	...	0.00	−12.02	24.0	...	...	...	...	1.6	N4945
KK 182	1.04	55	5	0.00	−11.89	24.5	0.73	6.9	6.48	0.34	1.2	N5128
N4945	17.41	80	174	1.23	−20.51	22.0	0.05	6.7	10.79	2.46	0.7	N5128
I4182	8.17	23	43	0.00	−16.40	24.5	0.61	6.8	9.24	3.06	0.6	N4736
DDO 165	4.20	59	23	0.00	−15.09	24.3	0.82	7.0	8.42	1.56	0.0	N4236
U8215	1.53	47	12	0.00	−12.75	24.5	1.60	7.2	7.44	1.41	−0.5	N5023
P45628	1.12	49	32	0.00	−14.28	22.3	0.83	7.8	8.13	1.66	−0.6	N5068
E269−058	3.12	55	38	0.00	−14.95	23.8	0.07	6.1	8.72	3.52	1.9	N4945
KKs 53	0.88	30	...	0.00	−10.86	25.2	...	...	...	...	1.2	N5128
N5023	7.50	85	81	0.75	−16.68	24.0	0.46	6.9	9.76	7.88	−0.5	N4736
KK 189	0.64	49	...	0.00	−10.52	24.8	...	...	...	...	2.0	N5128
E269−66	1.52	31	...	0.00	−13.56	23.7	...	...	...	...	1.7	N5128
DDO 167	1.20	58	11	0.00	−12.70	24.0	1.00	7.2	7.20	0.85	0.0	DDO 168
DDO 168	3.81	70	28	0.00	−15.28	23.9	1.31	7.4	8.55	1.75	−0.0	N4736
DDO 169	5.20	74	23	0.00	−15.15	24.7	1.34	7.1	8.49	1.75	1.1	N5194
N5068	14.76	28	61	0.02	−19.04	23.1	0.18	6.8	9.81	1.00	−1.4	N4594
KK 195	1.79	65	6	0.00	−11.76	25.8	4.47	7.1	6.84	0.88	−0.2	N5236
KKs 54	1.45	48	...	0.00	−10.47	26.6	...	...	...	...	1.0	N5236
KK 196	1.39	41	...	0.00	−12.00	25.0	...	...	...	...	2.2	N5128
N5102	7.20	75	92	0.47	−18.08	22.5	0.09	6.8	9.85	2.67	0.7	N5128
KK 197	1.03	40	...	0.00	−12.76	23.6	...	...	...	...	3.0	N5128
KKs 55	0.93	29	...	0.00	−9.91	26.2	...	...	...	...	3.1	N5128
KK 198	0.82	35	...	0.00	−10.96	24.9	...	...	...	...	0.8	N5236
KK 200	1.69	53	8	0.00	−11.96	25.5	0.90	6.6	7.07	1.23	1.2	N5236
N5128	28.88	40	398	0.29	−20.77	22.8	0.01	5.6	11.73	16.84	0.6	N4945
I4247	1.52	71	21	0.00	−14.18	23.0	0.20	6.9	7.88	1.02	1.5	N5236
KK 203	0.57	8	...	0.00	−10.22	24.9	...	...	...	...	2.1	N5128
E324−24	3.58	45	47	0.02	−15.45	23.6	0.98	7.4	8.96	3.91	2.4	N5128
P170257	1.14	42	...	0.00	−14.16	22.4	...	...	...	...	0.2	N5068
N5204	6.17	54	60	0.09	−16.75	23.5	0.75	7.3	9.41	3.29	−1.1	M81
N5194	24.44	52	112	0.24	−21.34	21.9	0.06	6.8	10.55	0.67	4.1	N5195
N5195	11.83	24	...	0.00	−19.22	22.5	0.02	6.1	...	...	5.0	N5194
U8508	1.15	55	26	0.00	−12.98	23.6	0.90	7.3	7.96	3.74	−1.0	M81
N5206	4.19	31	...	0.00	−16.66	22.8	...	...	...	...	1.1	N4945
N5229	3.61	84	57	0.31	−14.60	24.5	1.40	7.2	9.13	12.53	−0.6	N4736
N5238	2.50	35	26	0.00	−15.08	23.2	0.17	6.8	8.29	1.16	−0.9	N4736
P47885	0.73	38	79	0.07	−12.98	22.6	3.20	8.3	8.73	21.99	0.4	N5236
E444−78	1.42	68	15	0.00	−13.01	24.1	0.79	7.1	7.57	1.47	2.1	N5236
KK 208	7.14	68	...	0.00	−14.24	26.3	...	...	...	...	1.6	N5236
N5236	17.37	27	211	0.10	−20.43	22.1	0.36	7.5	10.96	3.90	0.8	N5264
DEEP J1337−33	0.27	19	13	0.00	−11.18	22.3	1.08	7.9	6.74	1.18	1.2	N5236
E444−84	1.76	41	41	0.00	−13.56	24.0	2.52	7.6	8.53	8.23	1.7	N5236
HIPASS J1337−39	0.88	35	27	0.00	−12.27	23.8	2.95	7.8	7.88	5.98	−0.3	N5236
N5237	2.09	34	64	0.04	−15.00	22.9	0.09	6.6	9.00	6.40	2.1	N5128
U8638	0.75	49	17	0.00	−12.43	23.3	0.33	7.0	7.39	1.69	−1.3	M31
DDO 181	1.81	57	20	0.00	−12.97	24.6	1.01	7.0	7.91	3.38	−1.3	M81
N5253	5.04	69	38	0.00	−17.38	22.4	0.14	7.0	8.92	0.59	0.5	N5236
I4316	1.97	52	31	0.00	−13.90	23.9	0.17	6.5	8.33	3.78	2.4	N5236
N5264	3.11	55	21	0.00	−15.90	22.9	0.17	6.9	8.22	0.47	2.6	N5236
KKs 57	0.88	30	...	0.00	−10.07	26.0	...	...	...	...	1.8	N5128
KK 211	0.89	30	...	0.00	−11.93	24.1	...	...	...	...	1.5	N5128
KK 213	0.60	62	...	0.00	−9.72	25.5	...	...	...	...	1.7	N5128
E325−11	2.49	64	33	0.00	−14.05	24.2	1.67	7.3	8.49	4.78	1.1	N5128
KKs 58	1.03	47	...	0.00	−10.64	25.7	...	...	...	...	0.6	N5128
KK 217	0.72	35	...	0.00	−10.87	24.7	...	...	...	...	1.1	N5128

TABLE 4—*Continued*

Name (1)	A_{25} (2)	i (3)	V_m (4)	A_i (5)	M_B (6)	Σ_B (7)	M_{HI}/L (8)	Σ_{HI} (9)	log M_{25} (10)	M_{25}/L (11)	Θ (12)	MD (13)
KK 218	1.32	66	...	0.00	−10.97	25.9	...	...	...	...	1.6	N5236
E174−01	6.88	67	28	0.00	−16.82	23.7	0.49	7.0	8.81	0.77	−1.5	N5128
Cen N	1.02	40	...	0.00	−10.89	25.5	...	...	...	...	0.9	N5128
HIPASS J1348−37	0.54	38	28	0.00	−11.21	23.7	0.64	7.1	7.68	10.12	0.8	N5128
KK221	1.82	49	...	0.00	−10.60	27.0	...	...	...	...	0.6	N5128
E383−87	2.00	37	21	0.00	−15.16	22.7	0.10	6.8	8.01	0.57	−0.8	Milky Way
DDO 183	2.64	75	12	0.00	−14.16	24.3	0.83	7.0	7.64	0.61	−0.8	N4736
HIPASS J1351−47	0.54	55	18	0.00	−10.90	24.1	2.95	7.7	7.29	5.43	0.9	N5128
KKH 86	0.51	46	5	0.00	−10.32	24.5	0.38	6.6	6.17	0.71	−1.5	M31
U8837	8.26	77	35	0.00	−15.90	25.0	0.91	6.8	9.07	3.32	−0.0	M101
U8833	0.82	28	22	0.00	−12.42	23.5	1.00	7.4	7.68	3.30	−1.4	N4736
E384−016	1.40	48	...	0.00	−13.06	24.0	...	...	...	...	0.3	N5128
M101	61.44	22	173	0.05	−21.23	24.0	0.47	6.9	11.33	4.39	0.6	N5474
N5408	2.38	68	33	0.00	−16.50	21.7	0.56	7.9	8.47	0.47	−0.5	N5236
N5474	9.95	26	39	0.00	−17.74	23.6	0.69	7.2	9.24	0.90	2.0	M101
N5477	3.64	41	35	0.00	−15.29	23.8	0.90	7.2	8.70	2.48	1.4	M101
KK 230	0.32	35	9	0.00	−8.55	25.3	5.30	7.4	6.50	7.74	−1.0	M31
Circinus	11.34	67	151	0.56	−19.90	21.7	0.25	7.5	10.48	2.11	−0.7	N4945
KKH 87	1.43	71	17	0.00	−13.30	23.8	1.34	7.4	7.67	1.45	1.0	M101
DDO 187	1.19	42	21	0.00	−12.71	24.0	0.93	7.2	7.78	3.20	−1.3	M31
N5585	8.92	51	83	0.16	−17.82	23.2	0.46	7.2	9.86	3.45	−0.8	M101
DDO 190	1.44	28	42	0.00	−14.18	22.9	0.59	7.4	8.47	4.03	−1.3	M81
P51659	2.31	71	16	0.00	−11.83	26.3	6.08	7.1	7.86	8.58	0.1	N5128
E222−10	1.62	66	31	0.00	−13.60	23.8	0.79	7.2	8.27	4.35	−1.4	N5128
DDO 194	3.24	73	39	0.00	−15.05	23.8	0.61	7.1	8.77	3.59	−0.0	M101
E272−25	5.17	61	38	0.00	−14.77	25.1	0.19	6.1	8.94	6.90	−1.5	N5128
E223−09	5.07	24	70	0.02	−15.98	23.8	2.24	7.6	9.46	7.43	−0.8	E274−01
UMin	0.49	52	...	0.00	−7.13	27.7	...	...	...	...	3.3	Milky Way
E274−01	14.55	86	80	0.77	−18.53	23.6	0.20	6.7	10.04	2.70	−1.0	N5128
KKR 25	0.54	55	5	0.00	−9.94	25.0	0.84	6.7	6.20	1.07	−0.7	M31
E137−18	5.41	73	68	0.26	−17.93	22.0	0.12	7.1	9.46	1.24	−1.8	N5128
Draco	0.79	48	...	0.00	−8.74	27.1	...	...	...	...	3.0	Milky Way
Milky Way	25.00	...	220	0.00	−20.80	...	0.14	7.0	11.15	4.33	2.5	LMC
I4662	1.55	57	53	0.06	−15.13	22.1	0.67	7.8	8.70	2.87	−0.9	Circinus
N6503	9.03	71	83	0.38	−18.08	23.0	0.39	7.2	9.86	2.71	−1.2	N6946
Sag dSph	2.60	64	...	0.00	−12.67	25.7	...	...	...	...	5.6	Milky Way
N6789	1.48	41	...	0.00	−14.32	22.8	...	...	...	...	−1.4	N6946
Sag DIG	0.92	45	18	0.00	−11.49	24.6	1.04	7.0	7.54	5.62	−0.3	Milky Way
N6822	2.71	30	55	0.02	−15.22	23.3	0.73	7.4	8.99	5.06	0.6	Milky Way
KK 246	2.14	68	25	0.00	−12.96	25.0	2.57	7.2	8.20	6.63	−2.2	N253
U11583	3.45	66	43	0.01	−14.28	24.7	2.03	7.2	8.88	9.39	3.3	N6946
KK 251	3.12	62	24	0.00	−13.63	25.1	2.43	7.1	8.33	4.81	3.4	N6946
KK 252	2.31	8	72	0.00	−14.11	24.0	0.17	6.4	9.14	20.37	3.0	N6946
N6946	25.96	32	144	0.10	−20.86	22.5	0.21	7.1	10.80	1.82	0.7	U11583
KKR 55	1.75	49	20	0.00	−14.79	22.7	0.34	7.3	7.93	0.66	2.3	N6946
DDO 210	0.55	62	6	0.00	−11.09	23.9	0.55	7.0	6.35	0.53	−0.1	M31
KKR 56	2.11	52	...	0.00	−14.39	23.5	0.52	7.1	...	...	1.6	N6946
Cepheus 1	10.38	62	45	0.02	−17.53	23.9	0.69	7.1	9.39	1.54	0.9	N6946
I5052	7.64	83	83	0.69	−18.23	22.5	0.32	7.3	9.79	2.00	−2.2	N253
KKR 59	7.96	54	34	0.00	−17.01	23.8	0.30	6.8	9.04	1.11	1.9	KKR 60
KKR 60	2.88	47	...	0.00	−15.42	23.2	...	...	...	...	2.5	KKR 59
I5152	2.91	53	50	0.04	−15.67	23.0	0.34	7.2	8.93	2.95	−1.1	M31
KK 258	1.60	62	...	0.00	−10.66	26.7	...	...	...	...	−0.9	N253
Tucana	0.64	69	...	0.00	−9.16	26.2	...	...	...	...	−0.1	Milky Way
UA438	0.94	38	23	0.00	−12.94	23.2	0.76	7.4	7.75	2.39	−0.7	N55
Cas dSph	0.65	38	...	0.00	−11.67	23.7	...	...	...	...	2.0	M31
Pegasus	1.04	59	8	0.00	−11.47	24.9	0.43	6.5	6.90	1.30	1.2	M31
UA 442	5.44	86	41	0.00	−14.64	25.3	1.26	6.8	9.03	9.60	−0.1	N7793
KKH 98	0.78	58	7	0.00	−10.78	25.0	1.80	7.1	6.67	1.47	−0.7	M31
Peg dSph	0.88	62	...	0.00	−10.80	25.2	...	...	...	...	1.7	M31
E149−003	2.94	90	24	0.00	−14.09	24.6	0.84	6.9	8.30	2.98	−1.7	N253
N7793	9.96	47	107	0.19	−18.53	22.8	0.26	7.1	10.12	3.28	0.1	N253

s^{-1} (the correction for turbulent motions for some dwarf galaxies with $W_{50} \leq 20$ km s^{-1} leads to negative values of V_m, and we adopt for them a formal value $V_m = 5$ km s^{-1}).

5. Internal absorption in the galaxy in the B band. According to Giovanelli et al. (1994) and Tully et al. (1998), the internal absorption, A_i, depends not only on inclination, but also on galaxy luminosity. Based on the data by Verheijen (2001) and Karachentsev, Karachentseva, & Kudrya (1999b), we determined the internal absorption as

$$A_i = [1.6 + 2.8(\log V_m - 2.2)] \log (a/b), \qquad (5)$$

if $V_m > 42.7$ km s^{-1}, otherwise, $A_i = 0$ (so, we considered dwarf galaxies to be fully transparent). For the gas-poor E, S0, and dSph galaxies without estimates of W_{50}, the A_i correction was also ignored.

6. Absolute magnitude M_b corrected for A_g and A_i.

7. Mean surface brightness, $\Sigma_B = (B_t - A_b - A_i) + 5 \log a_c + 8.63$, in magnitudes per square arcsecond.

8. H I mass-to-luminosity ratio in solar units where

$$\log (M_{\rm HI}/M_\odot) = \log F + 2 \log D + 5.37, \qquad (6)$$

(Roberts & Haynes 1994) and D is the galaxy distance in Mpc.

9. Logarithm of surface density of neutral hydrogen in $M_\odot$ kpc^{-2}, $\Sigma_{\rm HI} = 4M_{\rm HI}/\pi A_{25}^2$.

10. Logarithm of "indicative" mass of galaxy within its standard diameter A_{25}

$$\log (M_{25}/M_\odot) = 2 \log V_m + \log a_c + \log D + 4.52, \quad (7)$$

where V_m expressed in km s^{-1}, a_c in arcmin, and D in Mpc (Roberts & Haynes 1994).

11. Indicative mass-to-luminosity ratio in solar units.

12. Tidal index Θ, determined below in § 6. Negative values of Θ correspond to isolated galaxies, and positive values are typical of group members.

13. Name of a neighboring galaxy ("main disturber" = MD) producing the maximum tidal action on the given galaxy. In fact, a collection of galaxies with one and the same main disturber, and also the positive Θ, may be considered as a separate group (see also notes to Table 1).

4. ON THE SAMPLE COMPLETENESS

We hope that a more or less strict estimate of completeness of the catalog may be achieved in the nearest future, when precise distances are measured for the majority of galaxies within, say, 8 Mpc. At present, we restrict ourself to a rough estimate of completeness of the sample under consideration.

Figure 2 (bottom) presents the distribution of nearby galaxies according to their distances and absolute magnitudes corrected for the external and internal extinction. The dotted and dashed lines correspond to limiting apparent magnitudes $B_t = 15.5$ mag (the limit of CGCG; Zwicky et al. 1961) and 17.5 mag (an approximate limit for Karachentsevs' survey), respectively. The solid line shows a running median that increases from -12 mag within the Local Group to -14.5 mag at $D \sim 8$ Mpc. The behavior of the median indicates lack of the faintest galaxies toward the LV edge. Assuming our sample within $D = 2$ Mpc to be complete to nearly 100%, we derive an estimate of completeness within 8 Mpc to be roughly 70%–80%. Therefore, new more careful searches for faint neighboring galaxies may lead to discovery of $\sim$100 new galaxies within 8 Mpc around us.

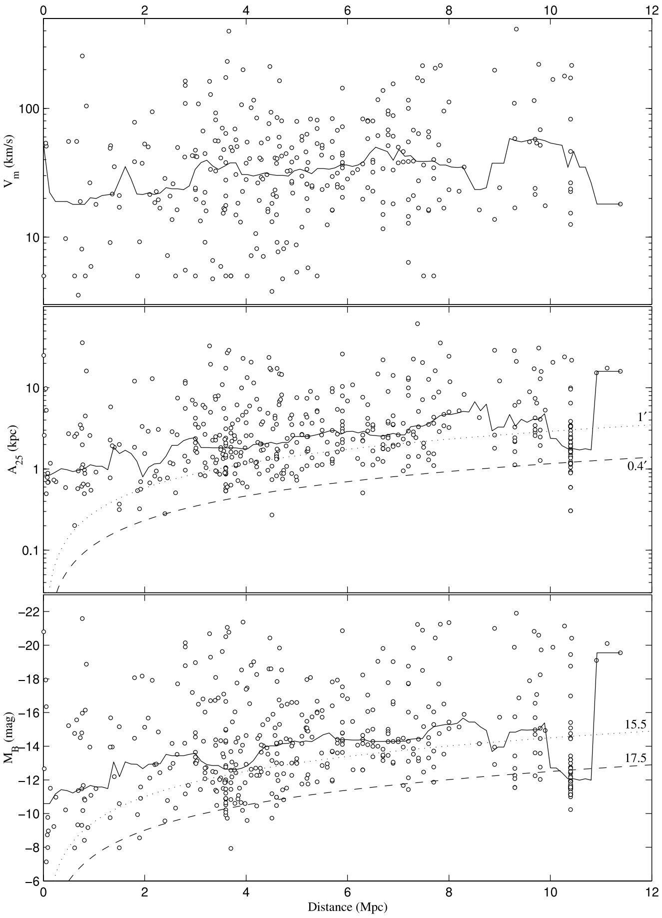

FIG. 2.—Distribution of the neighboring galaxies according to their absolute magnitudes (*bottom*), linear diameters (*middle*), and rotational velocities (*top*) vs. distances. The solid line corresponds to the median with a running window of 1 Mpc. The dotted and dashed lines on the bottom panel show limiting apparent magnitudes $B_t = 15.5$ and 17.5 mag for CGCG and Karachentsev's survey, respectively. The dotted and dashed lines on the middle panel correspond to limiting angular diameter 1′ and 0.′4 characteristic of Nilson's catalog (UGC) and Karachentsev's survey, respectively.

Apart from the absolute magnitude of a galaxy, its linear diameter A_{25} or rotation amplitude V_m may be treated as an indicator of its position on the scale "dwarf–giant". Distributions of neighboring galaxies versus these parameters are presented on the other two panels of Figure 2. The trend of the medians $A_{25}(D)$ and $V_m(D)$ with distance exhibit a slightly lower effect of incompleteness than the trend of luminosity. As it was noted by Cross & Driver (2002) and many others, the selectivity function depends not only on galaxy luminosity (diameter, rotation velocity), but also on its optical surface brightness. The absolute magnitude, the mean surface brightness relationship for 451 neighboring galaxies, is plotted in Figure 3. The mean surface brightnesses spreads over a Q5
range of 7 magnitudes. Evidently, the galaxies of very low surface brightness can easily be lost when situated in the Zone of Avoidance. Bothun, Impey, & McGaugh (1997), McGaugh & de Blok (1997), and O'Neil & Bothun (2000) supposed that the true distribution of galaxies according to their mean surface brightness extends far beyond $\Sigma_B \sim 27$ mag arcsec^{-2}, and, probably, about 80% of all galaxies lie for us beyond the present threshold of detection. Figure 3 shows that the mean surface brightness decreases from giant galaxies to dwarfs. Such tendency is expected when giant and dwarf galaxies have approximately one and the same mean volume density of

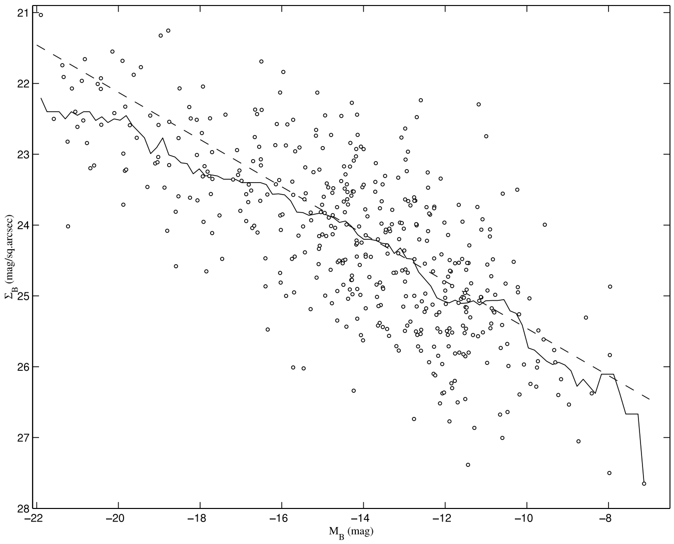

FIG. 3.—Mean surface brightness of 451 nearby galaxies versus their absolute magnitudes in the B band. The median with a running window of 1.5 mag is shown by the solid line. The expected relation $\Sigma_B \sim \frac{1}{3} M_B$ for galaxies with the constant luminosity density is traced by the dashed line.

stars, that corresponds to the relation $\Sigma_B \sim \frac{1}{3} M_B$ shown in Figure 3 by the dashed line.

5. GALAXY DISTRIBUTION WITHIN 8 Mpc

The sky distribution of 451 nearby galaxies is presented in Figure 4 in equatorial coordinates. The distribution looks extremely inhomogeneous, showing two large empty areas: the Local Void (Tully 1988) in Hercules-Aquila and the Local Mini-void (Karachentsev et al. 2002c) in the Orion constellation. Spatial distribution of the galaxies is seen in Figure 5 in the supergalactic coordinates. Galaxies with the distances less and more than 8 Mpc are shown by large and small circles, respectively. One can distinguish in Figure 5*a* (the supergalactic plane projection) some relatively compact groups around the Milky Way, M31, M81, IC 342, Cen A, M83, and also the Canes Venatici cloud. Remarkably, in the huge volume of the Tully void (~100 Mpc^3), there is not any galaxy with luminosity brighter than $L \sim 2 \times 10^6\ L_\odot$. The dwarf galaxy KK 246 (-12.96 mag) with the coordinates R.A. = $20^{\rm h}03^{\rm m}57^{\rm s}$, decl. = $-31°40'54''$ is situated just at the edge of the Tully void, but not inside it. The Local Supercluster center in Virgo (D = 17 Mpc, SGL = 104°, SGB = −4°) is characterized by the Cartesian coordinates: SGX = −4 Mpc, SGY = 16 Mpc, SGZ = −1 Mpc. A small density gradient is seen toward Virgo, but it is masked by strong density fluctuations caused by voids. The majority of the groups locate in a thin layer $|\,\mathrm{SGZ}\,| < 0.3$ Mpc of the Supergalactic plane. However, there are also groups around NGC 6946, M101, M96 (Leo I), situated at a distance of ~3–4 Mpc from the Local "pancake."

6. THE BASIC OPTICAL PROPERTIES OF THE LOCAL VOLUME SAMPLE

In space of parameters $\{M_B, A_{25}, V_m\}$, corrected for the Galactic extinction and inclination, the galaxies are concentrated toward the "fundamental surface" whose projections onto planes $\{M_B, A_{25}\}$, $\{M_B, V_m\}$ and $\{A_{25}, V_m\}$ are shown in Figure 6. The projections of these distributions onto the axes Q6 M_B, A_{25}, and V_m gives us a luminosity function $\varphi(M)$, as well as analogous functions, $\Psi(A_{25})$ and $\xi(V_m)$. Each of them can be represented by the Schechter (1976) function with three arbitrary parameters. It should be noted, however, that the rotational velocity V_m for 138 nearby galaxies out of 451 remains unknown because they are gas-poor systems undetectable in the H I line. This introduces additional selectivity when analyzing observational data.

It is well known that the morphological classifications of galaxies by de Vaucouleurs et al. (1976) and van den Bergh (1966) are luminosity-dependent. Figure 7 (*bottom*) presents a correlation between morphological class and absolute magnitude for the neighboring galaxies. The solid line,

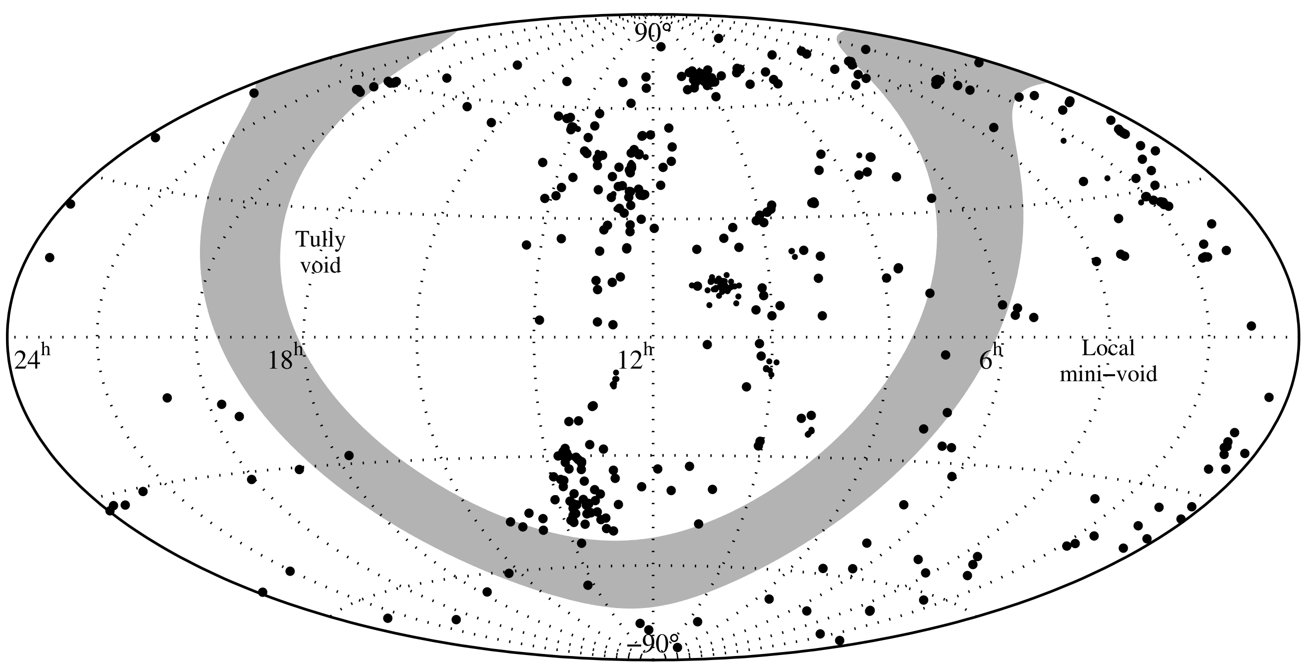

FIG. 4.—All-sky distribution of 451 neighboring galaxies from Table 1 in equatorial coordinates. The galaxies with distances $D > 8$ Mpc are shown as small circles. The shaded area marks the zone of avoidance in the Milky Way.

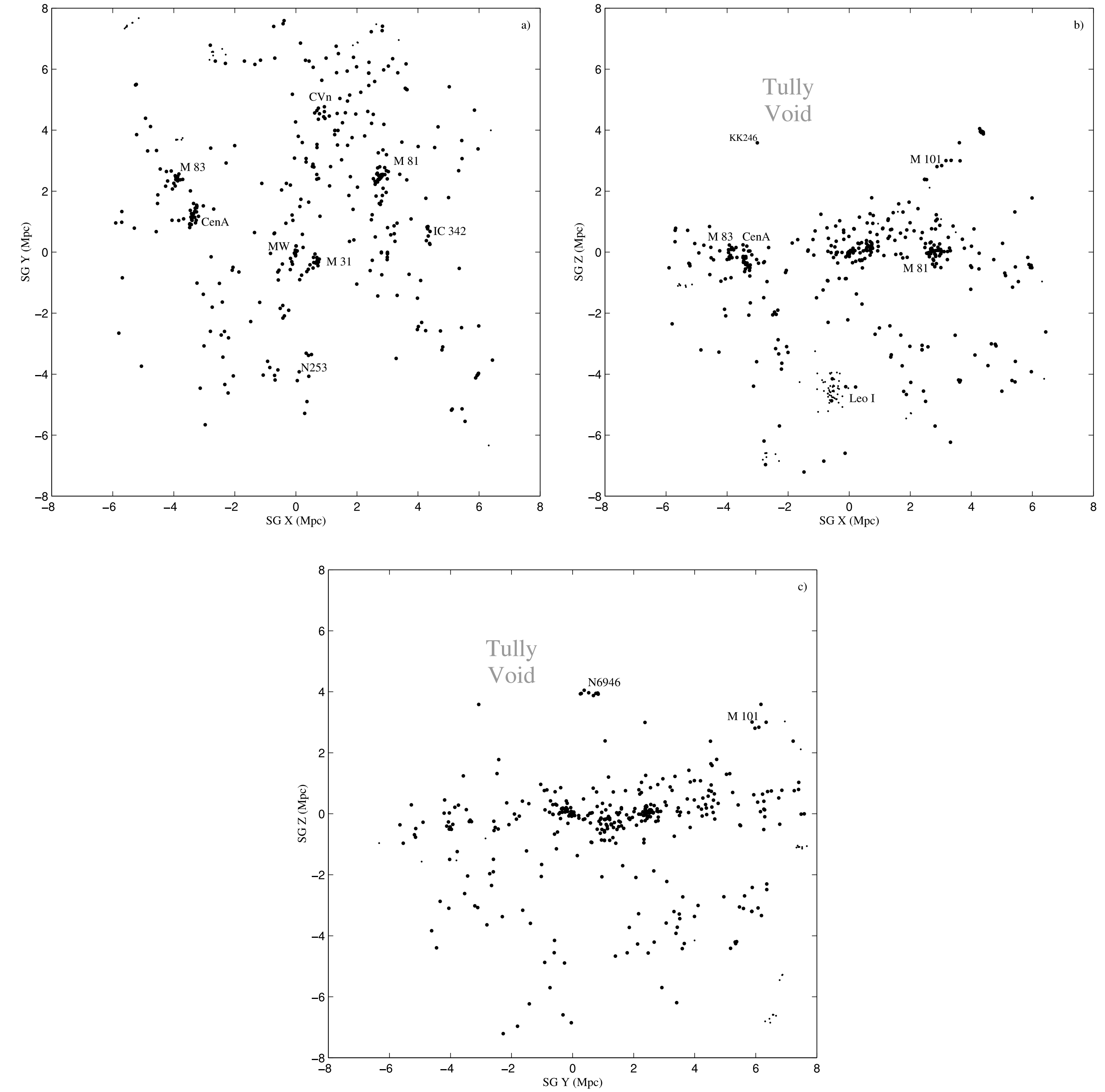

Fig. 5.—Panorama of the LV within a radius of 8 Mpc in Cartesian supergalactic coordinates. Galaxies from Table 1 with $D > 8$ Mpc are shown as small circles. (*a*) SGX-SGY, galaxies projected onto the plane of the Local Supercluster; (*b*) SGX-SGZ, the distribution in Z (perpendicular to the plane of the Local Supercluster) is dominated by the galaxies concentrated toward this plane. The Local Void in the top part of the diagram is evident; (*c*) SGY—SGZ, see the asymmetric distribution in the plane of the Local Supercluster toward Virgo. Some known groups are marked by names of their brightest members.

showing the median magnitude for each morphological class, reflects a decrease in luminosity from the Sa–Sbc types ($T = 2$–5) to both later and earlier types. Note that in the distance-limited sample about 70% of the galaxies belong only to two latest types, $T = 10$ and 9 (dIr, dIm). Another disadvantage of the existing morphological classifications relates to dwarf galaxies of ''transient'' type, dSph/dIr. Usually, dwarf spheroidal galaxies are attributed to the same types ($T < 0$) as giant elliptical galaxies. Therefore, slight misclassification of a dSph/dIr galaxy can throw it from one side of the Hubble diagram to the opposite side.

Figure 7 (*top*) presents a distribution of nearby galaxies according to their absolute magnitude and type of the circumnuclear region. As can be seen, the cases of apparent stellar-like nuclei (''K'' = kern) and the bulgelike central regions (''B''), as well as intermediate cases (KB, BA), occur exceptionally in galaxies of high luminosity (both spiral and elliptical). Among dwarf galaxies fainter than -15.7 mag, there are only systems without a nucleus (''A'' = absent), irrespective of their location in groups or the general field. One exception to the rule is PGC 47885 (-12.96 mag). However, this may be a distant galaxy with an erroneous radial

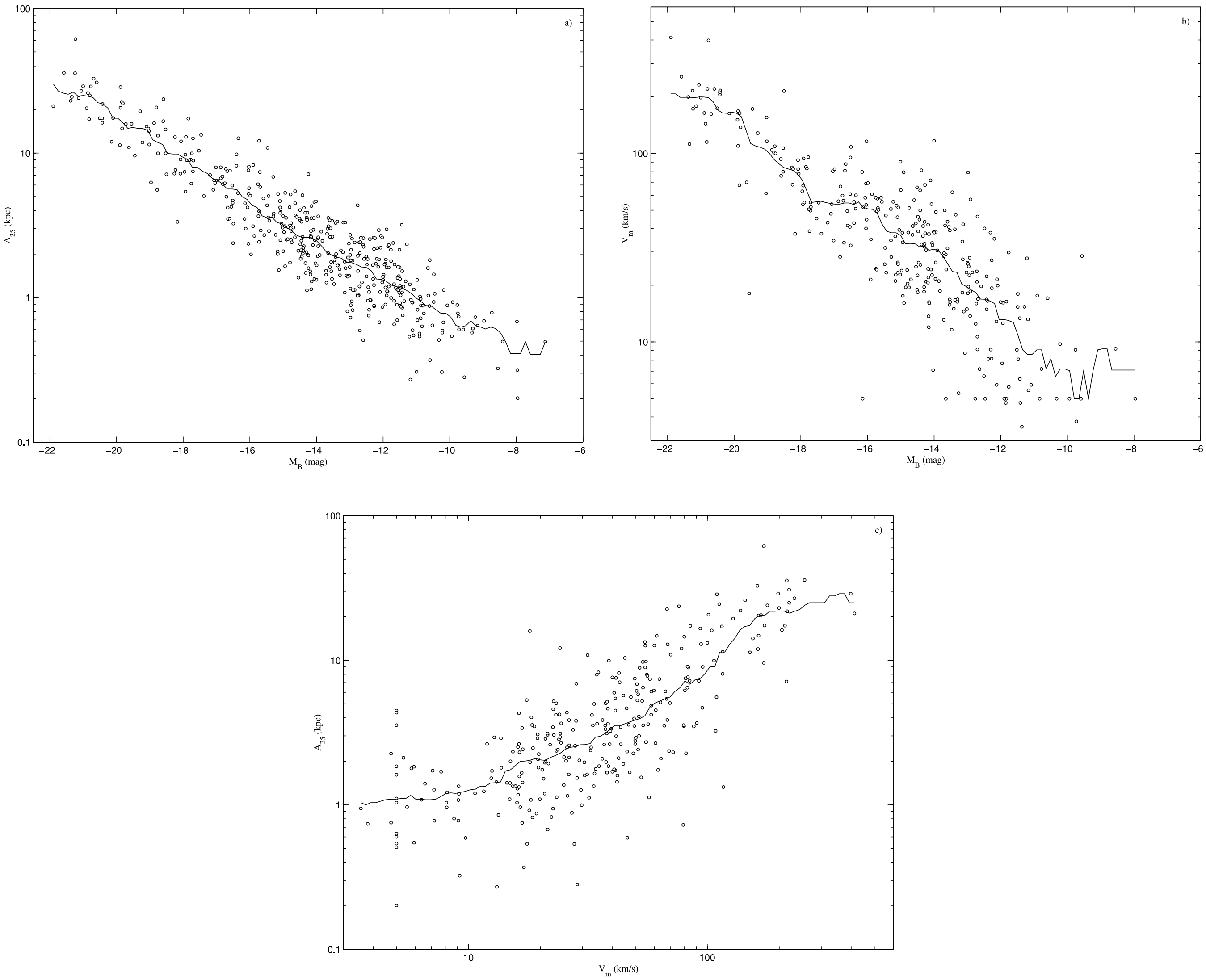

FIG. 6.—Distribution of the LV galaxies according to their (*a*) absolute magnitudes and linear diameters, (*b*) absolute magnitudes and rotational velocities, and (*c*) rotational velocities and linear diameters. The running medians are shown by the solid lines.

velocity because of H I confusion with the bright galaxy M83 situated 30′ southeast. On the other hand, there are some rather bright galaxies: NGC 925 (−19.88 mag) and NGC 2403 (−19.29 mag) without apparent stellar-like nuclei. The absence of "nucleated" systems among galaxies with luminosity below a fixed threshold may give the key to understanding of formation of black holes in galaxy cores.

If the scheme of corrections of L, A_{25}, and V_m for extinction and inclination was right, we would expect to see the total mass-to-luminosity ratio and hydrogen mass-to-luminosity ratio to be uncorrelated with angle i. Figure 8 shows that both the ratios, M_{25}/L and $M_{\rm HI}/L$, are practically independent of galaxy inclination. As it also follows from Figure 8, individual values of M_{25}/L for the neighboring galaxies are distributed in a rather wide range from 0.03 to 34 $M_\odot/L_\odot$ with the median 3.1 $M_\odot/L_\odot$. Only a small part of this scatter is caused by measurement errors of V_m, A_{25}, L, and i, affecting the mass-to-luminosity ratios.

Figure 9*a*–9*e* present correlation of M_{25}/L with different galaxy parameters. The bottom panel shows that the mean total mass-to-blue luminosity ratio does not depend, in fact, on the morphological type. The dispersion of M_{25}/L is minimal for early-type galaxies ($T = 0$–4), increasing appreciably toward irregular galaxies. The last circumstance may be caused by the uncertainty in inclination correction for objects of irregular shape. As seen from other panels of Figure 9, the mean mass-to-luminosity ratio reveals a slight correlation with the mean surface brightness and other galaxy parameters. Obviously, the correlations of M_{25}/L with V_m, A_{25}, and L have partially a statistical cause, which is the more essential the higher random errors of V_m, A_{25}, and L are. Nevertheless, the data in Figure 9 gives no support to the widespread opinion (Mateo 1998) that the total mass-to-luminosity ratio increases steadily from giant galaxies toward dwarf ones. As a zero-hypothesis, we adhere to the assumption that the mean M_{25}/L actually does not depend on galaxy luminosity, and individual variations of M_{25}/L from galaxy to galaxy reflect different amount of dark matter per unit galaxy luminosity.

Curiously, most of the galaxies having the highest ratios, $M_{25}/L > 10$, are concentrated within a narrow range of absolute magnitudes, [−12, −16 mag], and rotational velocities, [30–100] km s^{-1}. All of them are rather isolated objects. The nature of this peculiar population of dwarf galaxies seems to us enigmatic.

7. THE BASIC HYDROGEN PROPERTIES OF THE SAMPLE

A detailed survey of H I data on several thousand galaxies from Nilson (1974) catalog has been presented by Roberts & Haynes (1994). However, being angular diameter-limited but not distance-limited, their sample is strongly biased toward giant galaxies. H I properties of a distance-limited sample have been considered by Huchtmeier & Richter (1988) and Karachentsev et al. (1999b). Here we briefly discuss the basic hydrogen characteristics of the most complete sample of neighboring galaxies.

Figure 10*a* presents the H I surface density within the standard optical diameter (in $M_\odot$ kpc^{-2}) as a function of rotational velocity V_m. As distinct from others, the early-type galaxies ($T < 3.5$) with predominant bulges are shown by large circles. In its top part, the distribution is restricted by $\Sigma_{\rm HI} = 1 \times 10^8$ $M_\odot$ kpc^{-2}. Two galaxies situated above this limit, KKSG 9 and PGC 47885, are probable cases of H I confusion with nearby brighter objects. The median, plotted for disklike galaxies, remains practically the same over the whole range of rotational velocities. The constancy of $\Sigma_{\rm HI}$ for both small and large disks of galaxies has been **noted** by Roberts (1969).

Comparison of the $M_{\rm HI}/L$ and $M_{\rm HI}/M_{25}$ ratios with V_m is shown in Figures 10*b* and 10*c*, respectively. The median hydrogen mass-to-luminosity ratio is $\sim$1 $M_\odot/L_\odot$ for dwarf galaxies and decreases slowly to $\sim$0.1 $M_\odot/L_\odot$ for the most fast rotating objects. Therefore, many dwarf systems have a large amount of gas for further star formation. Carignan & Beaulieu (1989) noted that the dwarf galaxy DDO 154 in the Canes Venatici cloud with $M_{\rm HI}/L = 9.4$ $M_\odot/L_\odot$ is enveloped in a ''H I coat'' extending over 6 optical diameters of the galaxy. In the LV, there are other galaxies (Ho IX, KKs 40, PGC 51659, KKSG 19, KK 191, KK 230, UA 292) comparable with DDO 154 in their $M_{\rm HI}/L$. All of them are worthy of further study in the H I line with high angular resolution. As it follows from Figure 10*c*, the hydrogen mass–to–indicative mass ratio decreases smoothly from dwarf galaxies toward giant ones. The average fraction of hydrogen mass drops by 2 orders when V_m increases by almost 2 orders. The median $M_{\rm HI}/M_{25}$ ratio for the entire sample is $\sim$0.3. As shown by Broeils (1992) and Swaters et al. (2002), for many dwarf galaxies their flat rotation curve extends to $(2–3)a_{25}$. Consequently, the total mass of these galaxies exceeds their indicative mass 2–3 times. In the left corner of the $M_{\rm HI}/M_{25}$ versus V_m diagram, one can see cases where the $M_{\rm HI}/M_{25}$ ratio reaches a value about 5. Apparently, an essential fraction of the total mass of the galaxies are concentrated in gaseous disks. It should be noted, however, that the correction for turbulent motions for many dIr galaxies with the observed H I line $W_{50} \leq 20$ km s^{-1} leads to negative values of V_m. Therefore, the derived estimates of M_{25} and $M_{\rm HI}/M_{25}$ for them are rather uncertain. To such infradwarf objects, we ascribed formally one and the same value, $V_m = 5$ km s^{-1}. Detailed study of the velocity field for the galaxies with $M_{\rm HI}/M_{25} > 1$ would give important information Q7

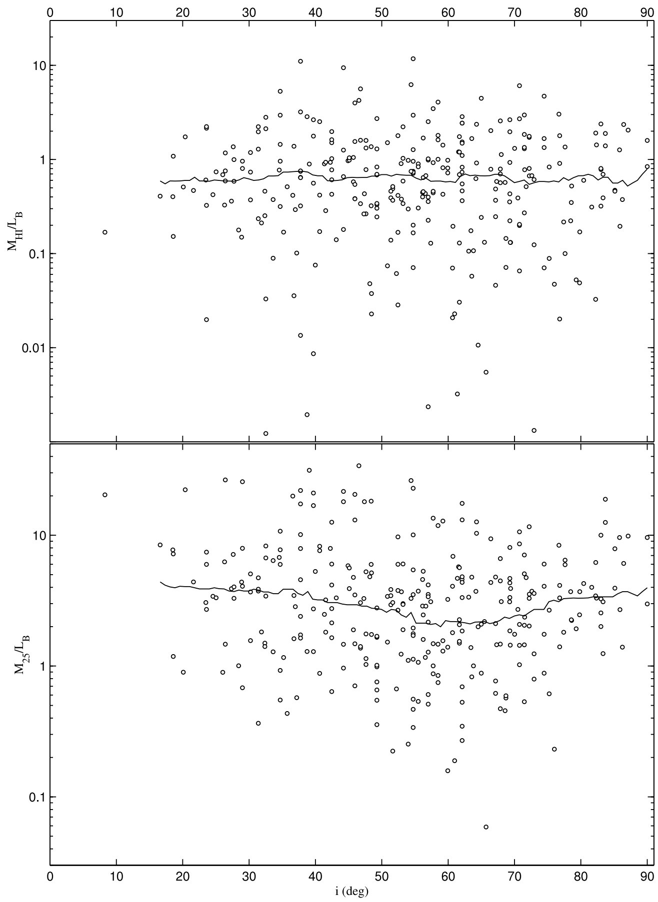

FIG. 8.—Indicative mass-to-luminosity ratio (*bottom*) and H I mass-to-luminosity ratio (*top*) for the LV galaxies vs. their inclination angle. The running medians (window = 20°) are shown by the solid lines.

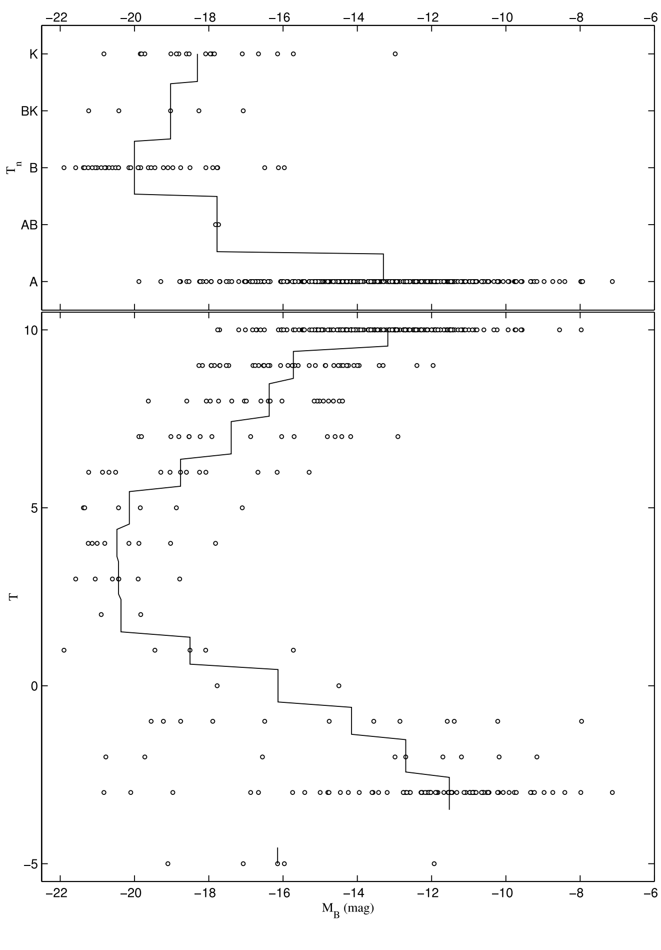

FIG. 7.—Morphological type vs. absolute magnitude (*bottom*) and circumnuclear region type vs. absolute magnitude (*top*) for 451 neighboring galaxies. The solid lines show the running medians.

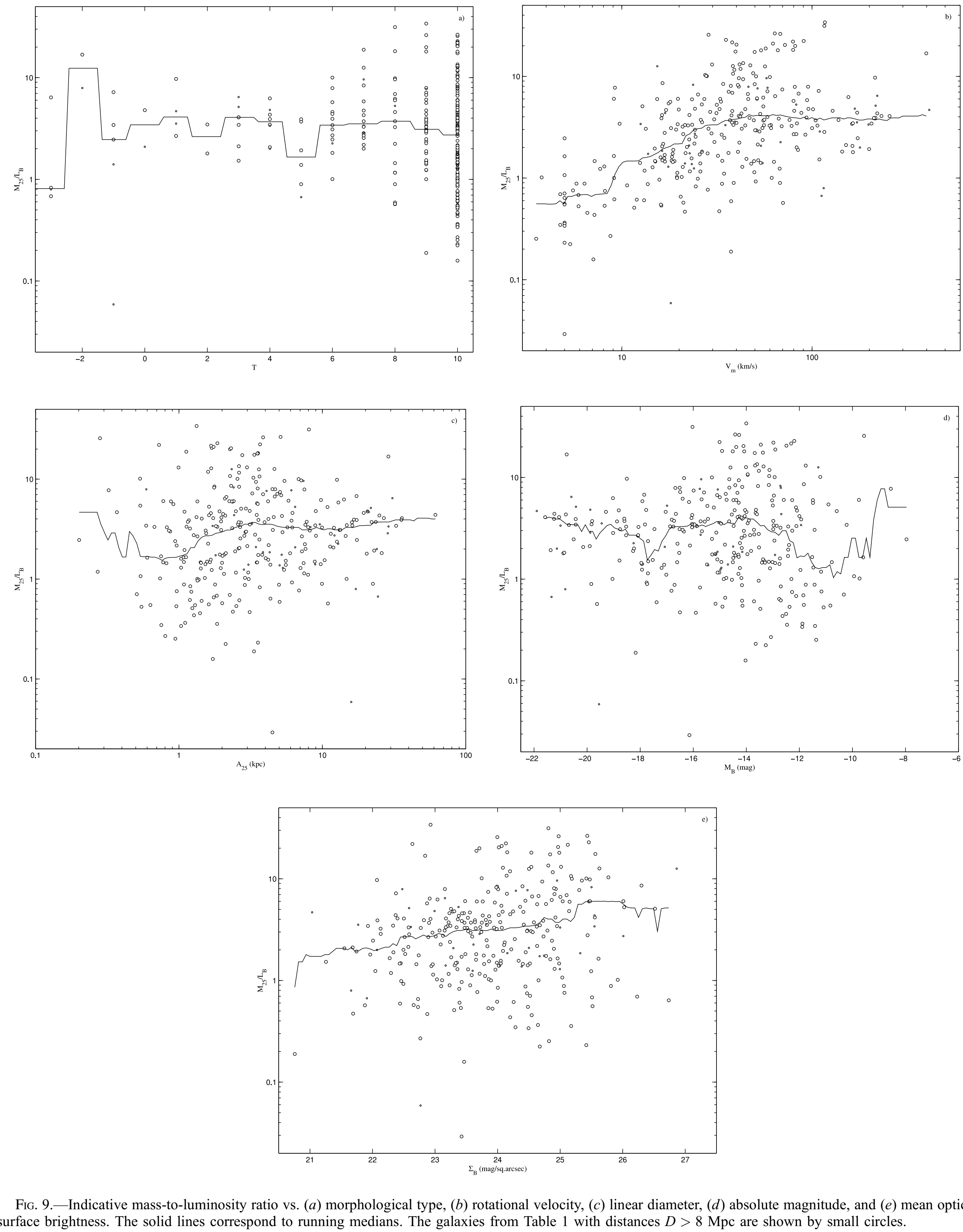

FIG. 9.—Indicative mass-to-luminosity ratio vs. (*a*) morphological type, (*b*) rotational velocity, (*c*) linear diameter, (*d*) absolute magnitude, and (*e*) mean optical surface brightness. The solid lines correspond to running medians. The galaxies from Table 1 with distances $D > 8$ Mpc are shown by small circles.

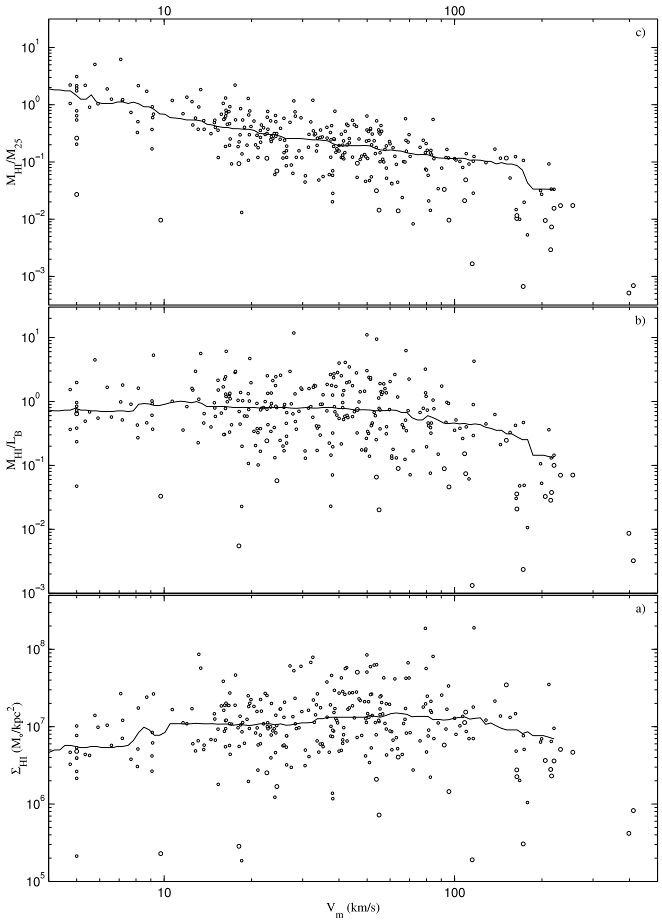

FIG. 10.—(*a*) H I surface density within the standard optical diameter, (*b*) the H I mass-to-luminosity ratio, and (*c*) the H I mass-to-indicative mass ratio as a function of rotational velocity for the neighboring galaxies. The bulgelike ($T < 3.5$) galaxies are shown by large circles. The running medians (*solid lines*) correspond to disklike objects only.

about the distribution of dark and luminous matter on scales of 0.1–1 kpc.

According to Zasov (1974) and Zasov & Rubtsova (1989), active star formation occurs in gaseous disks of galaxies situated just above the threshold of gravitational instability. This condition corresponds to the linear proportionality between the total mass of gas in a disk and its angular momentum: $M_{\rm HI} \propto V_m A_{25}$. The relation between these quantities for 282 nearest disklike ($T > 3.5$) galaxies is presented in Figure 11, which demonstrates that the running median (*solid line*) follows well Zasov's relation (*straight dashed line*) in the range of angular momentum covering four orders. Whether the empirical relation $M_{\rm HI}/V_m A_{25} =$ const remains valid or not if the structure and kinematics of gaseous disks are driven by the process of consecutive mergings of galaxies is open to question.

As it was shown by Dressler (1984), Giovanelli & Haynes (1991), and many others, the H I abundance in disklike galaxies depends on their environment. Spiral galaxies in the cores of rich clusters demonstrate significant H I deficiency with respect to the field galaxies of the same morphological type. However, observational data on the H I deficiency outside rich clusters look rather controversial. To describe the local mass density around a galaxy "*i*," Karachentsev & Makarov (1998) have introduced the so-called tidal index:

$$\Theta_i = \max\,[\log\,(M_k/D_{ik}^3)] + C, i = 1, 2 \cdots N, \qquad (8)$$

where M_k is the total mass of any neighboring galaxy separated from the considered galaxy by a space distance D_{ik}. For every galaxy i, we found its main disturber, producing the highest tidal action or a maximum density enhancement, $\Delta\rho_k \sim M_k/D_{ik}^3$. In order to take account of all surrounding galaxies (but not only ones with measured V_m), we determined the total mass of every galaxy from its luminosity. Based on the data of Figure 9*c*, we accept the mean luminosity weighted value $M_{25}/L = 3.8\ M_\odot/L_\odot$ for all morphological types and also adopted that the total mass of each galaxy is on average 2.5 times its indicative mass, $M_T = 2.5 M_{25}$. The value of the constant C is chosen so that $\Theta = 0$ when the Keplerian cyclic period of the galaxy with respect to its main disturber equals the cosmic Hubble time, $1/H$. In this sense, galaxies with $\Theta < 0$ may be considered as undisturbed (isolated) objects. For instance, the value $\Theta = 0$ is characteristic of NGC 2403, having the highest tidal action from M81. The most disturbed objects in the LV are NGC 221 ($\Theta = 6.8$) and Sag dSph ($\Theta = 5.6$), which have strong gravitational disturbance from M31 and Milky Way, respectively. Data on the tidal index for each galaxy, as well as the name of its main disturber (MD), are given in the last two columns of Table 4. It should be noted here that galaxies at the periphery of the LV have a systematically lower index Θ because we ignored many objects with $D > 10$ Mpc. To characterize the H I deficiency of galaxies in groups, we considered two parameters: the morphological type and the H I surface brightness, $\Sigma_{\rm HI} = M_{\rm HI}/A_{25}^2$, having lower dispersion with respect to $M_{\rm HI}/L$ and $M_{\rm HI}/M_{25}$. The relations T versus Θ and $\Sigma_{\rm HI}$ versus Θ are plotted on the top and bottom panels of Figure 12, respectively. The bulge-dominated galaxies with ($T < 3.5$) are shown by large circles. As can be seen, the data on the top panel are indicative of the known morphological segregation effect when E, S0, and dSph occur usually in groups, but not in the general field. Among the most isolated galaxies with $\Theta < -2.0$, there are exceptionally objects of the latest types, $T = 7$–10. For spiral and irregular galaxies themselves, one can see a slight (30%) decrease in the H I surface brightness from the most isolated ($\Theta < 0$) galaxies toward the most disturbed. This observational fact may be used to choose between two extreme

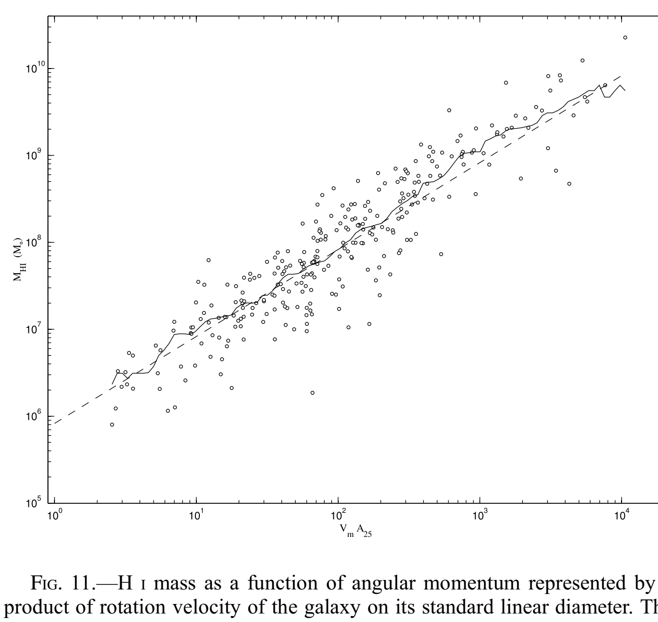

FIG. 11.—H I mass as a function of angular momentum represented by a product of rotation velocity of the galaxy on its standard linear diameter. The running median (*solid line*) corresponds to disklike galaxies ($T > 3.5$). Most of them follow the linear Zasov relation (*dashed line*).

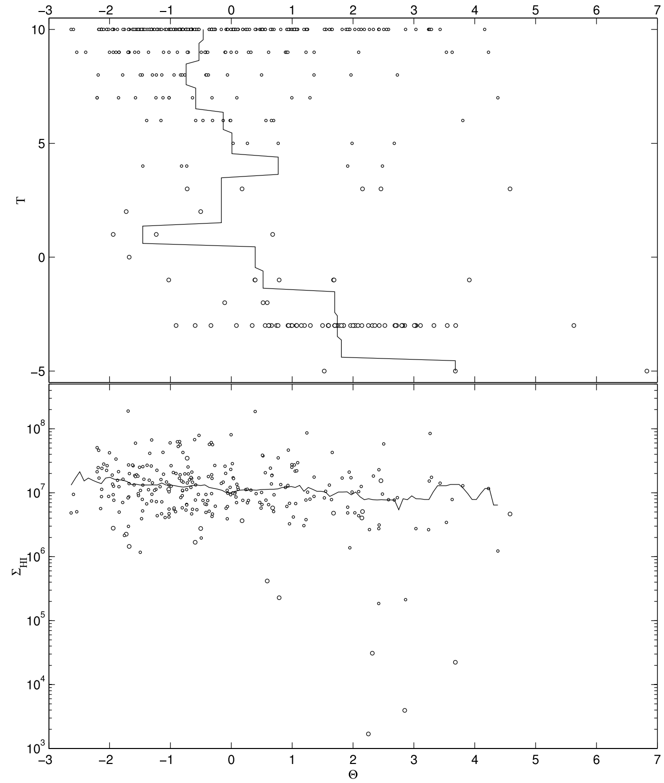

FIG. 12.—Plot of morphological type (*top*) and the mean H I surface brightness (*bottom*) vs. the "tidal index" defined in the text. The medians for disklike galaxies (*small circles*) are shown by solid lines.

scenarios of galaxy evolution when (1) star formation in a galaxy is driven by its gas consumption or (2) the structure of the stellar and gaseous components of a galaxy is defined by the process of recurrent merging of galaxies.

8. SOME COSMOLOGIC PARAMETERS EXTRACTED FROM THE LV SAMPLE

In different scenarios of galaxy evolution under the influence of internal mechanisms (conversion of gas into stars) or external causes (merging), it is suggested that the highest evolutional rate is typical of the largest, the most massive galaxies. The position of a galaxy on the scale "dwarf-giant" may be determined by different parameters: absolute magnitude (M_B), standard linear diameter (A_{25}), rotational velocity (V_m), indicative mass (M_{25}), or hydrogen mass ($M_{\rm HI}$). Three "champions" of the LV according to each of the five mentioned parameters are presented in Table 5. The most luminous galaxy is NGC 4594 ("Sombrero"), and M31 and NGC 253 follow it at 0.32 and 0.53 mag intervals. According to its dimension, M101 stands out significantly against other galaxies. Note that Sandage (1993) considered M101 as a typical Sc galaxy, which led to an untypically low Hubble parameter, $\sim$45 km s^{-1} Mpc^{-1}. Among the most fastrotating galaxies, the first position in the LV is occupied by Sombrero, surpassing Cen A and M 31. For three most massive galaxies—Cen A, NGC 4594, and M31—their indicative masses differ from each other within 0.3 dex. Among the galaxies with the highest hydrogen masses, three Sc galaxies—M101, IC 342, and NGC 5236—are leading. Each of them is surrounded by a suite of dwarf irregular companions (see the last column in Table 4).

Apart from supergiant galaxies in the LV, the dwarf galaxies with unusual global parameters are of great interest too. To select them, it is convenient to use the diagrams $M_{\rm HI}/L$ versus M_{25}/L or $M_{\rm HI}/M_{25}$ versus M_{25}/L presented in Figure 13. Bulgelike galaxies on the diagrams are shown by large circles. As seen, the galaxies are concentrated along a diffuse bifurcated sequence resembling the figure of a running horse. The above-mentioned population of isolated dIr galaxies with absolute magnitudes of [-12,-16] and $M_{25}/L > 10$ corresponds to the "horsehead." Significant deviations of some galaxies from the "main sequence" on both the panels may be caused either by variations of initial conditions (say, dark-to-baryonic mass ratio) or by the difference in galaxy evolution tracks (bursts of star formation, mergings). We did not find in the literature any distinct suggestions predicting distribution of galaxies on such diagrams based on the existing ideas of formation and evolution of galaxies of different masses. On the other hand, photometric and kinematic data on many galaxies in Figure 13 remain too rough, which restricts their evolutional interpretation.

Based on the data of Tables 1 and 4, we estimate some parameters important for cosmology. Figure 14 (*bottom*) gives the mean local luminosity density within a sphere of radius D in $L_\odot$ Mpc^{-3}. Apart from the peak produced by the Local Group, the luminosity density shows a secondary peak at $D = 3.7$ Mpc, caused by the M81 and Cen A groups, and then decreases smoothly to the value 3.5×10^8 $L_\odot$ Mpc^{-3} at $D = 8$ Mpc. Subsequent decrease in the luminosity density occurs because of the fractal nature of the matter distribution. Another reason might be caused by an incompleteness of our sample. The global value of $\rho_{\rm lum}$ in the B band was estimated by Efstathiou, Ellis, & Peterson (1988) and others on the basis of essentially deeper samples. Recently, Blanton et al. (2003) and Liske et al. (2003) have derived from the Sloan Digital Sky Survey and the Millennium Galaxy Catalogue the values 1.23×10^8 $L_\odot$ Mpc^{-3}

TABLE 5
PRINCIPAL GALAXIES IN THE LV

Galaxy	M_B (mag)	A_{25} (kpc)	V_m (km s^{-1})	log ($M_{25}/M_\odot$)	log ($M_{\rm HI}/M_\odot$)
IC 342	...	...	...	...	10.09
M31	−21.58	36	255	11.43	...
M101	...	61	...	...	10.36
NGC 253	−21.37	...	...	...	...
NGC 4258		36	...	...	...
NGC 4594	−21.90	...	413	11.62	...
NGC 5128	...	...	398	11.73	...
NGC 5236	...	...	...	...	9.92

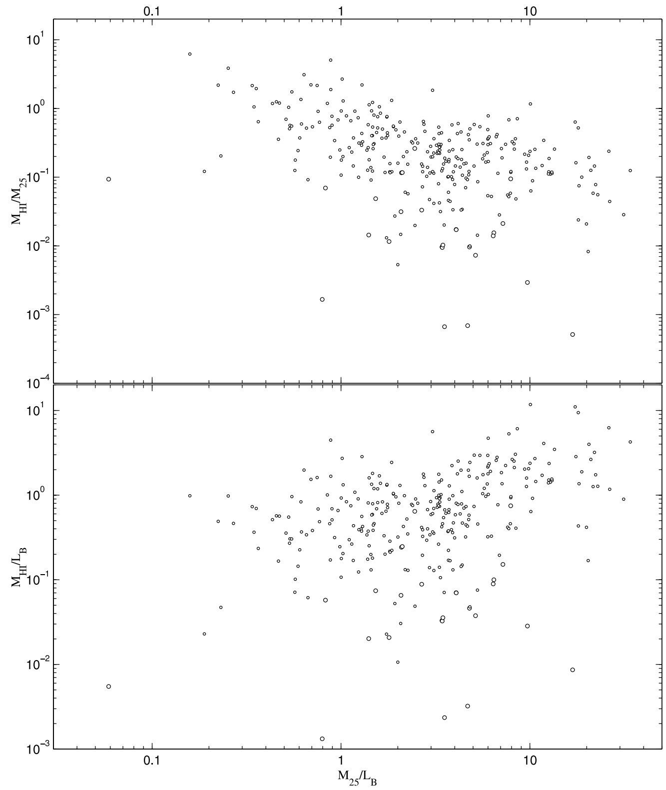

Fig. 13.—H i mass-to-luminosity ratio (*bottom*) and the H i mass-to-indicative mass ratio (*top*) for the galaxies with distances $D < 8$ Mpc vs. their indicative mass-to-luminosity ratio. The early-type objects with $T < 3.5$ are shown by large circles.

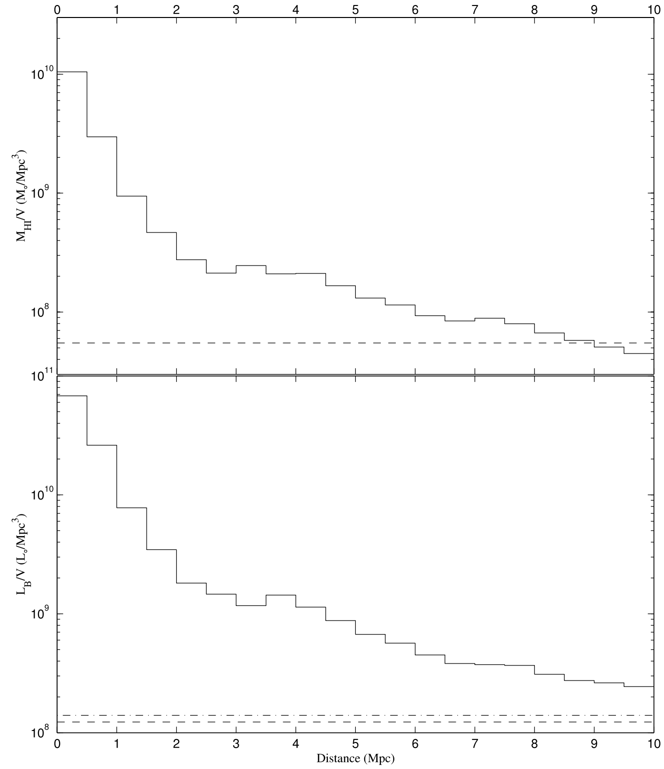

Fig. 14.—Mean local luminosity density (*bottom*) and the mean local H i mass density (*top*) as a function of the radius of the local sphere D. The global values from (Blanton et al. 2003, Liske et al. 2003, and Zwaan et al. 2003) are shown by the horizontal lines.

and $1.43 \times 10^8\ L_\odot$ Mpc^{-3}, respectively. Both the quantities are reduced to the local value of the Hubble parameter, $H = 72$ km s^{-1} Mpc^{-1}. It should be noted, however, that the global estimates of $\rho_{\rm lum}$ were derived when internal extinction in galaxies was not taken into account. Ignoring the internal extinction, we obtain $\rho_{\rm lum}(<8\ {\rm Mpc}) = 2.5 \times 10^8\ L_\odot$ Mpc^{-3}, then the mean local luminosity density exceeds 1.7–2.0 times the global density in spite of the presence of the Tully void and the absence of rich clusters in the LV. Dwarf galaxies with $M_B > -17.0$ contribute only 4% to the local luminosity density, although their relative number within 8 Mpc is 85%. Therefore, the mentioned incompleteness of our sample on its faint luminosity end does not affect essentially the derived value of local luminosity density.

Figure 14 (*top*) presents the mean local density of neutral hydrogen as a function of radius D. The dashed line corresponds to the "global" estimate, $\Omega_{\rm HI} = (3.8 \pm 0.6) \times 10^{-4}$, derived by Zwaan et al. (2003) from the HIPASS. Here we took into account that the critical density $\rho_c = 0.97 \times 10^{-29}$ g cm^{-3} = $1.43 \times 10^{11}\ M_\odot$ Mpc^{-3}, when $H = 72$ km s^{-1} Mpc^{-1}. Therefore, in the sphere of radius $D = 8$ Mpc, the mean local density of H i mass is $0.78 \times 10^8\ M_\odot$ Mpc^{-3}, which exceeds 1.44 times the HIPASS estimate. (We consider, however, that the shallow HIPASS survey with its typical depth of $V \sim$ 2300 km s^{-1} overestimates appreciably the global value of the H i density). We found above that the mean indicative mass–to—blue luminosity ratio for nearby galaxies is 3.8 $M_\odot/L_\odot$. Assuming the total baryon mass of each galaxy to be 2.5 times its indicative mass within A_{25}, we derive the mean local baryon density $\Omega_b(<8\ {\rm Mpc}) = 2.3\%$. According to Turner (2002), the mean global baryon density is $\Omega_b = (3.9 \pm 0.8)\%$.

Therefore, there is a paradox of local baryon density that consists of only one-half of the global baryon density. It is likely that the remaining half of the local Ω_b spreads as warm gas in the intergalactic medium (Sembach et al. 2003; Braun & Thilker 2003). Q8

Based on the Table 1 data, we estimated another important cosmologic parameter, an expected number of intersections of the line of sight with optical images of galaxies. Figure 15 shows the integral area of the sky, $S(D)_{\rm sky} = \sum_i (\pi a_i b_i/4)$,

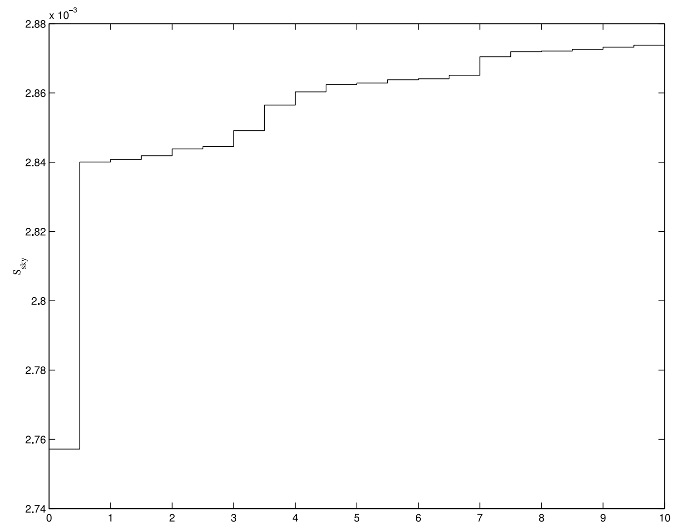

Fig. 15.—Integrated area of the sky occupied by optical images of galaxies within a distance D from us.

occupied by galaxies within a distance D from us. Here, a and b stand for the major and minor diameter of galaxy taken from columns (3) and (4). Within the distance $D = 8.3$ Mpc, which corresponds to the redshift $Z = 0.002$ ($H = 72$ km s^{-1} Mpc^{-1}), all optical images of galaxies cover a fraction of the entire sky of 2.87×10^{-3}. Extrapolating this value over a sample with the depth $Z = 1$ and ignoring any evolution effect, we obtain the expected number of the line of sight intersections with galaxy images to be 1.44. However, the Magellanic Clouds make an overwhelming contribution to the estimate of $S_{\rm sky}$. Excluding the Magellanic Clouds drops the relative sky area covered by galaxies to $S_{\rm sky}(Z < 1) = 0.06$. This quantity fits approximately the view of the Hubble Deep Field (Bernstein, Freedman, & Madore 2002), where optical images of galaxies occupy about 10% of the sky. If one adopts the H I size of spiral and irregular galaxies to be 2.5 times the standard optical diameter, the expected number of the line-of-sight intersections with the H I galaxy images increases to $dn/dz \sim 0.4$. This estimate does not contradict the observed number of absorptions in QSO spectra.

As it was mentioned above, there is a pessimistic opinion that about 80% dwarf galaxies are hidden from us because of their extremely low surface brightness (McGaugh 1996; Impey & Bothun 1997). Such a strong argument can be checked, in principle, by observations. In Table 6, we present a relative contribution of dwarf galaxies having absolute magnitudes fainter than -17.0 to different characteristics of the LV. As it is seen, the dwarf galaxies (with their relative number of 85%) make a minor contribution, roughly 2%–5%, to the integrated luminosity or integrated "indicative" mass of the neighboring galaxies. However, their contribution in the H I mass (10–15)%, as well as in the sky area covered by galaxies (31%), turns out to be more significant. The supposed presence of an about 5 times larger population of ELSB dwarf galaxies would lead to an unusually high number of absorp-
Q9 tion lines seen in QSO spectra. Note also that the recent very deep optical study of the M83 group by Minchin et al. (2003) does not confirm the existence of a huge number of hidden ELSB dwarfs.

To characterize the average virial galaxy motions on different scales, the mean-square difference of radial velocities of two galaxies, $\langle \Delta V_{12}^2(R_p) \rangle$, as a function of their projected linear distance is often used. It is supposed that this quantity allows one to estimate the dark matter density on corresponding scales. The existing estimates of $\langle \Delta V_{12}^2(R_p) \rangle$, obtained from different catalogs of galaxies and surveys, differ essentially from each other but lie in the range from 200 km s^{-1} (Branchini et al. 2001) to 600 km s^{-1} (Jing, Mo, & Borner 1998; Zehavi et al. 2002), which leads to $\Omega_{\rm DM} \sim (0.1–0.3)$ on a scale of $\sim$1 Mpc. It should be emphasized that the derived quantities were obtained for galaxy samples restricted by the flux, but not the distance. Based on the data of Table 1, we calculated $\langle \Delta V_{12}^2 \rangle$ as a function of space separation, as well as

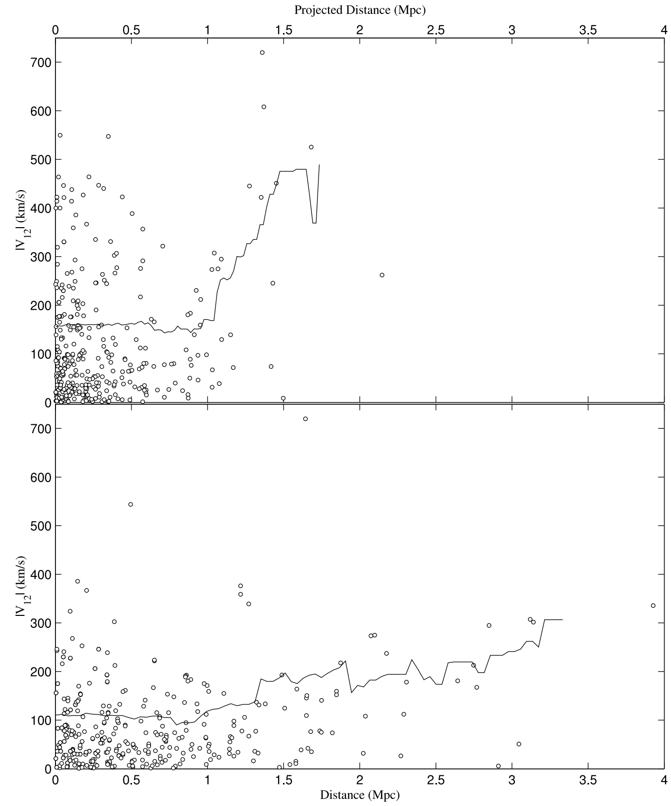

Fig. 16.—Mean-square difference of radial velocities of two galaxies in the LV as a function of their spatial (*bottom*) and projected (*top*) separation.

projected separation between the galaxies. The results are presented by solid lines in the bottom and top panels of Figure 16, respectively. The mean-square difference of radial velocities changes slightly from 110 to 90 km s^{-1} on a scale of $R < 1$ Mpc, apparently reflecting the Keplerian motions in tight galaxy pairs. The increase in $\langle \Delta V_{12}^2 \rangle$ to $\sim$300 km s^{-1}, seen on scales of 1–4 Mpc, is caused by the increasing role of the systematic Hubble component at larger mutual galaxy separations. The two pairs of galaxies, NGC 4600 + UGC 7857 and UGC 8061 + UGC 7857, with mutual velocities of greater than 500 km s^{-1}, are situated in the direction toward the Virgo cluster and are considered to benonphysical ficti-
tious pairs. The solid line on the top panel shows a steep rise Q10
of the mean-square velocity difference to 450 km s^{-1} at projected separations of 1–2 Mpc. As one can see, the projection effects distort essentially the observed behavior of $\langle \Delta V_{12}^2(R) \rangle$, which makes the interpretation of the relation in terms of $\Omega_{\rm DM}$ rather uncertain.

We hope that this catalog will help observers to select the nearby galaxies for study of their detailed kinematics and stellar population, as well as proper kinematics of the local universe.

TABLE 6

Integrated Contribution of Dwarf Galaxies with $M_B > -17$ mag to Different Local Quantities

Quantity	$D \leq 2.0$ Mpc (%)	$D \leq 5.0$ Mpc (%)	$D \leq 8.0$ Mpc (%)	$D \leq 10$ Mpc (%)
Luminosity density	2.8	4.0	4.3	3.7
"Total" mass density	1.7	3.0	5.1	4.2
H I mass density	9.5	15.2	15.4	15.4
Relative area of the sky covered by dwarf galaxies	31.2	31.1	31.0	31.0

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Q1 Au: according to journal style, the abbreviation "HIPASS" is to be spelled out at its first mention in the abstract and again in the text. Treatment OK?
Q2 Au: revisions to this sentence OK?
Q3 Au: OK to spell out "LSB" as "low surface brightness"?
Q4 Au: since the Notes to Table 1 are in themselves a separate table, this table has been named Table 2, and the rest of the tables have been renumbered accordingly.
Q5 Au: Revisions to this sentence OK?
Q6 Au: OK to replace subscript "*M*" with " *m*"?
Q7 Au: revision to this sentence OK?
Q8 Au: revisions to this sentence OK?
Q9 Au: please define what the abbreviation "ELSB" stands for.
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Q11 Au: please note that private communications are not included in the reference list. These have been moved to their appearances in the text and tables.
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